UNIVERSITÉ DE GENÈVE

OBSERVATOIRE DE GENÈVE

Stochastic acceleration in blazars

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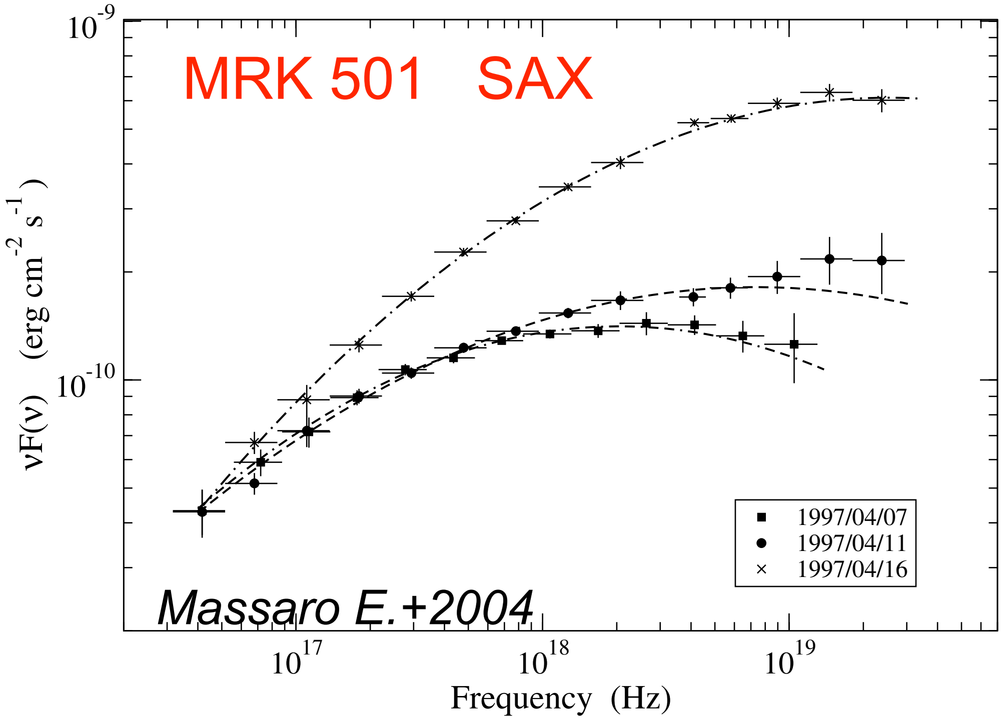
Geneva

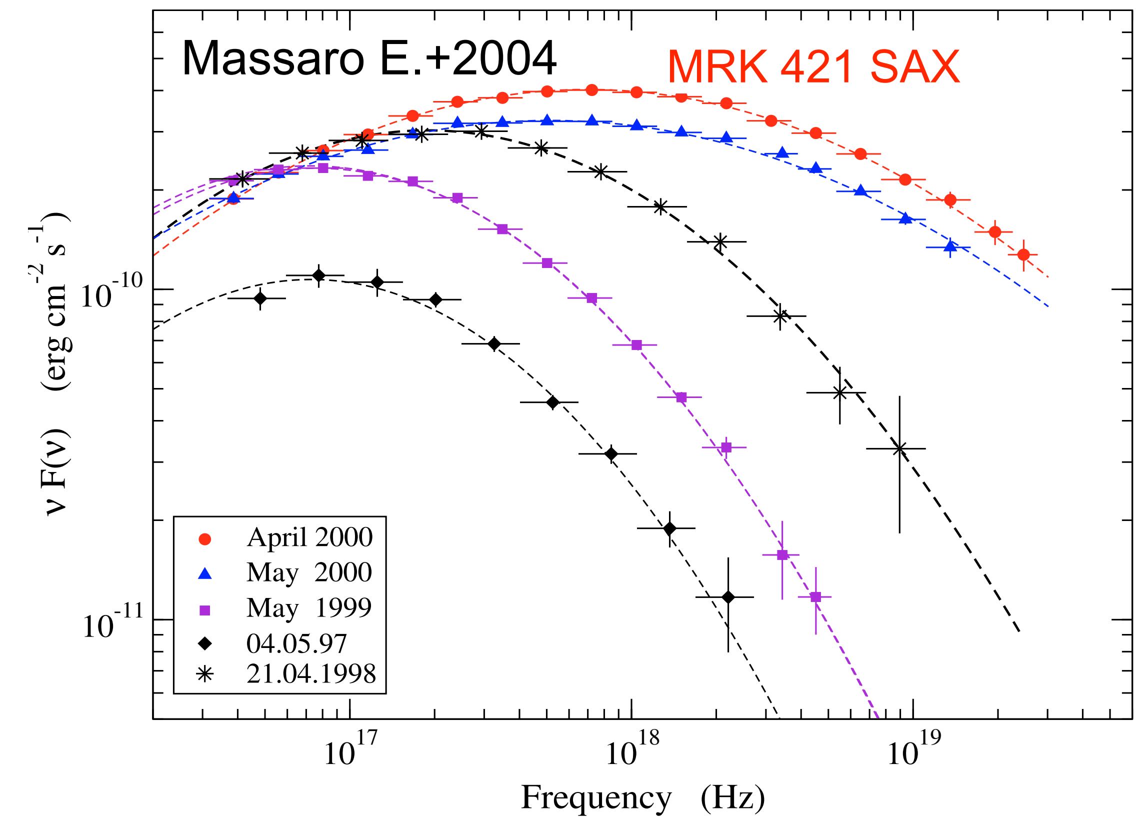
12-February 2019

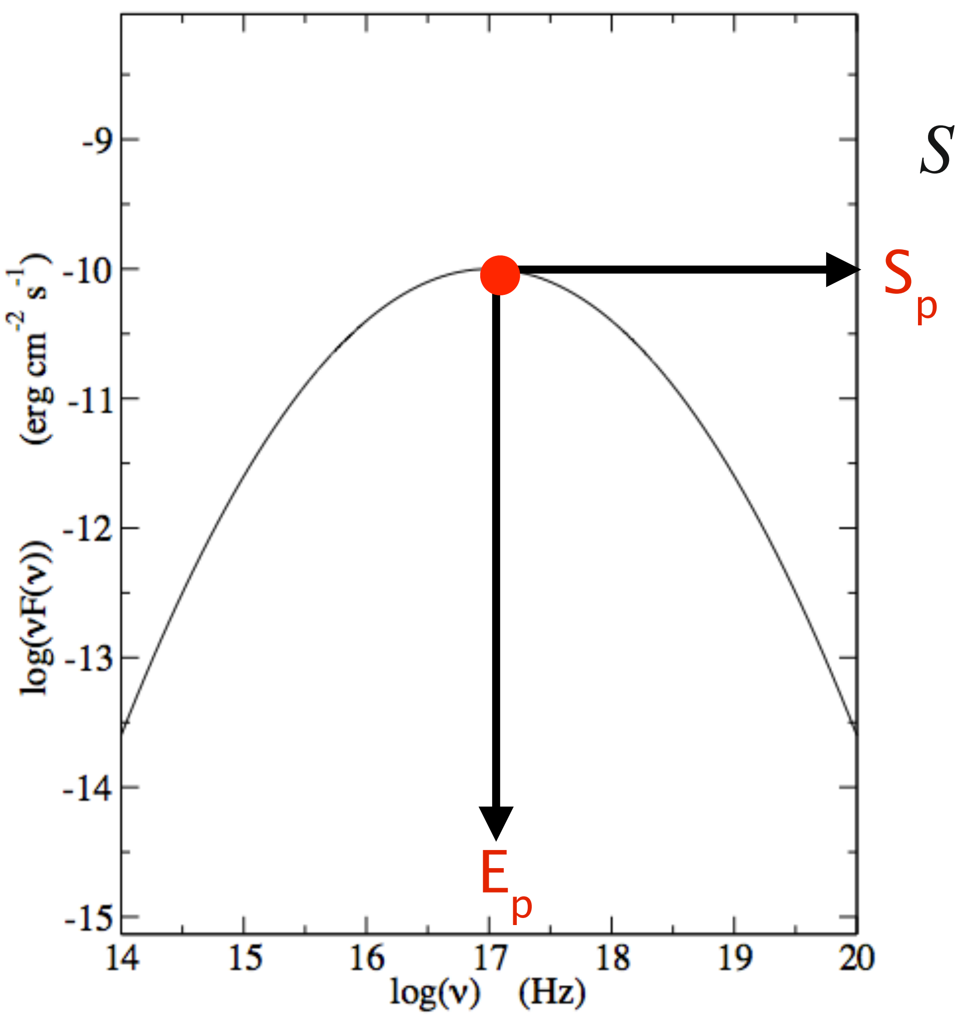
Outline

- Phenomenological signatures
- setup of Theory/Numerical framework for stochastic acceleration
- Self-consistent reproduction of Long Term Trends
- numerical modeling, numerical fit (no eyeball fit) no analytical approximations

SPECTRAL DISTRIBUTION OF HBLs





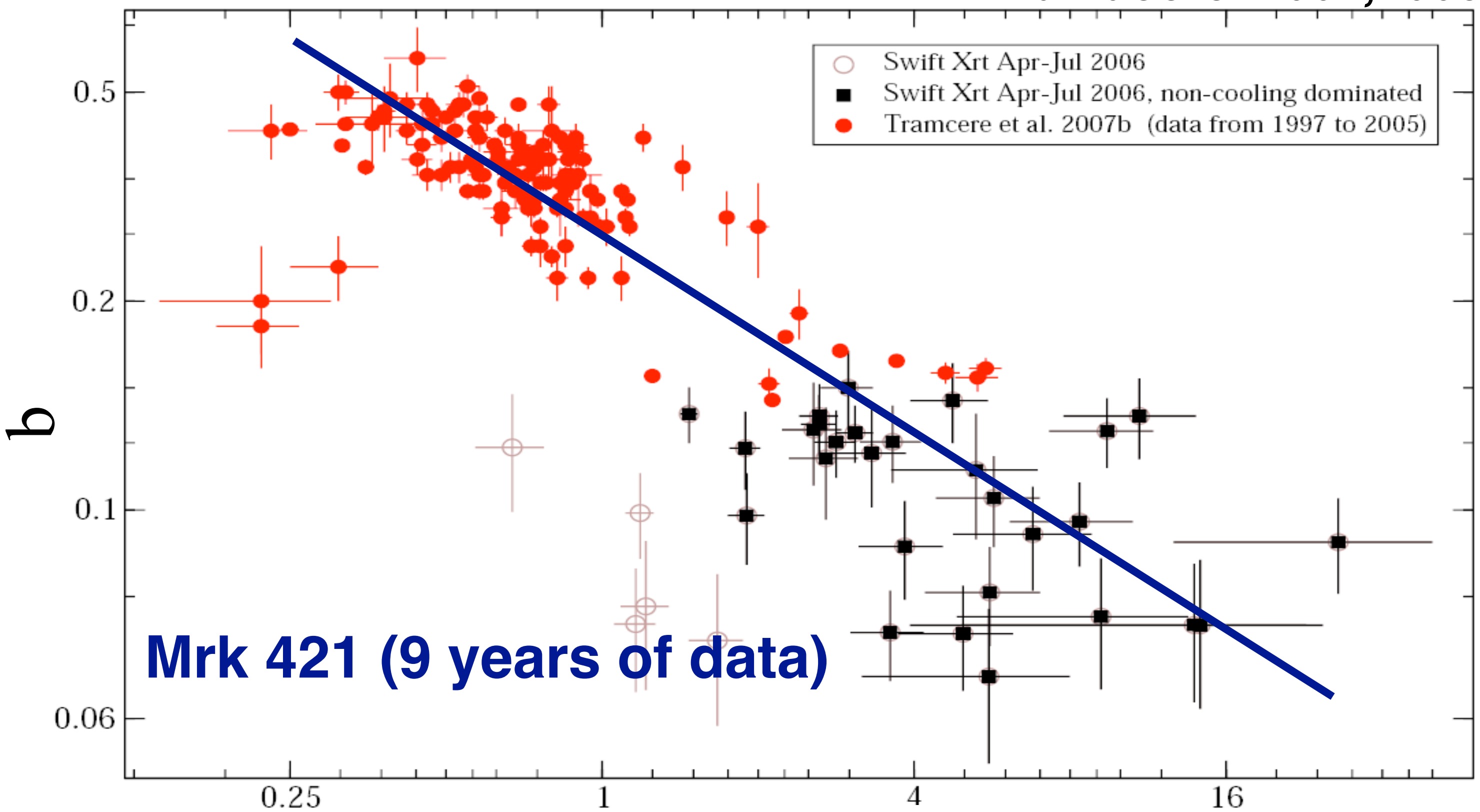


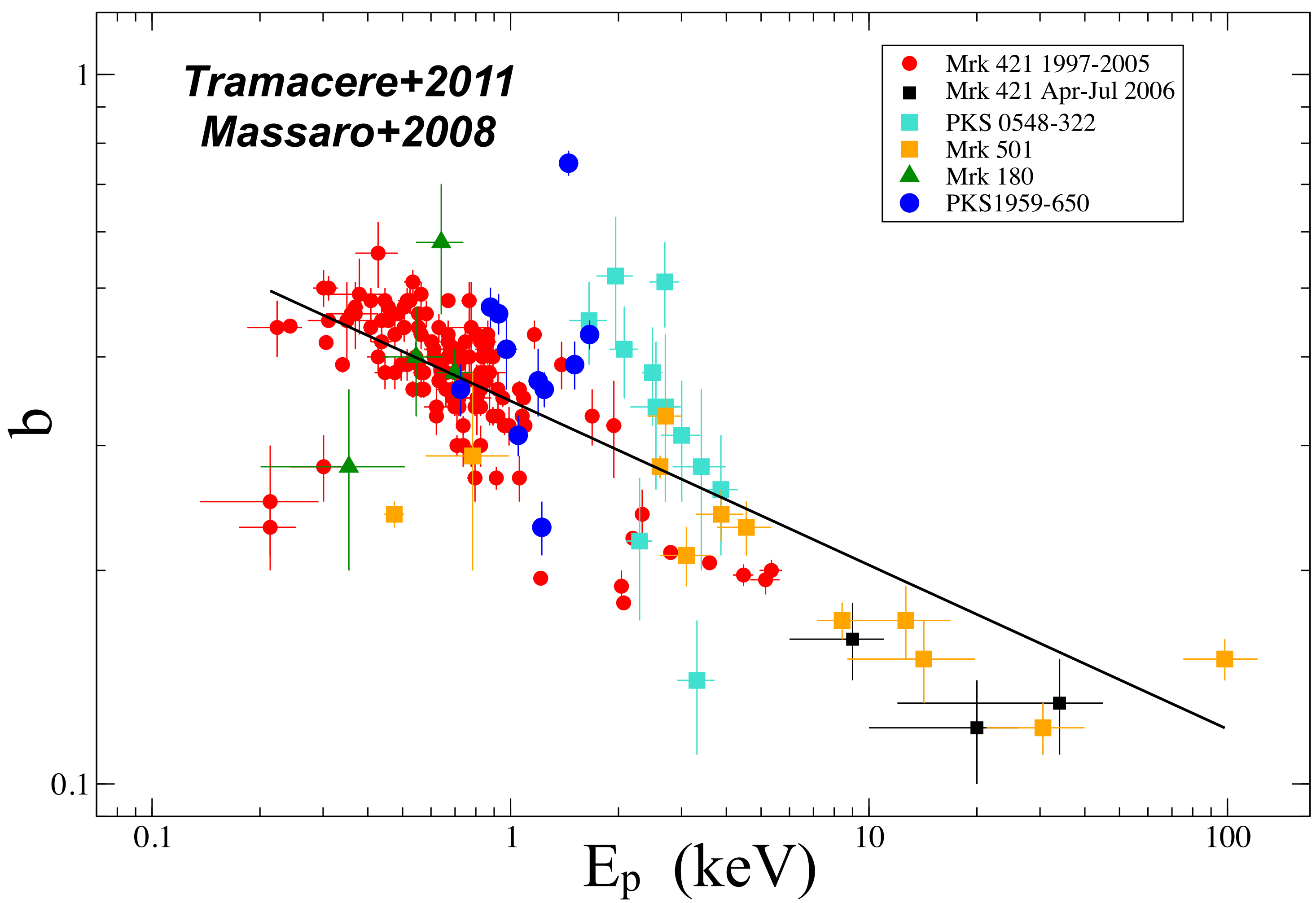
$$S(E) = S_{\rm p}\, 10^{-b\,(\log(E/E_{\rm p}))^2}$$

- b: curvature at peak
- E_p: peak energy
- S_p: SED height @ E_p

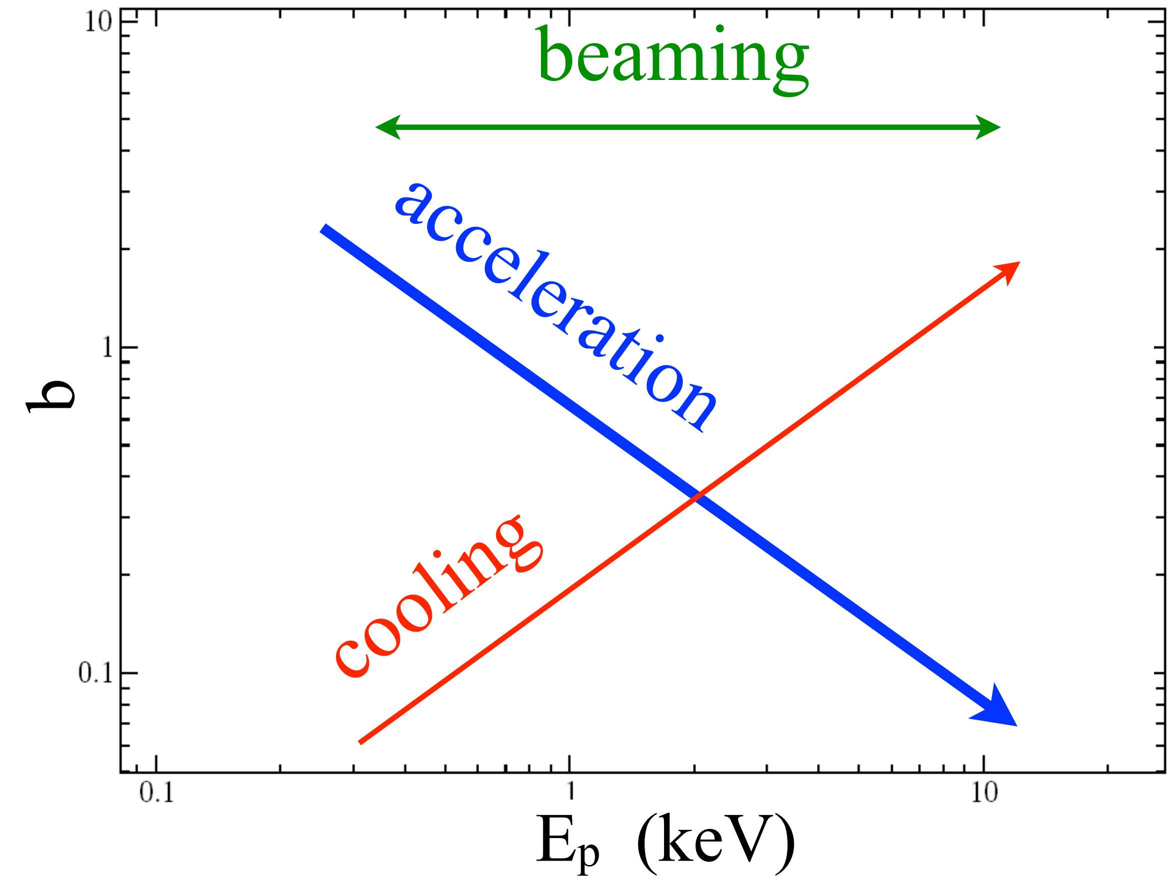
acceleration signature in the E_S-vs-b trend

Tramacere+2007,2009





Ep-vs-b, different scenarios

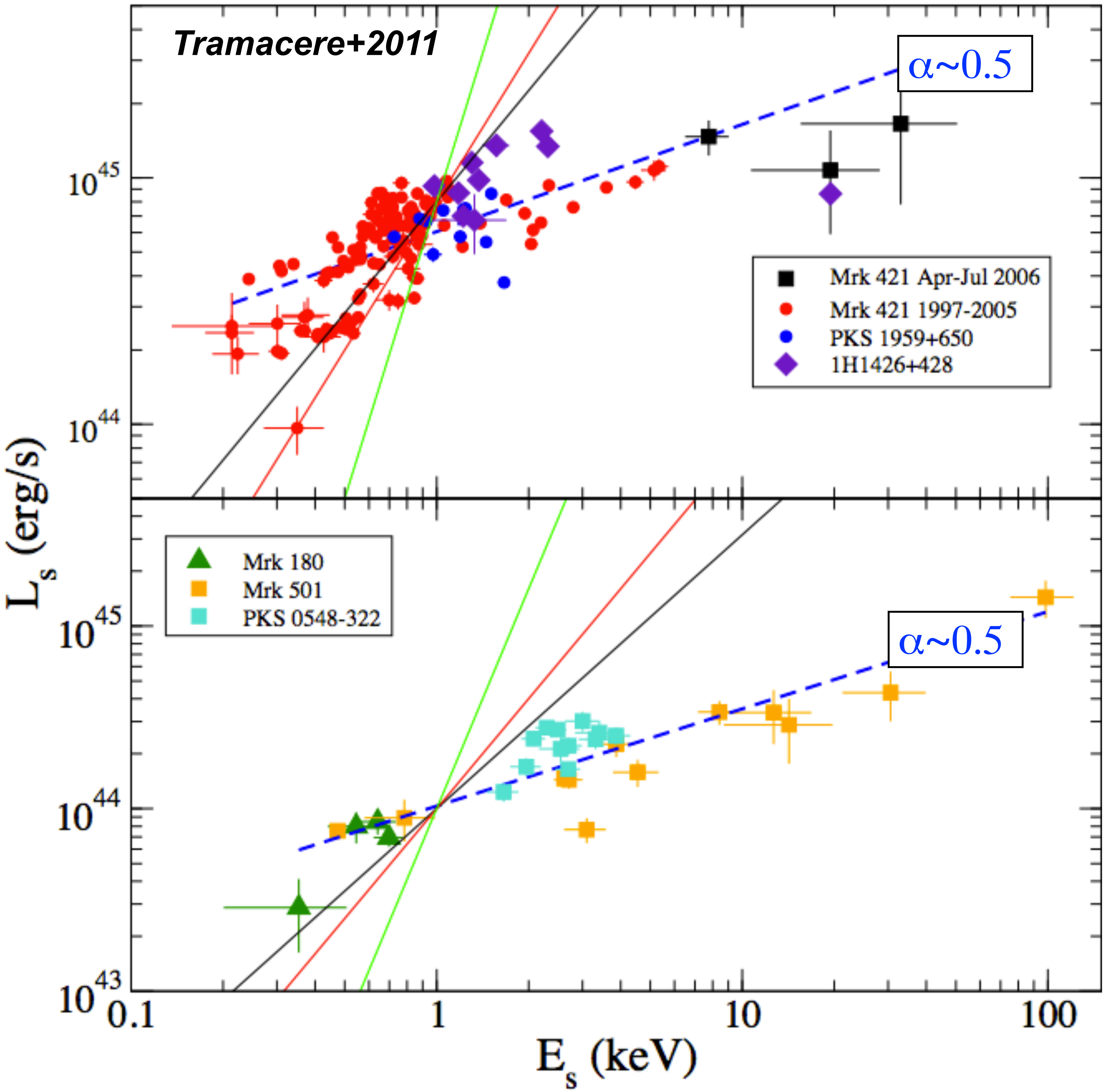


11 years of data:
PKS 0548-322,1H1426+418,
Mrk 501 ,1ES1959+650,
PKS2155-34

Long term (overall 13 years of data)
Ep-vs-b trends hint for an acceleration
dominated scenario

acceleration signature in the E_S-vs-L_s trend

long-trend main drivers



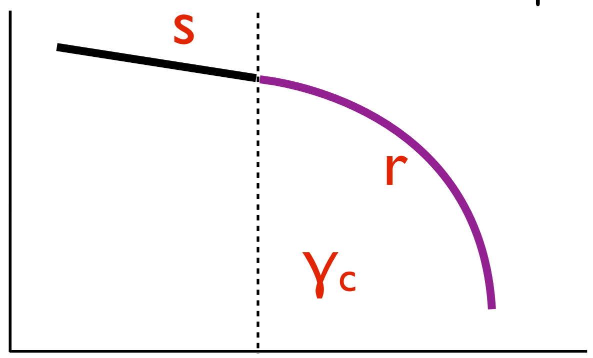
- γ_{3p}↑ and $n(\gamma 3p)$↓ => $\alpha<1.5$
  acceleration+energy conservation

- B–>α=2.0,
- δ–>α=4

incompatible as long-trend main driver

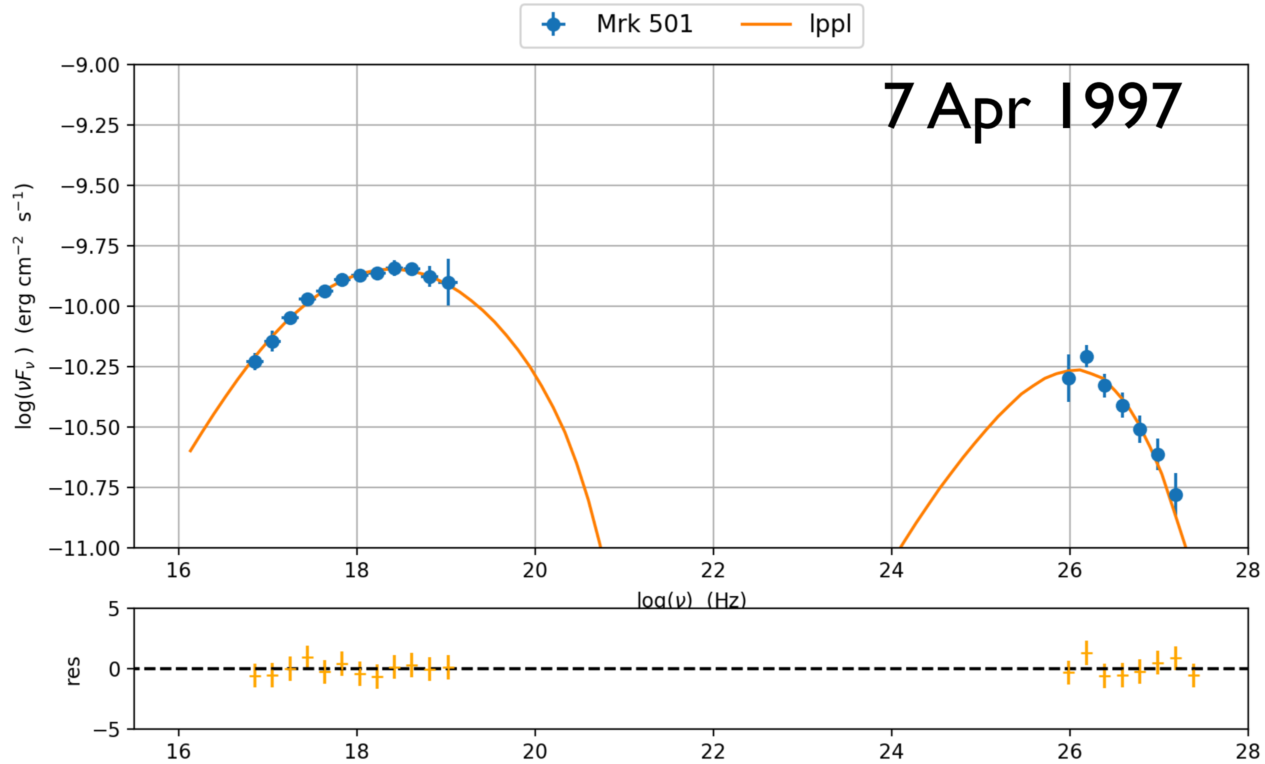
Hard spectra $s<<2.00$

$$s = 1 + \frac{t_{acc}}{2t_{esc}}$$



Mrk 501 1997 Flare

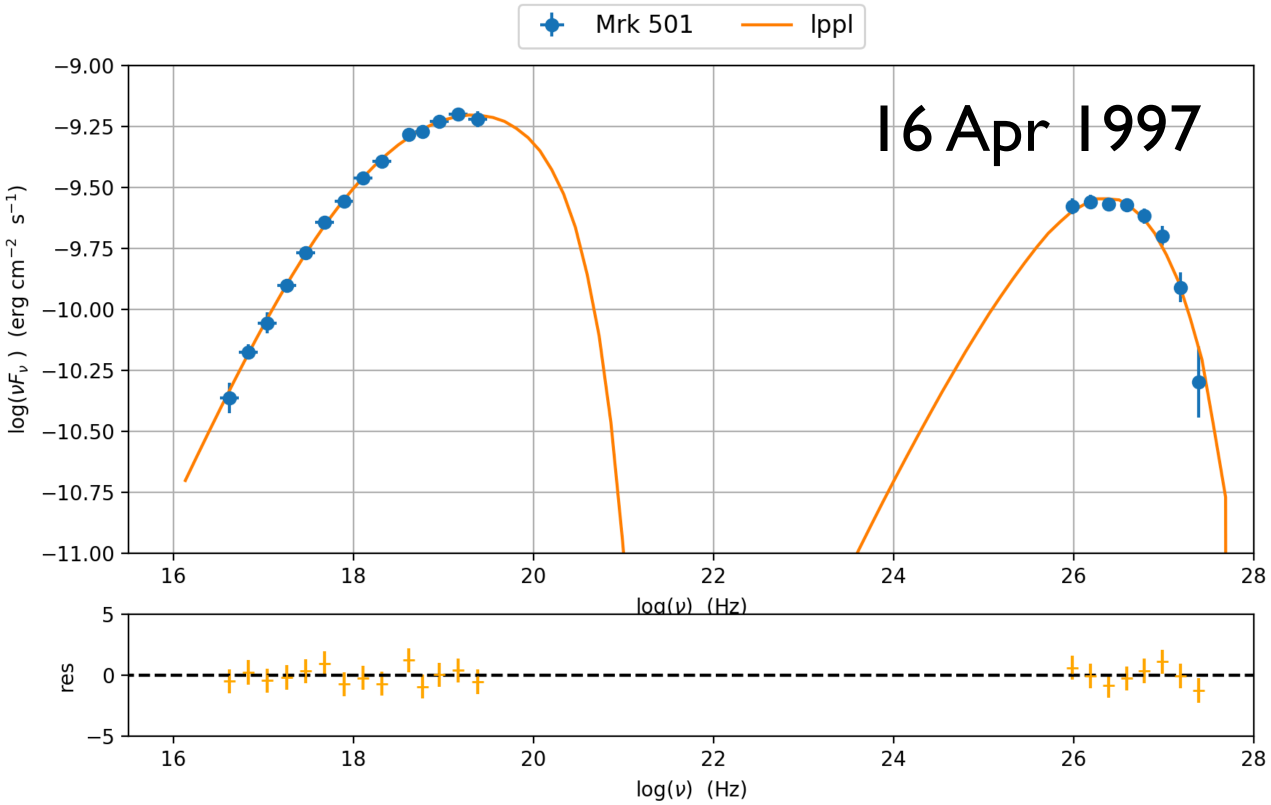
Massaro & Tramacere +2006



best fit pars

best-fit parameters:

Name	best-fit value	best-fit err +
B	+1.072178e-01	+5.436622e-03
N	+4.585348e+00	+4.756569e-01
R	Frozen	Frozen
beam_obj	+2.450884e+01	+7.642113e-01
gamma0_log_parab	+6.609649e+04	+7.427709e+03
gmax	+1.860044e+14	+5.881595e+14
gmin	+1.404527e+03	+2.198648e+02
r	+7.513452e-01	+5.059815e-02
s	+1.638026e+00	+3.170384e-02
z_cosm	Frozen	Frozen



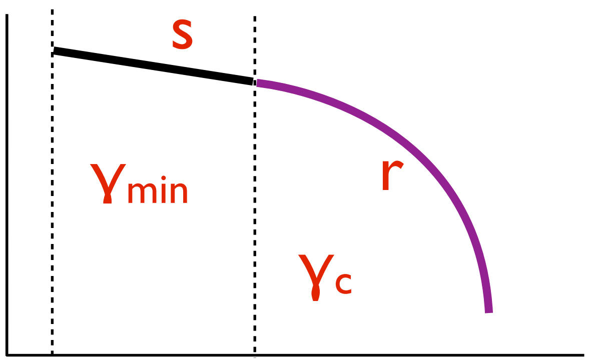
best-fit parameters:

Name	best-fit value	best-fit err +
B	+3.065207e-01	+1.159567e-02
N	+1.079944e+02	+7.375385e+00
R	Frozen	Frozen
beam_obj	+2.722013e+01	+5.889626e-01
gamma0_log_parab	+6.493888e+04	+5.410315e+03
gmax	+1.902146e+06	+2.216666e+02
gmin	+3.003970e+02	+5.686711e+01
r	+6.778727e-01	+3.526656e-02
s	+1.321307e+00	+1.844825e-02
z_cosm	Frozen	Frozen

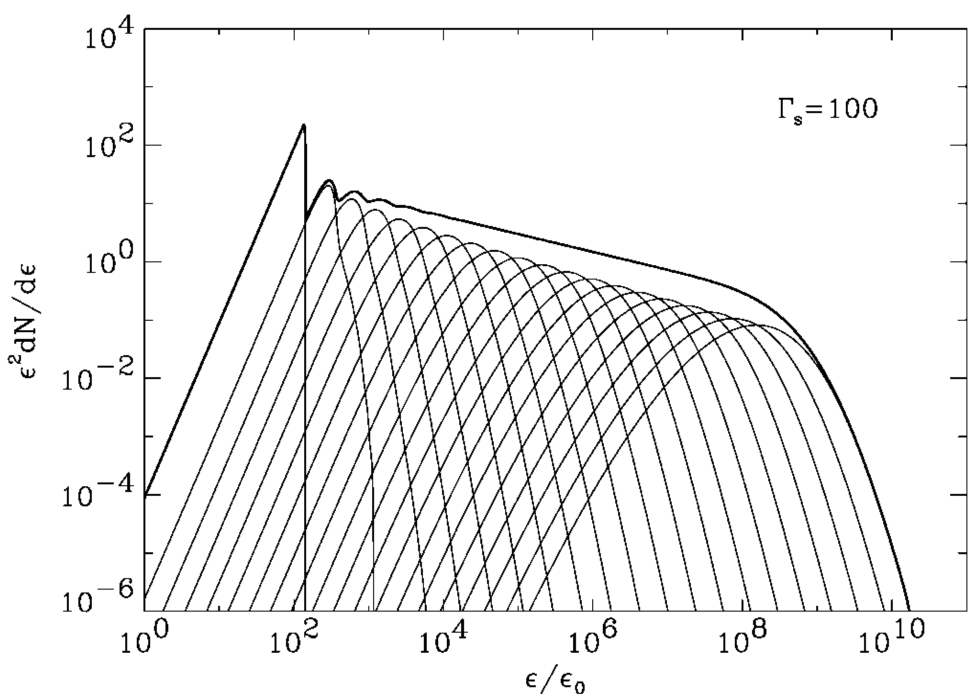
Fermi I+Fermi II Mrk 421 2006

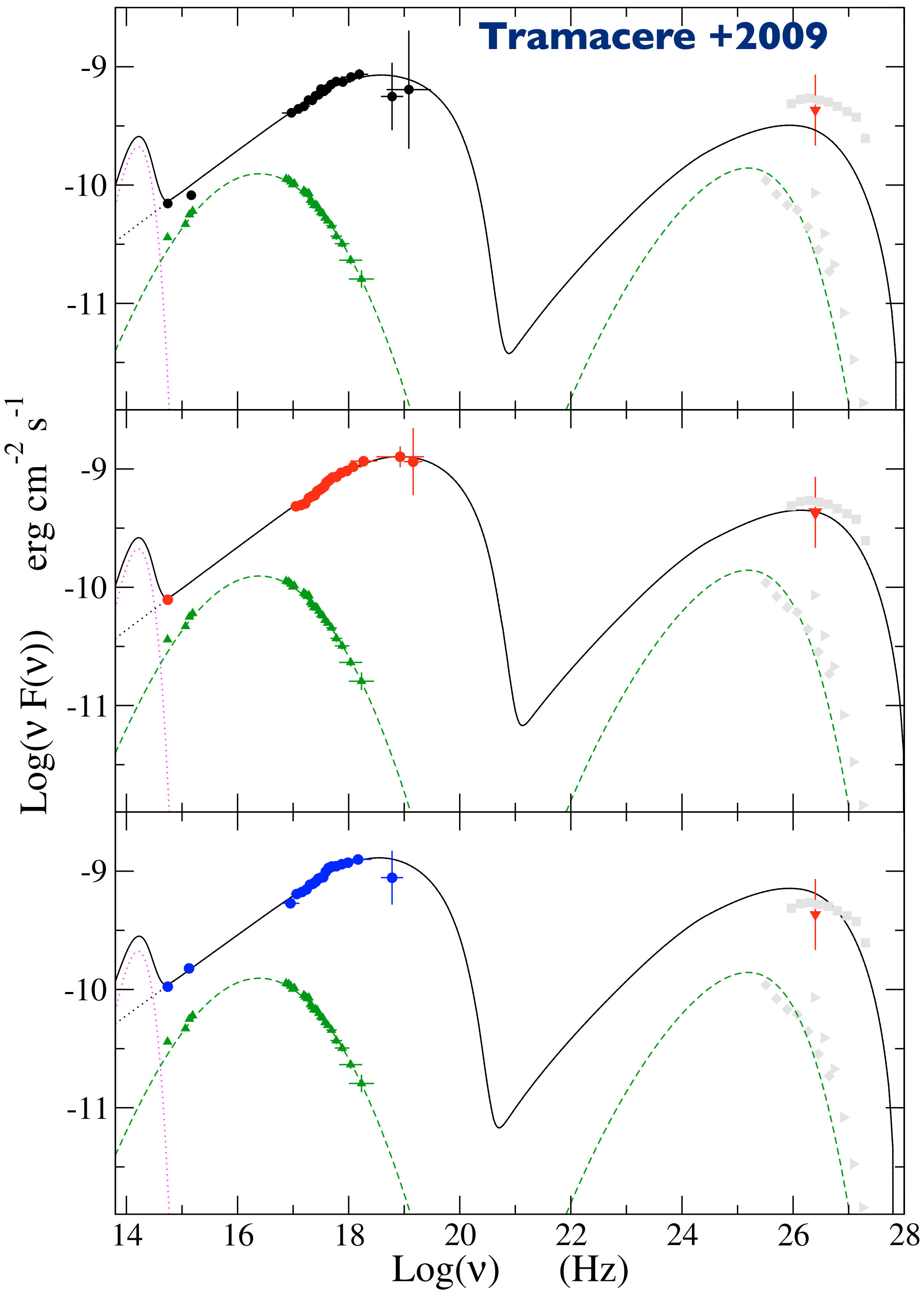
LP+PL spectra

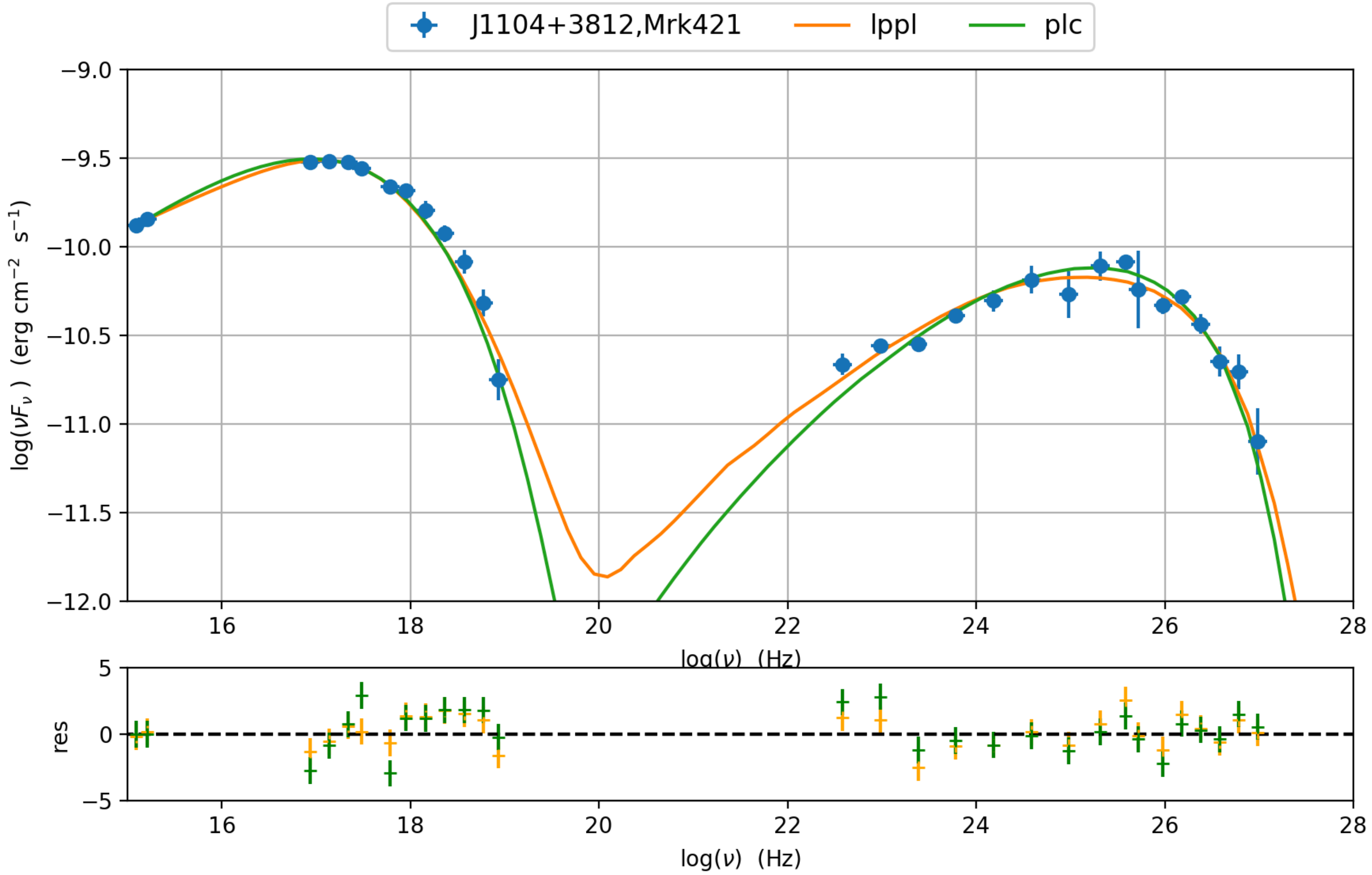
Synch index~[1.6-1.7]=>s~[2.2-2.4]



Lemoine,Pelletier 2003








```
dof=21
chisq=39.696427, chisq/red=1.890306 null hypothesis

best fit pars
---------------------------------------------------
best-fit parameters:
  Name              | best-fit value| best-fit err +|
---------------------------------------------------
 B                  | +2.096016e-02 | +5.744998e-05 |
 N                  | +1.152143e-01 | +1.545857e-03 |
 R                  | Frozen        | Frozen        |
 beam_obj           | +2.619674e+01 | +8.501912e-02 |
 gamma0_log_parab   | +1.884210e+05 | +1.891713e+03 |
 gmax               | +3.492780e+08 | +6.130842e+08 |
 gmin               | +1.929302e+03 | +2.109472e+01 |
 r                  | +1.681768e+00 | +3.032664e-02 |
 s                  | +2.509224e+00 | +2.902511e-03 |
 z_cosm             | Frozen        | Frozen        |
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Mrk 421 2009 data

data from Abdo et al 2011
Fermi-LAT+Magic coll.

lppl/plc p-value= 6.8E-6

The log-parabola origin:
physical insight

The origin of the log-parabolic shape: statistical derivation

fluctuation

$$\varepsilon = \bar{\varepsilon} + \chi$$

systematic

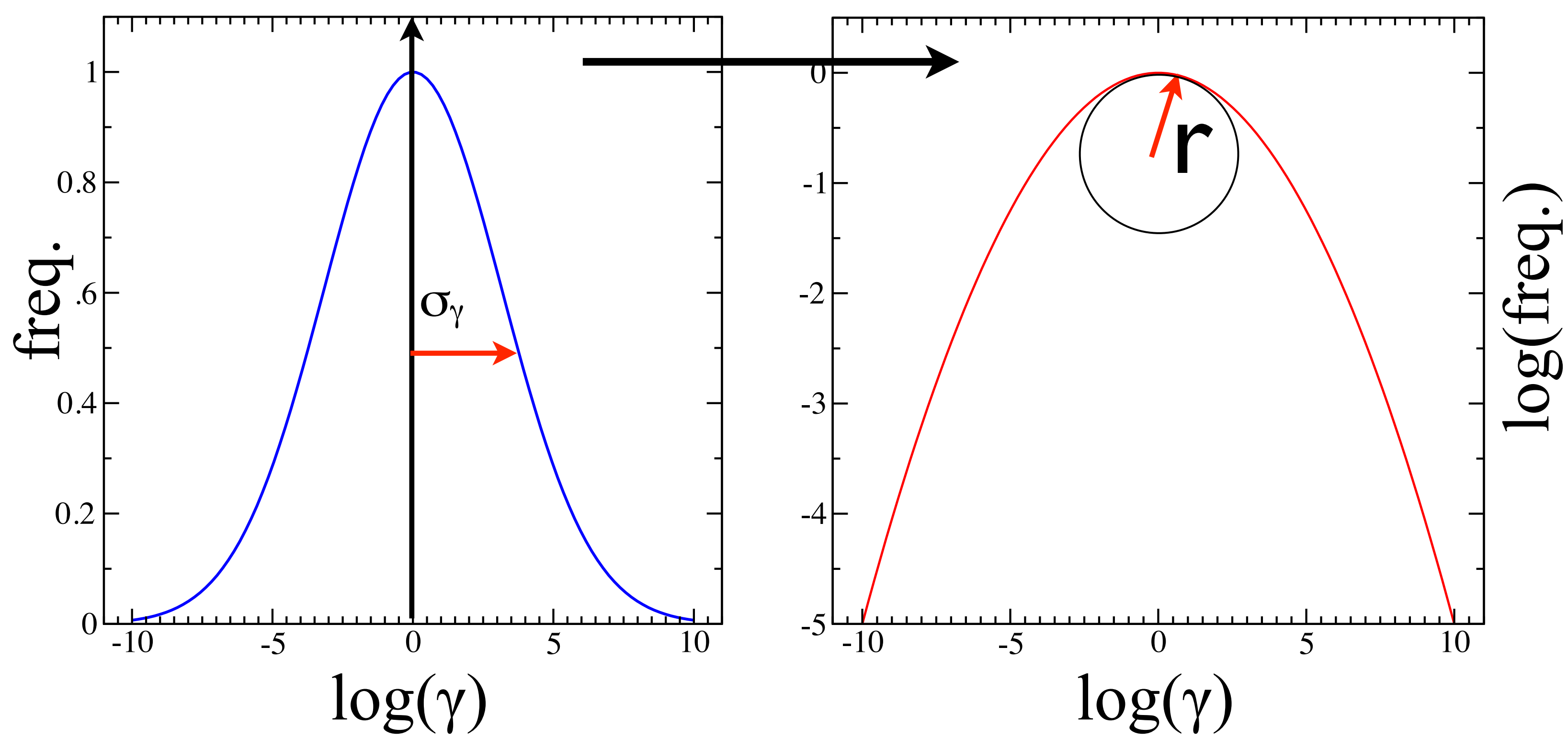
ε_i is a R.V.

$$\gamma_{n_s} = \gamma_0 \Pi_{i=1}^{n_s} \varepsilon_i$$

C.L. Theorem
multipl. case

log-normal distribution

Log-Parabolic representation



$$\log(n(\gamma)) \propto \frac{(\log\gamma - \mu)^2}{2\sigma_\gamma^2} \propto r \; [\log(\gamma) - \mu]^2$$

$$\frac{\partial n(\gamma,t)}{\partial t} = \frac{\partial}{\partial\gamma}\Big\{-[S(\gamma,t)+D_A(\gamma,t)]n(\gamma,t)+D_p(\gamma,t)\frac{\partial n(\gamma,t)}{\partial\gamma}\Big\} - \frac{n(\gamma,t)}{T_{esc}(\gamma)} + Q(\gamma,t)$$

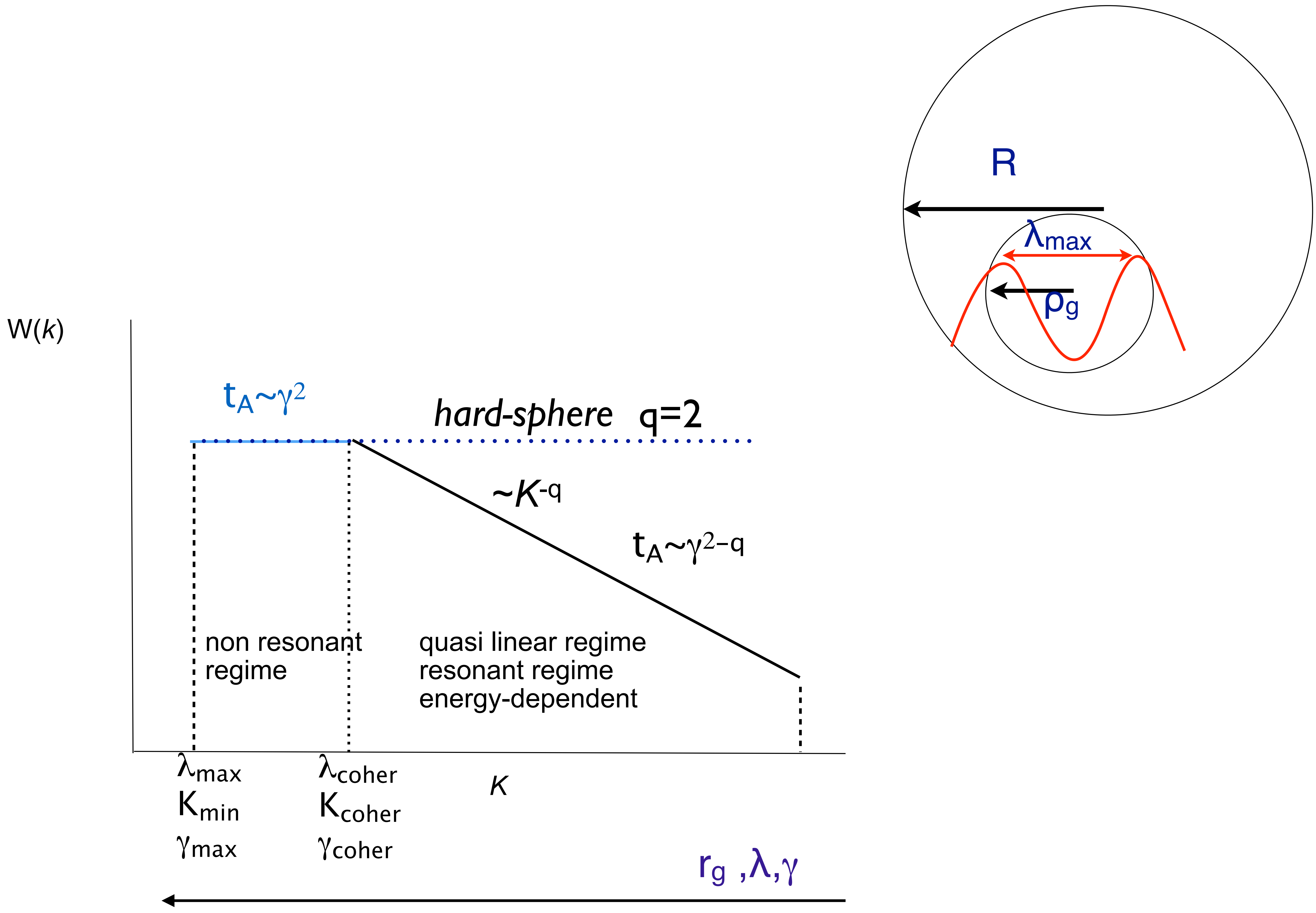
analytical solution for:
$D_p \sim \gamma^q$, $q=2$

“hard-sphere” case no cooling

Melrose 1968,

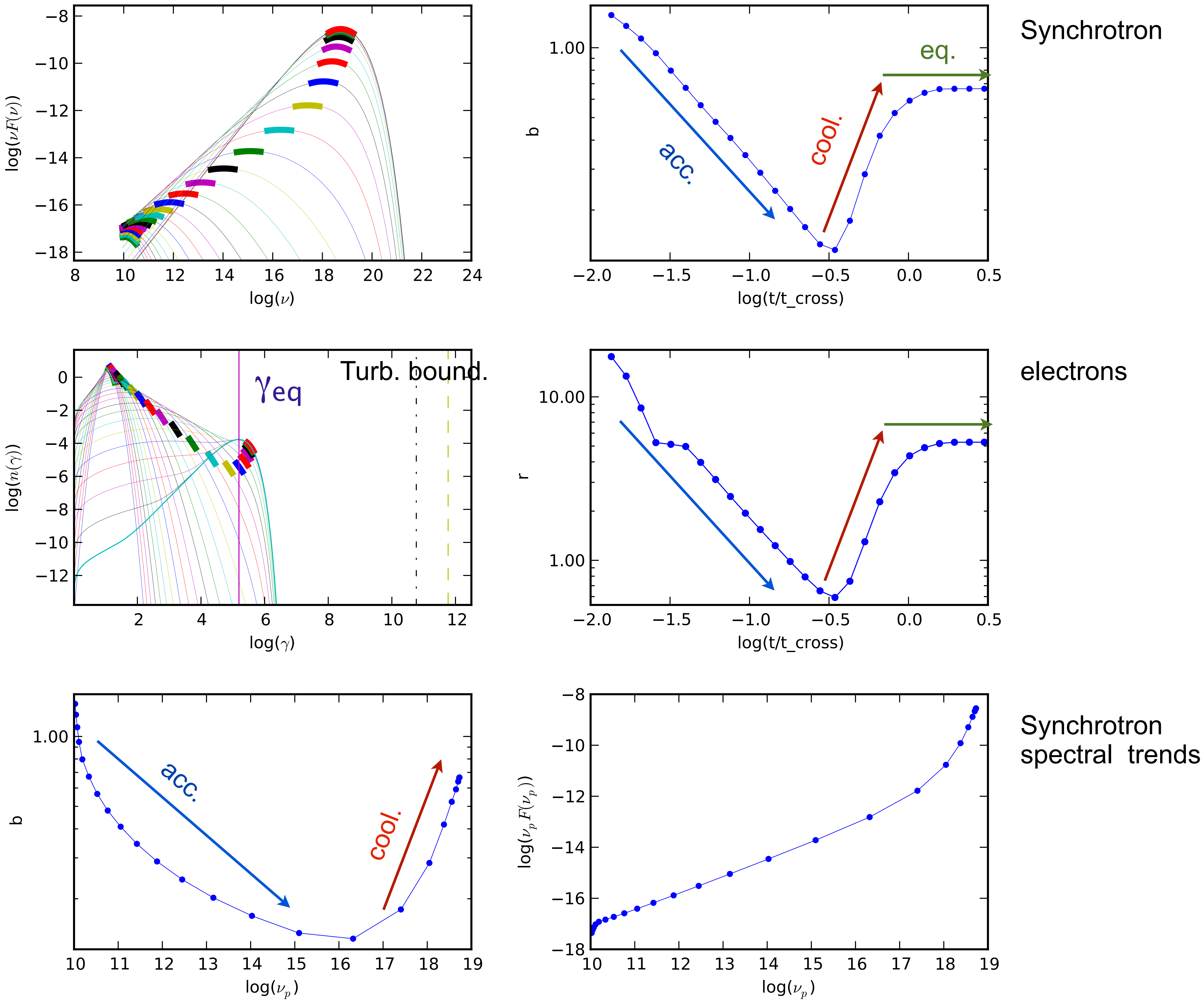
$$\boxed{n(\gamma,t) = \frac{N_0}{\gamma\sqrt{4\pi D_{p0}t}}\exp\Big\{-\frac{[\ln(\gamma/\gamma_0)-(A_{p0}-D_{p0})t]^2}{4D_{p0}t}\Big\}}$$

set-up of the accelerator



spectral trends

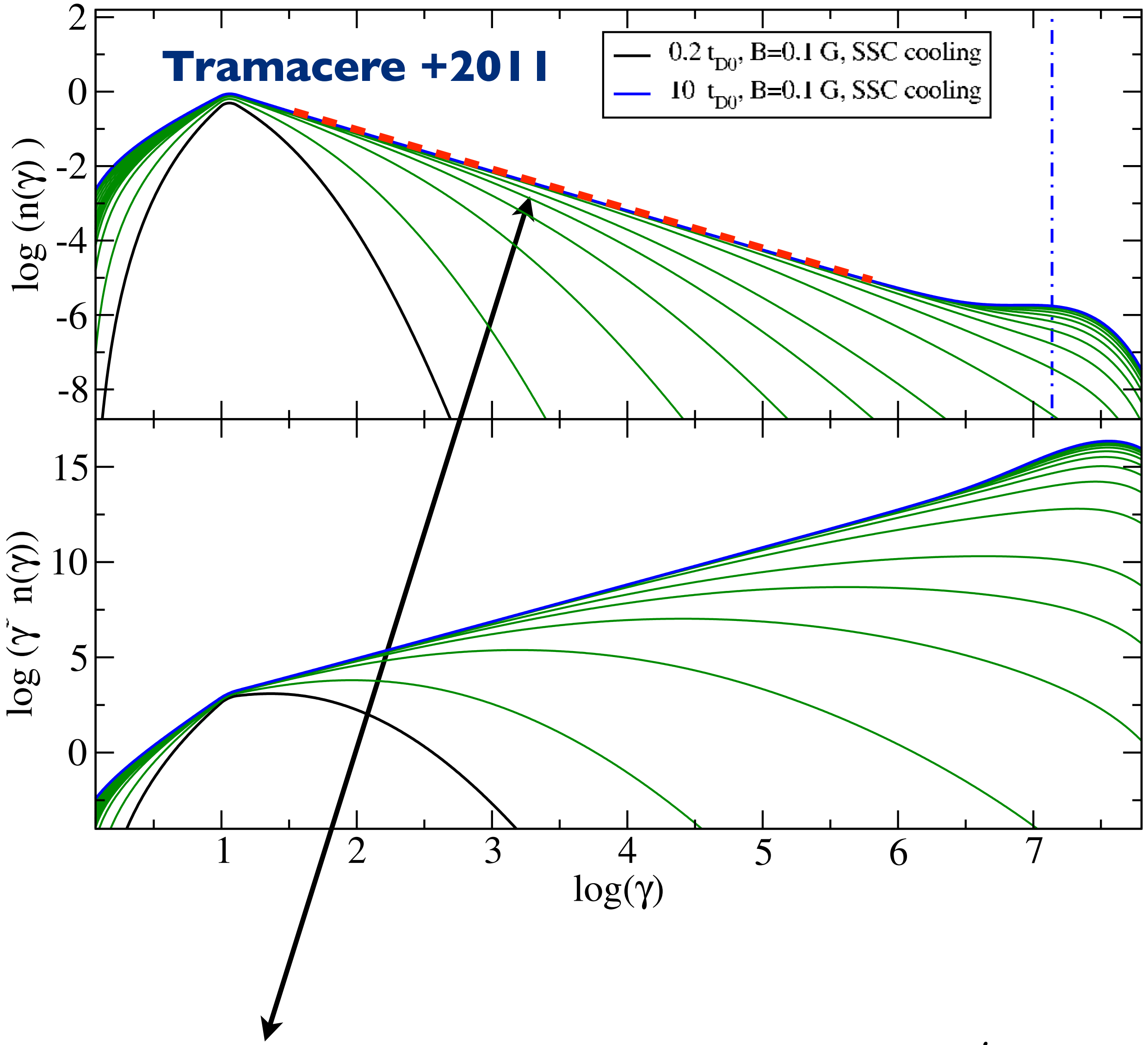
single flare



Synchrotron
electrons
Synchrotron
spectral trends
log(νF(ν))
log(ν)
b
log(t/t_cross)
acc.
cool.
eq.
log(n(γ))
log(γ)
γeq
Turb. bound.
r
log(ν_p)
log(ν_p F(ν_p))

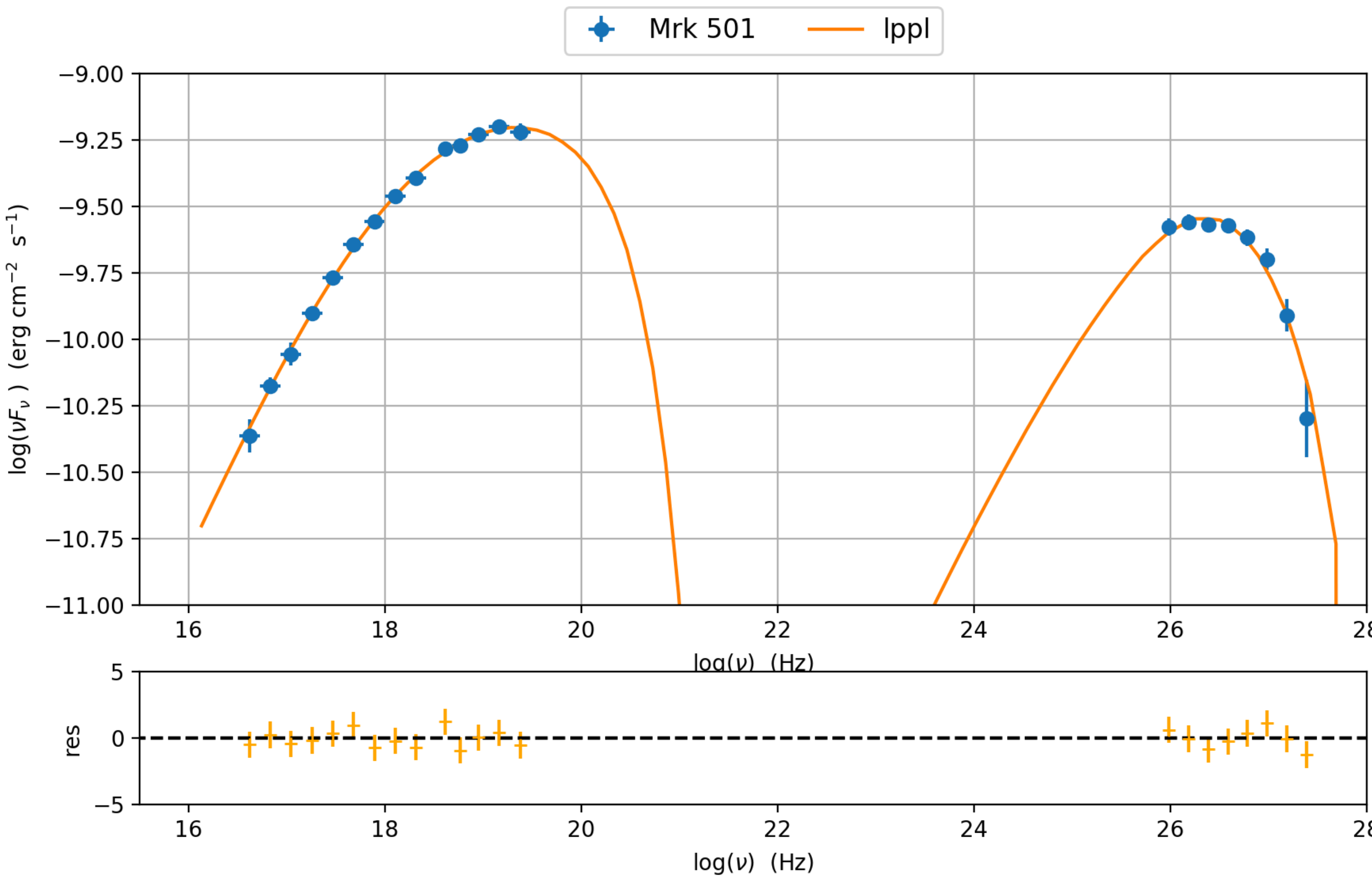
Pile-up and hard spectra

q=2, R=10^{15} cm, B=0.1 G, t_{inj}=t_D=10^4 s



s in agreement with $s = 1 + \frac{t_{acc}}{2t_{esc}}$

Mrk 501 1997



Massaro & Tramacere +2006

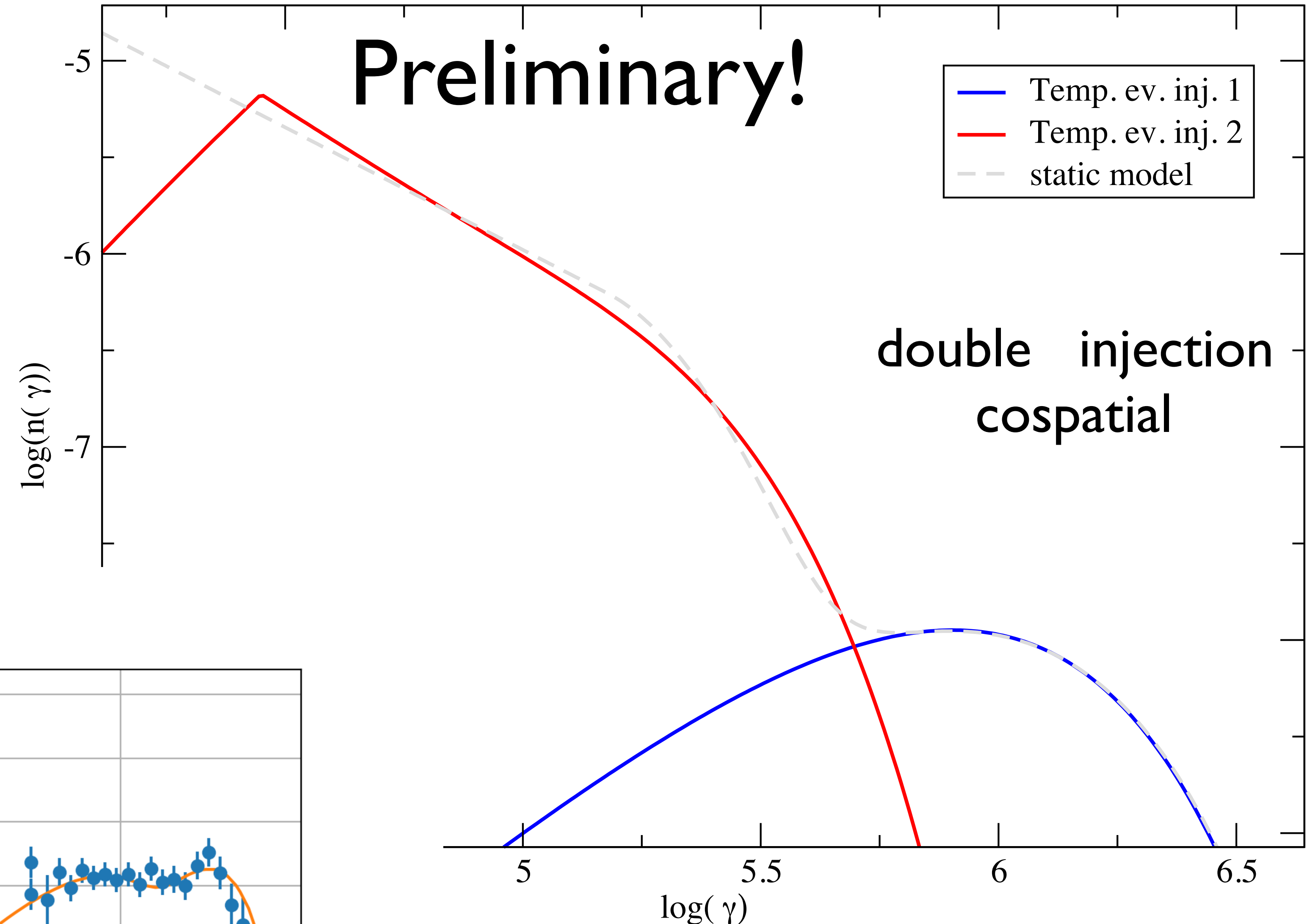
s~1.6

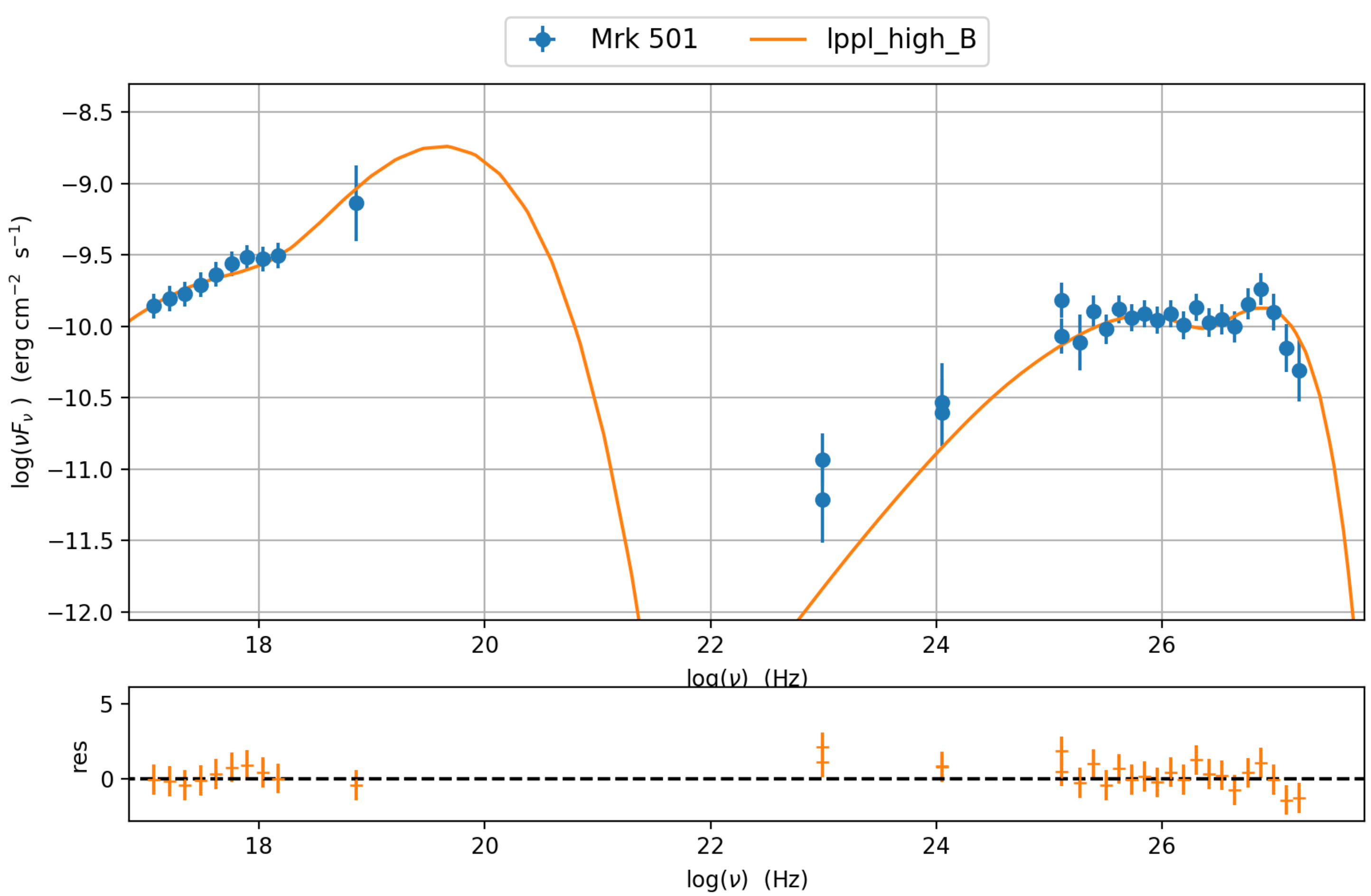
r~0.7-0.8<<r_{eq}~6

s<<s_{FI}~2.3

Pile-up and hard spectra

Mrk 501 2014 Flare
MAGIC paper
(submitted)





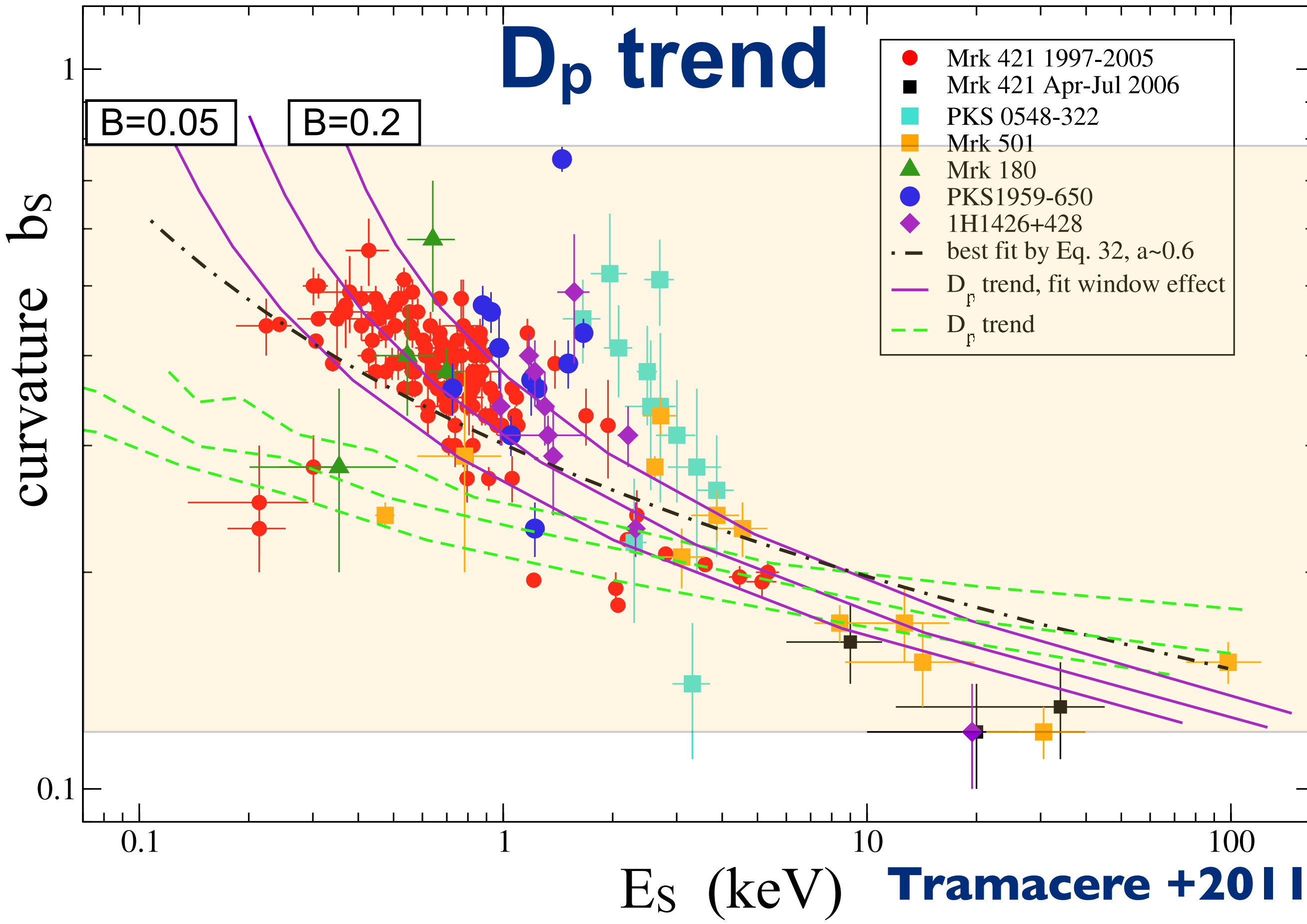
Summary of Stochastic signatures from self-consistent modeling

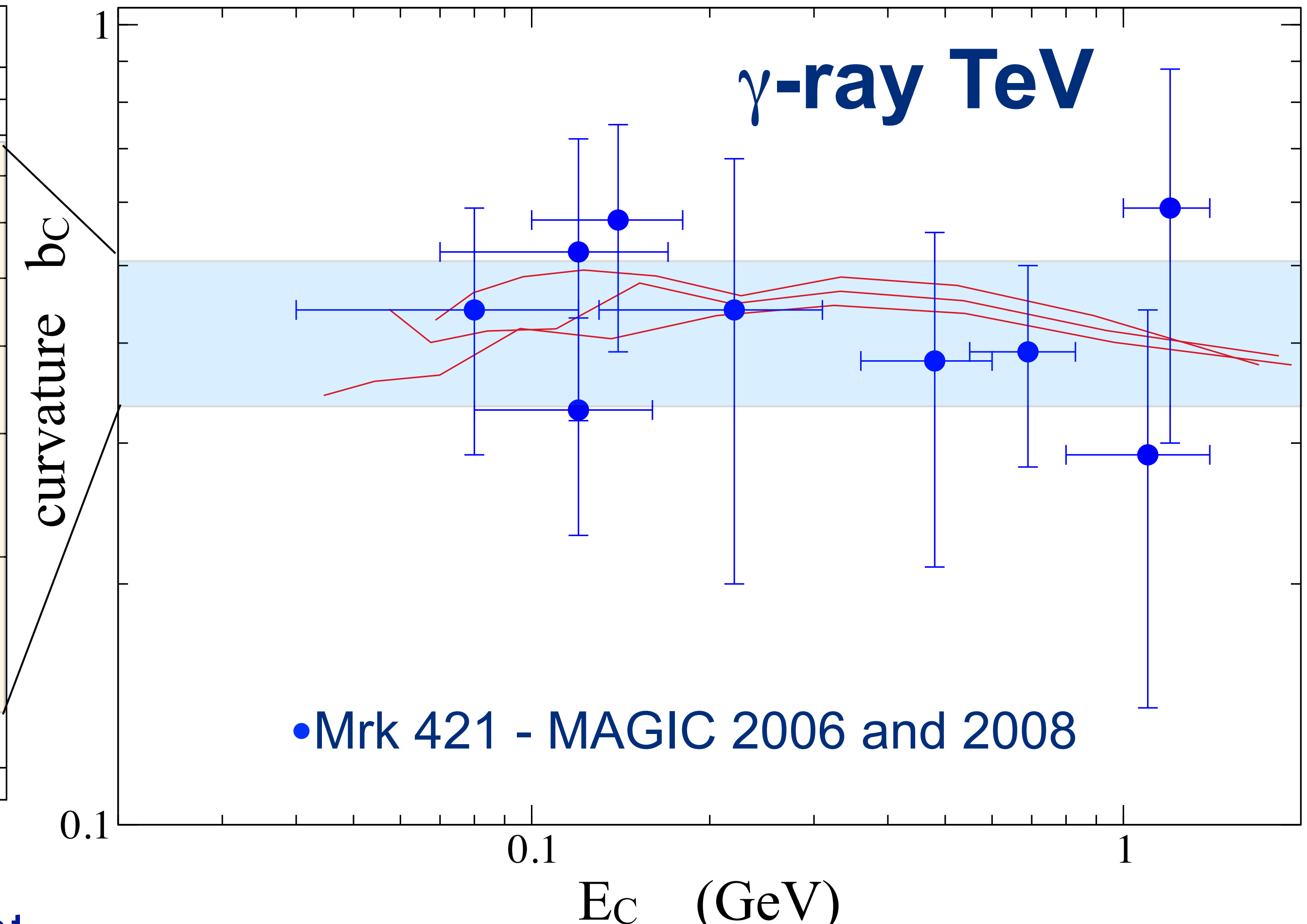
	Acceleration dominated	Equilibrium
curvature trend	curvature decreasing trend *b-Ep*	curvature stable or increasing (r~7,b~1.3)
spectral shape	LPPL or LP	PL+exp-cutoff or Maxwellian

spectral trends

multiple flares and population trends

E_S-b_S X-ray trend and γ-ray predictions

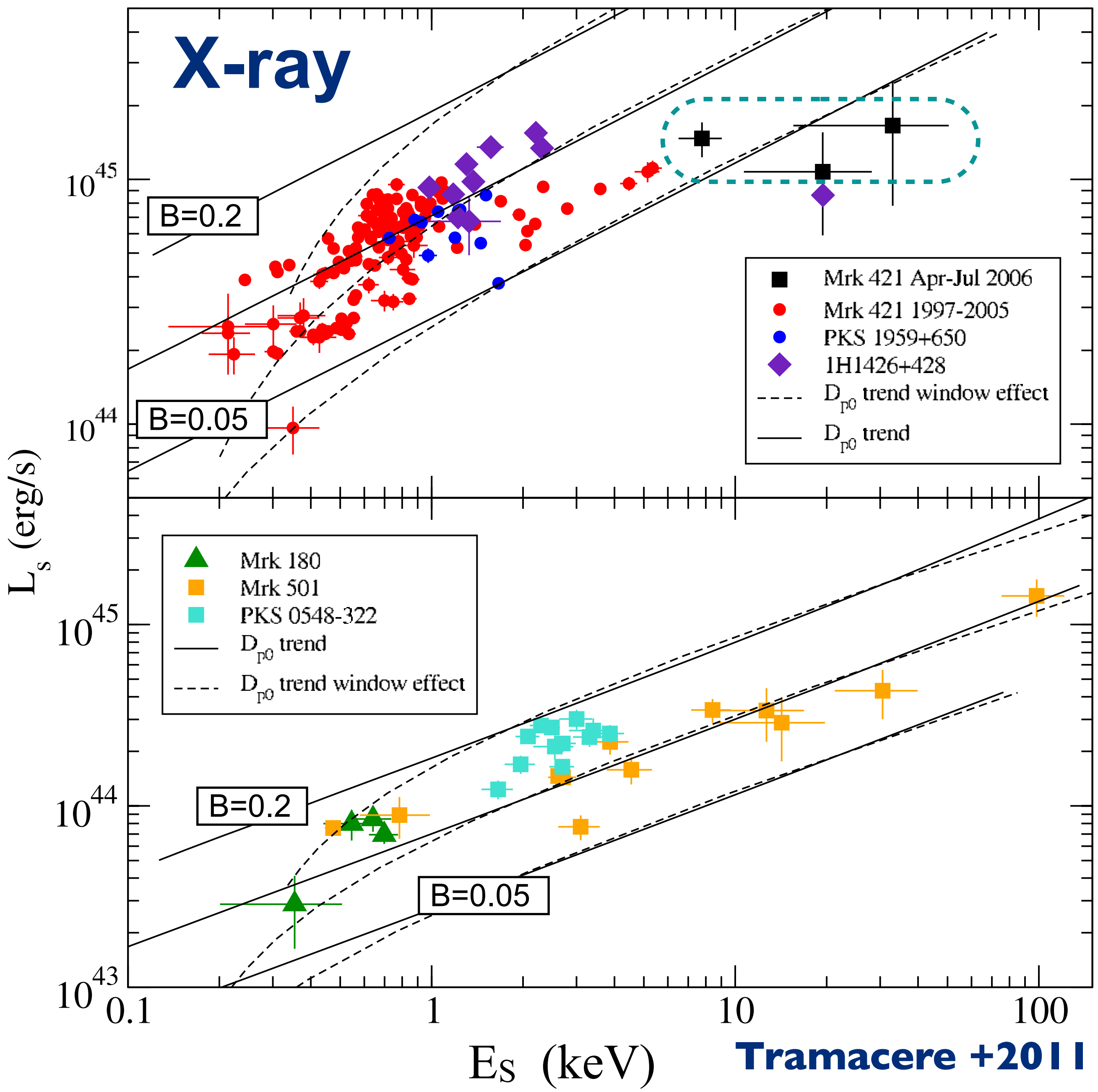


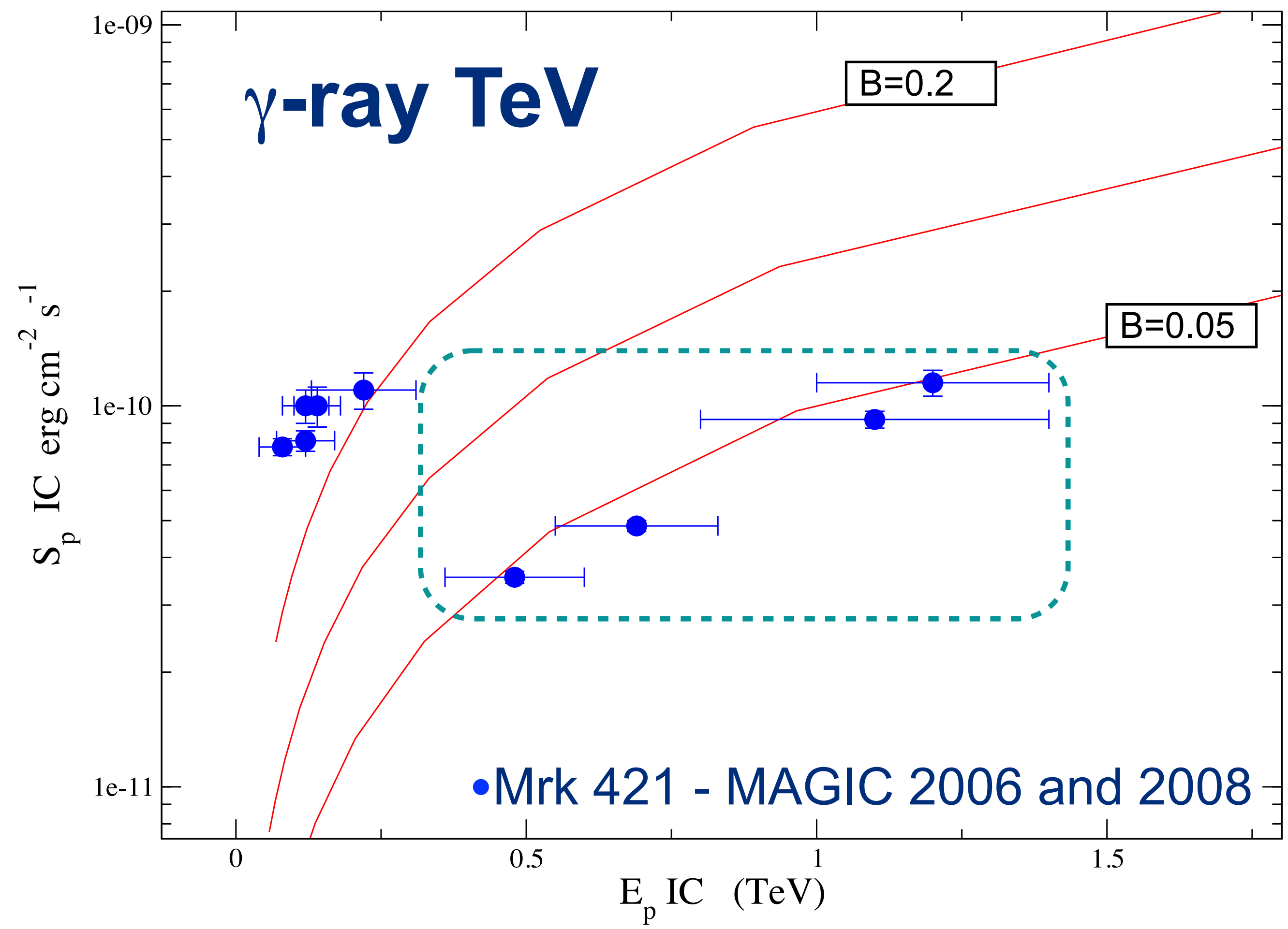


- data span **13 years**, both flaring and quiescent states
- We are able to reproduce these long-term behaviours, by changing the value of only one parameter (D_p)
- for q=2, curvature values imply distribution far from the equilibrium (b~[1.0–0.7])
- More data needed at GeV/TeV, curvature seems to be cooling-dominated
- Similar trend observed in GRBs (Massaro & Grindlay 2001)

L_{inj} (E_s–b_s trend)	(erg s^{-1})	5×10^{39}
L_{inj} (E_s–L_s trend)	(erg s^{-1})	5×10^{38}, 5×10^{39}
q		2
t_A	(s)	1.2×10^3
$t_{D_0} = 1/D_{P0}$	(s)	$[1.5 \times 10^4, 1.5 \times 10^5]$
T_{inj}	(s)	10^4
T_{esc}	(R/c)	2.0

E_S-L_S X-ray trend and γ-ray predictions





- the E_s–S_s (E_s–L_s) relation follows naturally from that between E_s and b_s
- the low L_{inj} objets (Mrk 501 vs Mrk 421) reach a larger E_S, compatibly with larger γ_{eq}
- Mrk 421 MAGIC data on 2006 match very well the Synchrotron prediction with simultaneous X-ray data
- the average index of the trend $L_s \propto E_S^{\alpha}$ with $\alpha \sim 0.6$, is compatible with the data, and with a scenario in which a typical constant energy (L_{inj} xt_{inj}) is injected for any flare (jet-feeding problem), whilst the peak dynamic is ruled by the turbulence in the magnetic field.

https://jetset.readthedocs.io/en/latest/

https://github.com/andreatramacere/jetset/archive/stable.tar.gz

to get the beta release write to

- andrea.tramacere@gmail.com
- andrea.tramacere@unige.ch

JetSeT 1.0 JetSeT Page Installation » Source Search

- installation
- user guide
- code documentation (API)
- Source

JetSeT Documentation

JetSeT

Jets SED modeler and fitting Tool

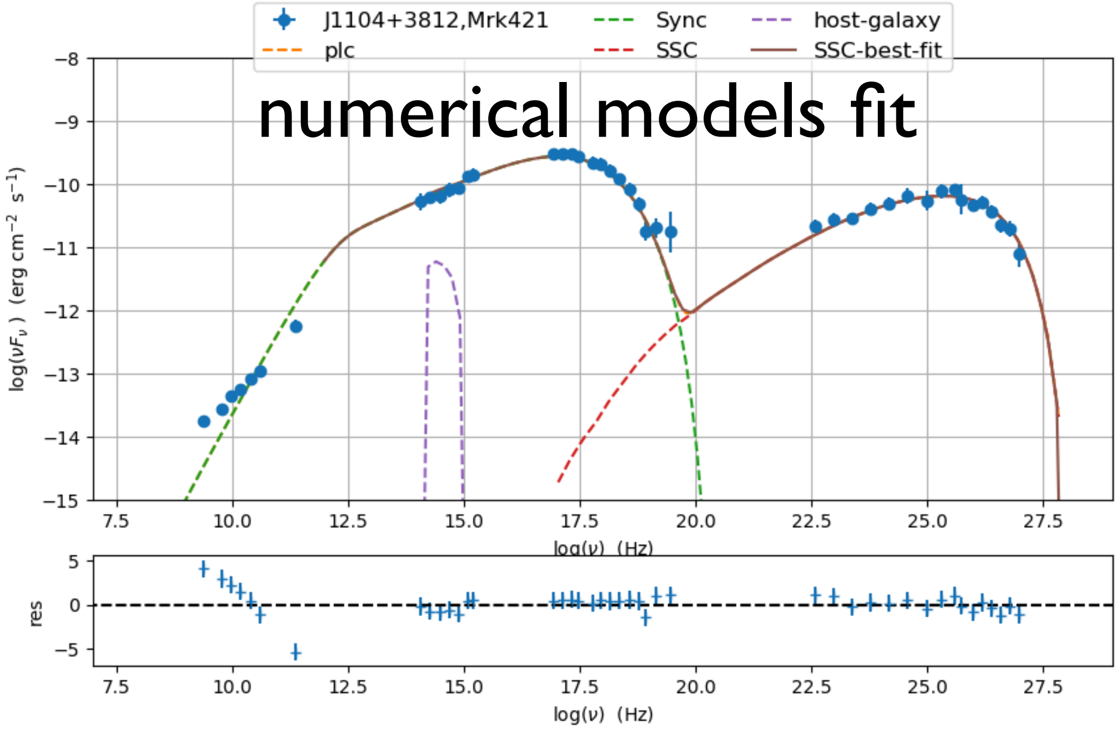
Author: Andrea Tramacere

JetSeT is an open source C/Python framework to reproduce radiative and accelerative processes acting in relativistic jets, allowing to fit the numerical models to observed data. The main features of this framework are:

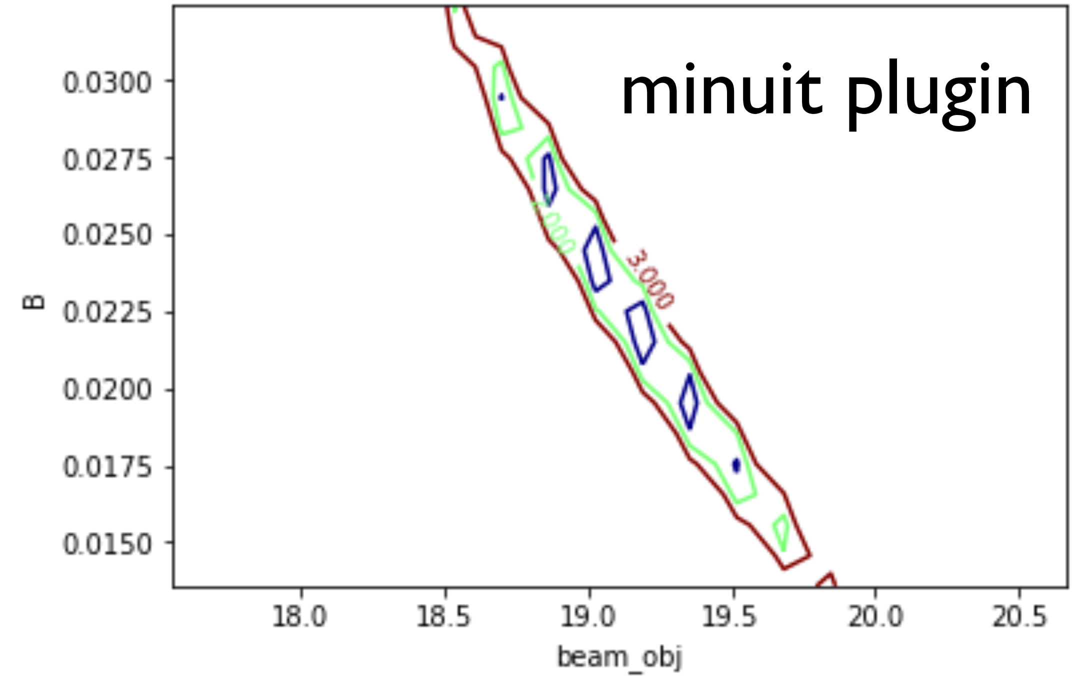
- handling observed data: re-binning, definition of data sets, bindings to astropy tables and quantities definition of complex numerical radiative scenarios: Synchrotron Self-Compton (SSC), external Compton (EC) and EC against the CMB
- Constraining of the model in the pre-fitting stage, based on accurate and already published phenomenological trends. In particular, starting from phenomenological parameters, such as spectral indices, peak fluxes and frequencies, and spectral curvatures, that the code evaluates automatically, the pre-fitting algorithm is able to provide a good starting model,following the phenomenological trends that I have implemented. fitting of multiwavelength SEDs using both frequentist approach (iminuit) and bayesian MCMC sampling (emcee)
- Self-consistent temporal evolution of the plasma under the effect of radiative and accelerative processes, both first order and second order (stochastic acceleration) processes.

v: latest

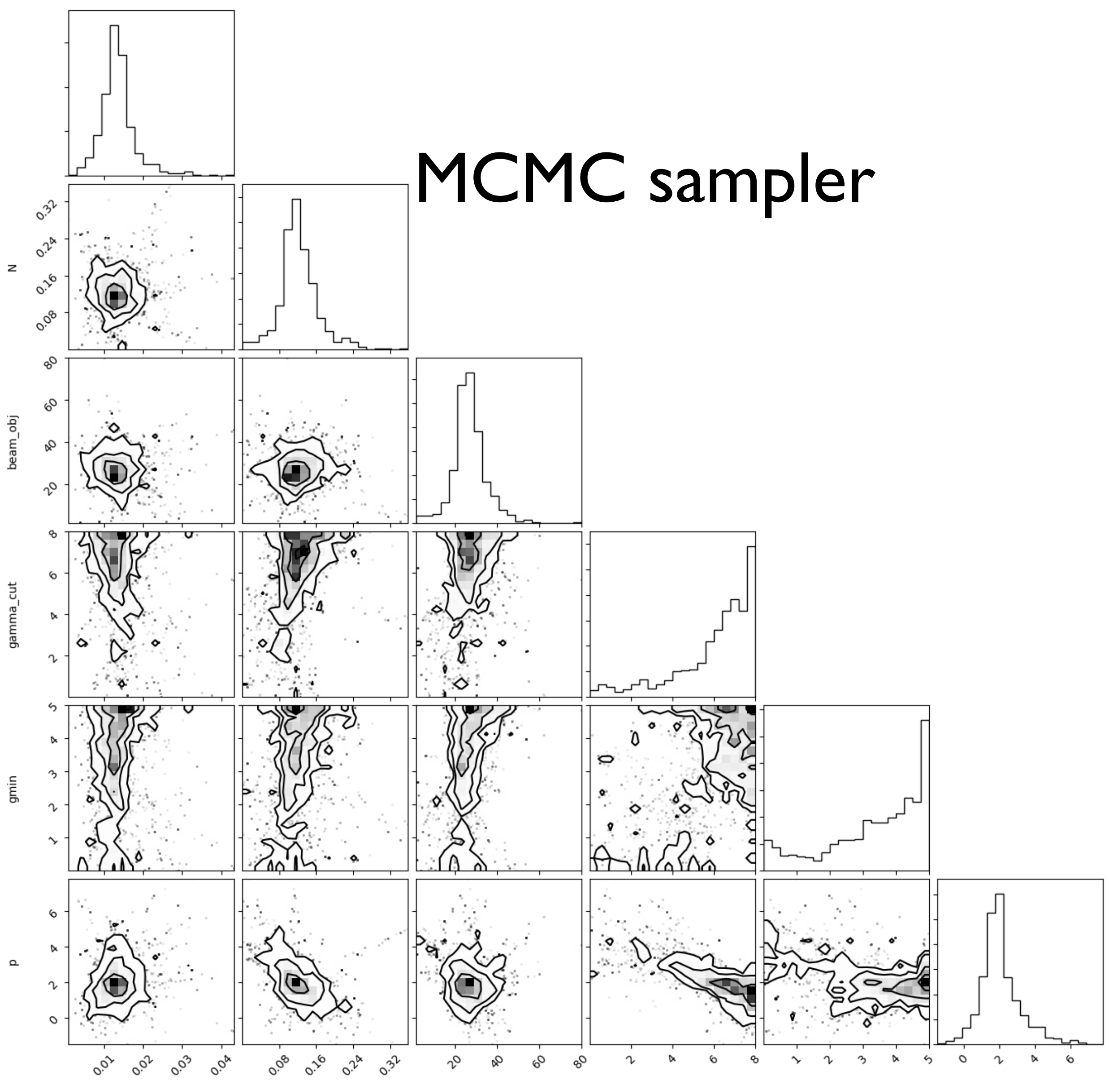
numerical models fit



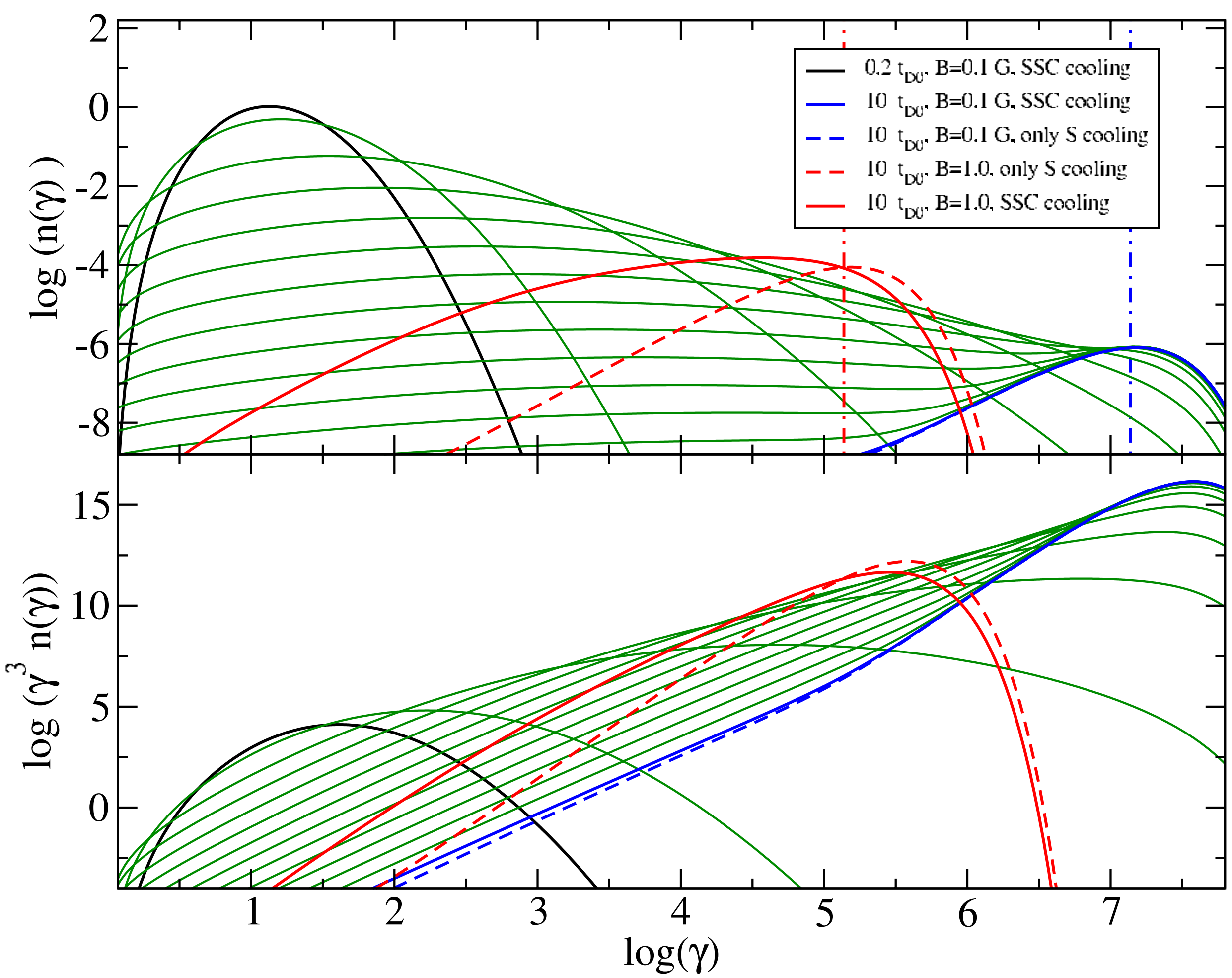
minuit plugin



MCMC sampler



Temp. ev. of the plasma



backup slides

injection term

$$L_{inj} = \frac{4}{3}\pi R^3 \int \gamma_{inj} m_e c^2 Q(\gamma_{inj}, t) d\gamma_{inj} \quad (erg/s)$$

systematic term

$$S(\gamma, t) = -C(\gamma, t) + A(\gamma, t)$$

cooling term

$$C(\gamma) = |\dot{\gamma}_{\mathrm{synch}}| + |\dot{\gamma}_{\mathrm{IC}}|$$

syst. acc. term

$$A(\gamma) = A_{p0}\gamma,\ t_A = \frac{1}{A_0}$$

$$\frac{\partial n(\gamma,t)}{\partial t} = \frac{\partial}{\partial \gamma}\left\{ -[S(\gamma,t) + D_A(\gamma,t)]n(\gamma,t) + D_p(\gamma,t)\frac{\partial n(\gamma,t)}{\partial \gamma}\right\} - \frac{n(\gamma,t)}{T_{\mathrm{esc}}(\gamma)} + Q(\gamma,t)$$

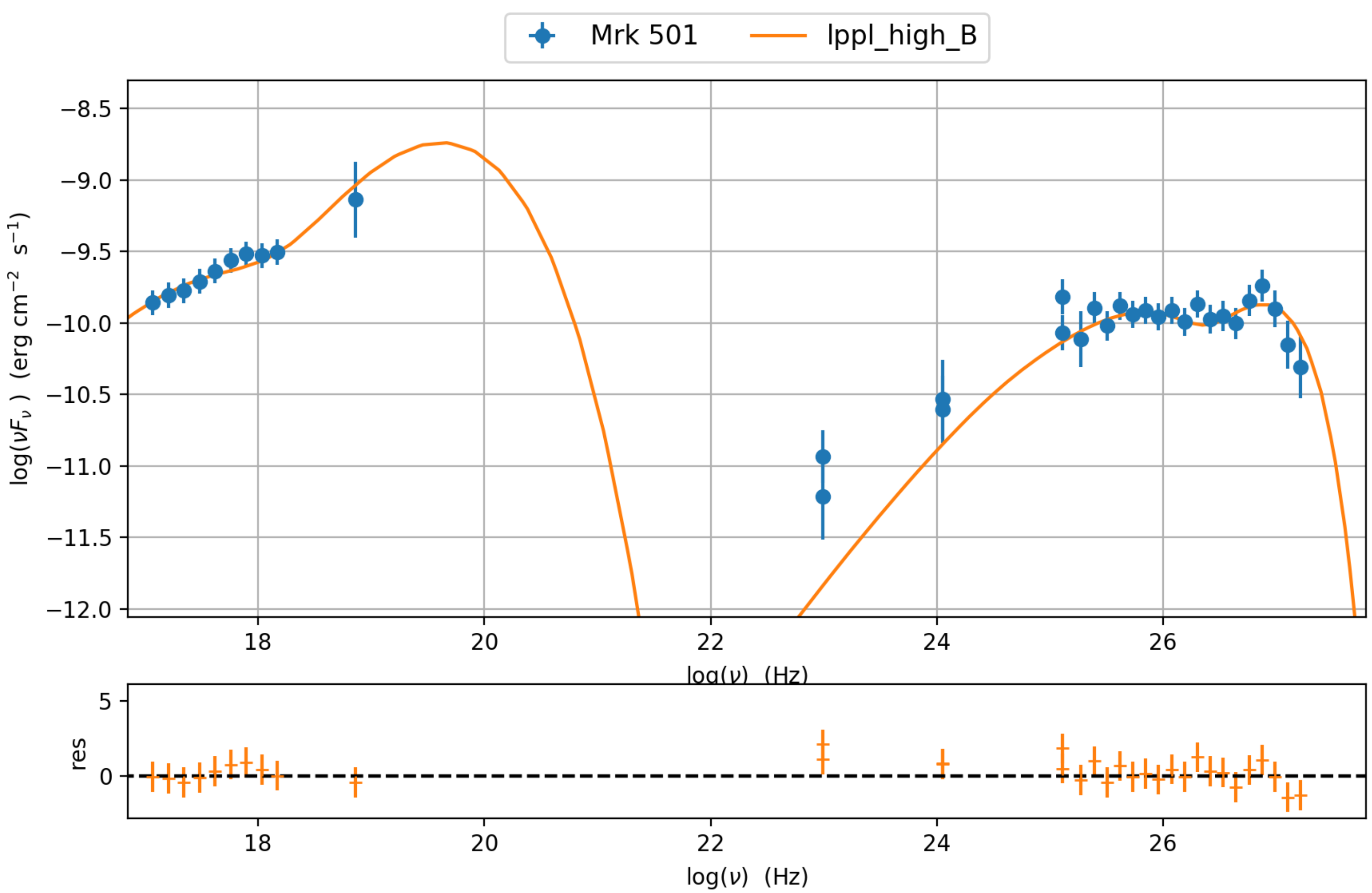
Turbulent magnetic field $\longrightarrow$ **momentum diffusion term**

$$W(k) = \frac{\delta B(k_0^2)}{8\pi}\left(\frac{k}{k_0}\right)^{-q}$$

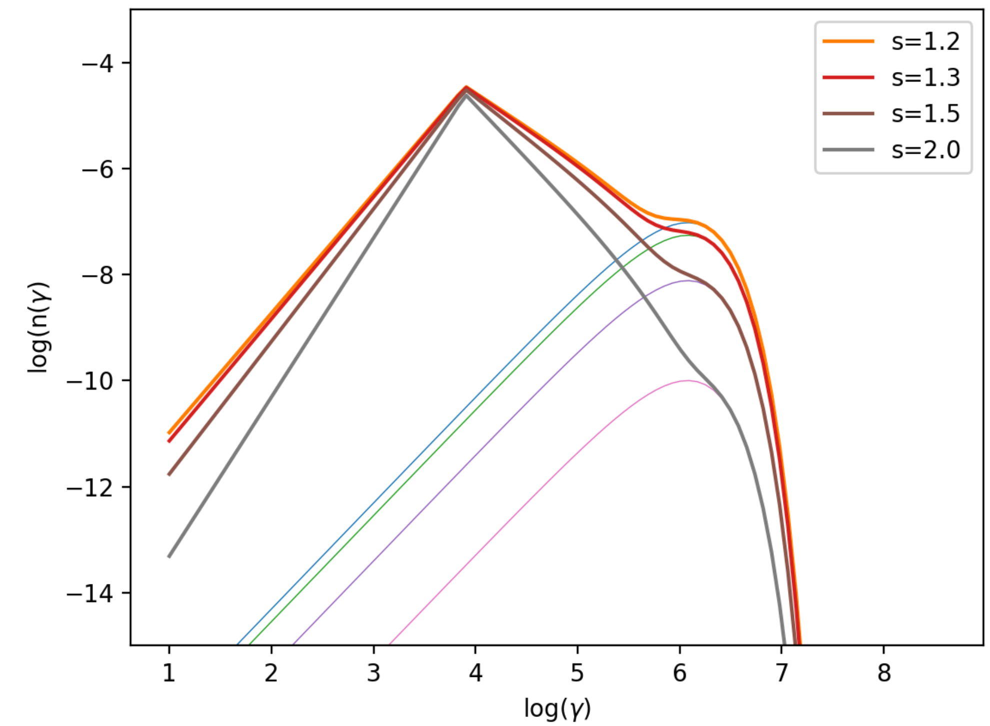
Mrk 501 2014 Flare MAGIC paper (submitted)

model parameters:

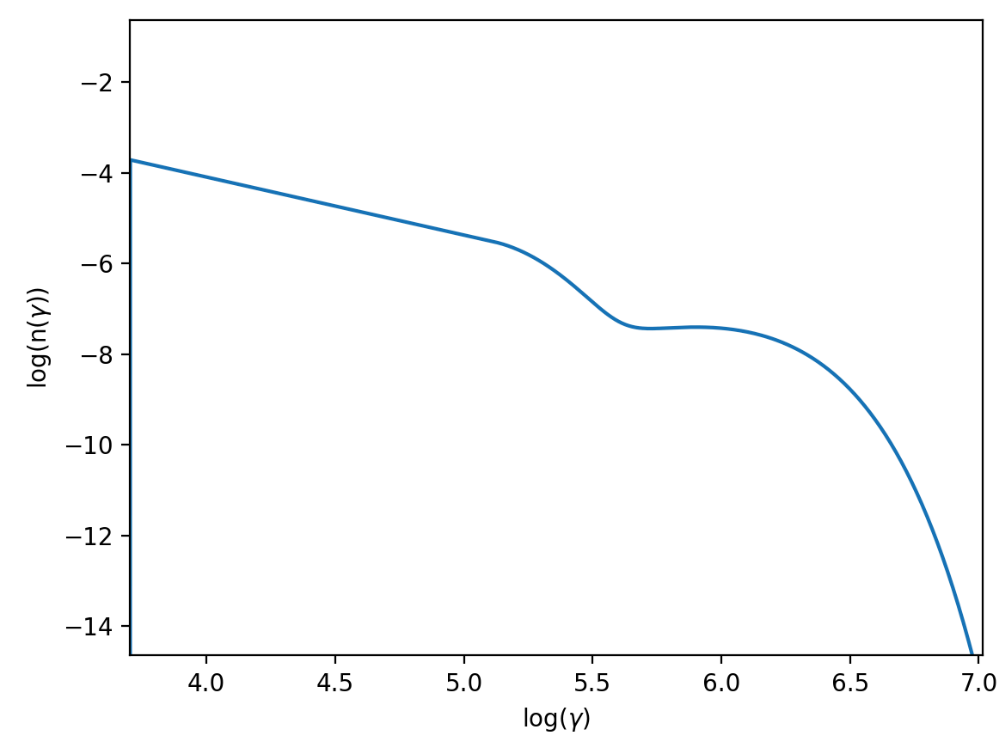
Name	Type	Units	value
B	magnetic_field	G	+3.000000e-01
N	electron_density	cm^-3	+2.360060e+00
R	region_size	cm	+1.551851e+01
alpha_pile_up	turn-over-energy		+1.000000e+00
beam_obj	beaming		+1.000000e+01
gamma0_log_parab	turn-over-energy	Lorentz-factor	+1.300000e+05
gamma_inj	turn-over-energy	Lorentz-factor	+5.000000e+03
gamma_pile_up	turn-over-energy	Lorentz-factor	+4.000000e+05
gmax	high-energy-cut-off	Lorentz-factor	+1.000000e+07
gmin	low-energy-cut-off	Lorentz-factor	+5.000000e+03
r	spectral_curvature		+6.100000e+00
ratio_pile_up	turn-over-energy		+7.000000e-18
s	LE_spectral_slope		+1.280000e+00
z_cosm	redshift		+3.364200e-02



cont. single injection (Stawarz&Petrosian 2009) not compatible with MW data

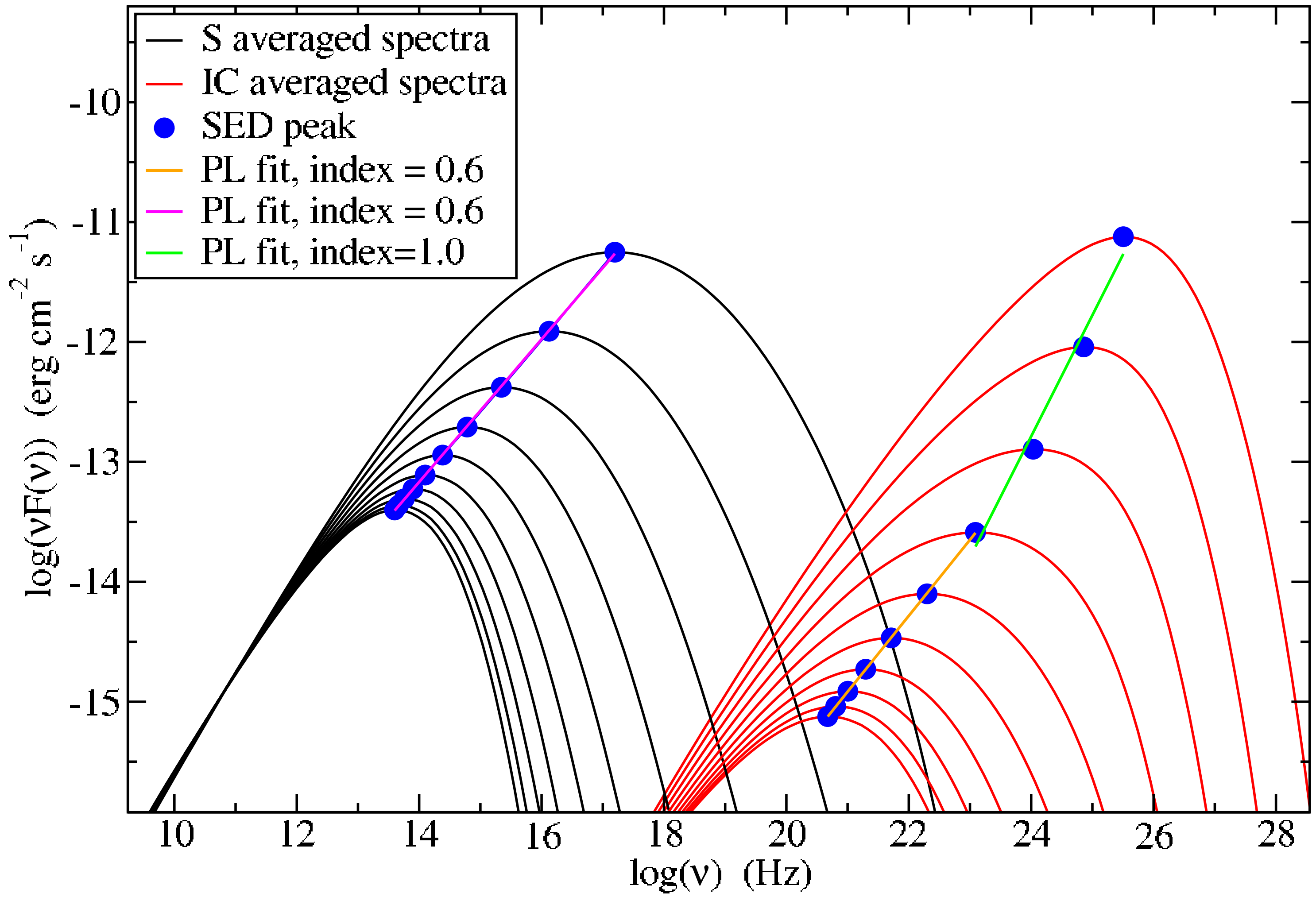


double cospatial injection compatible with data

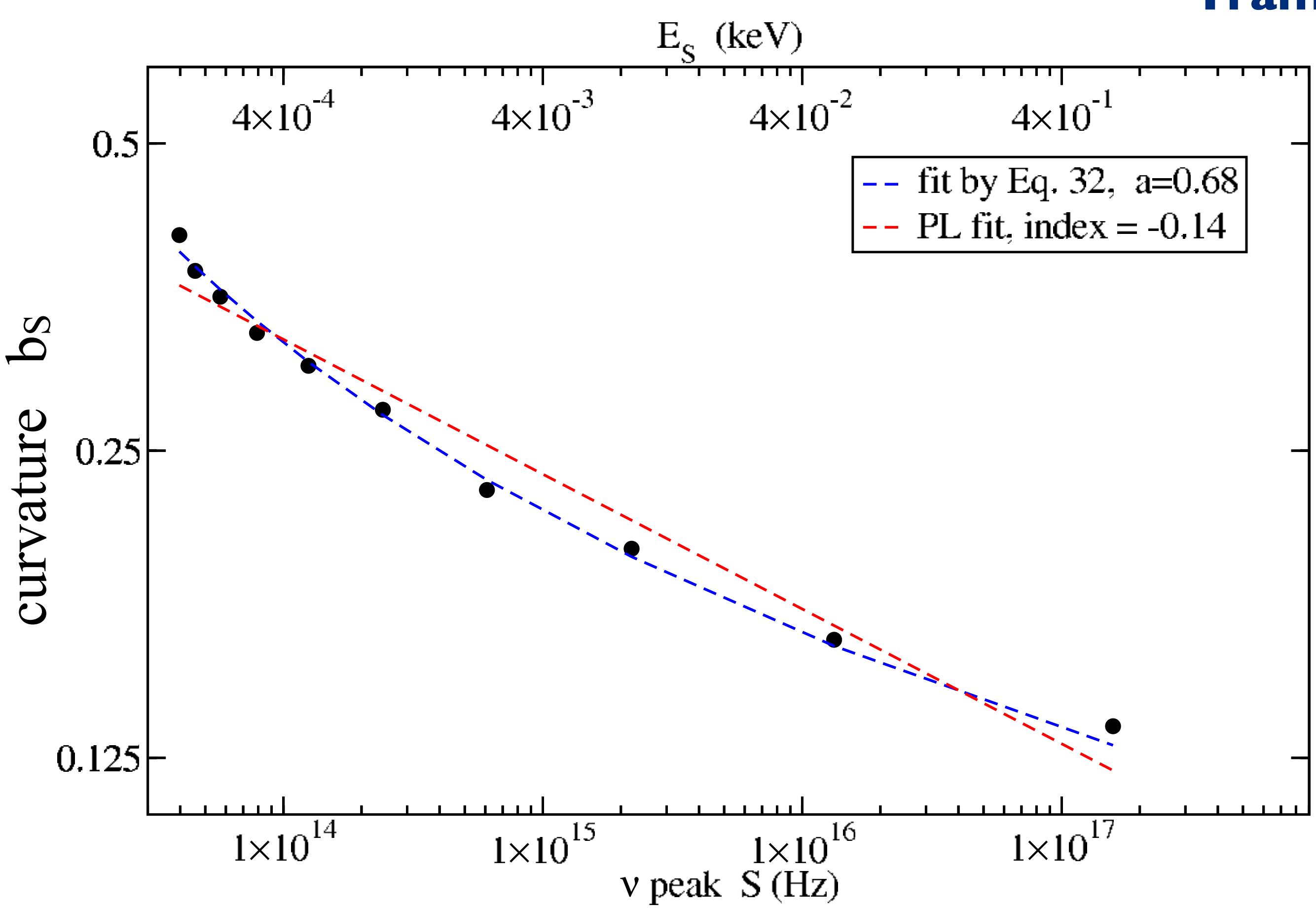


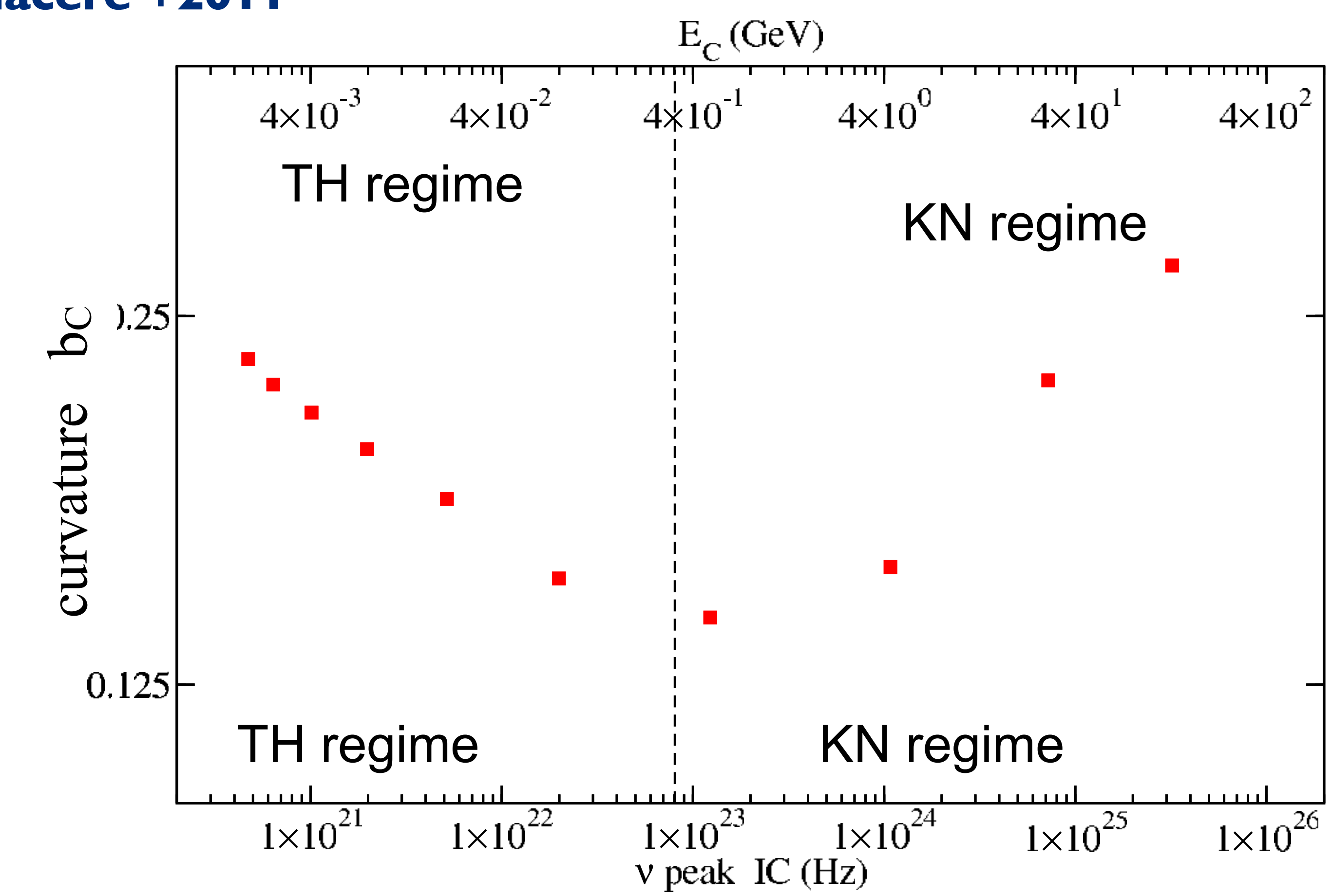
S vs IC

Tramacere+2011

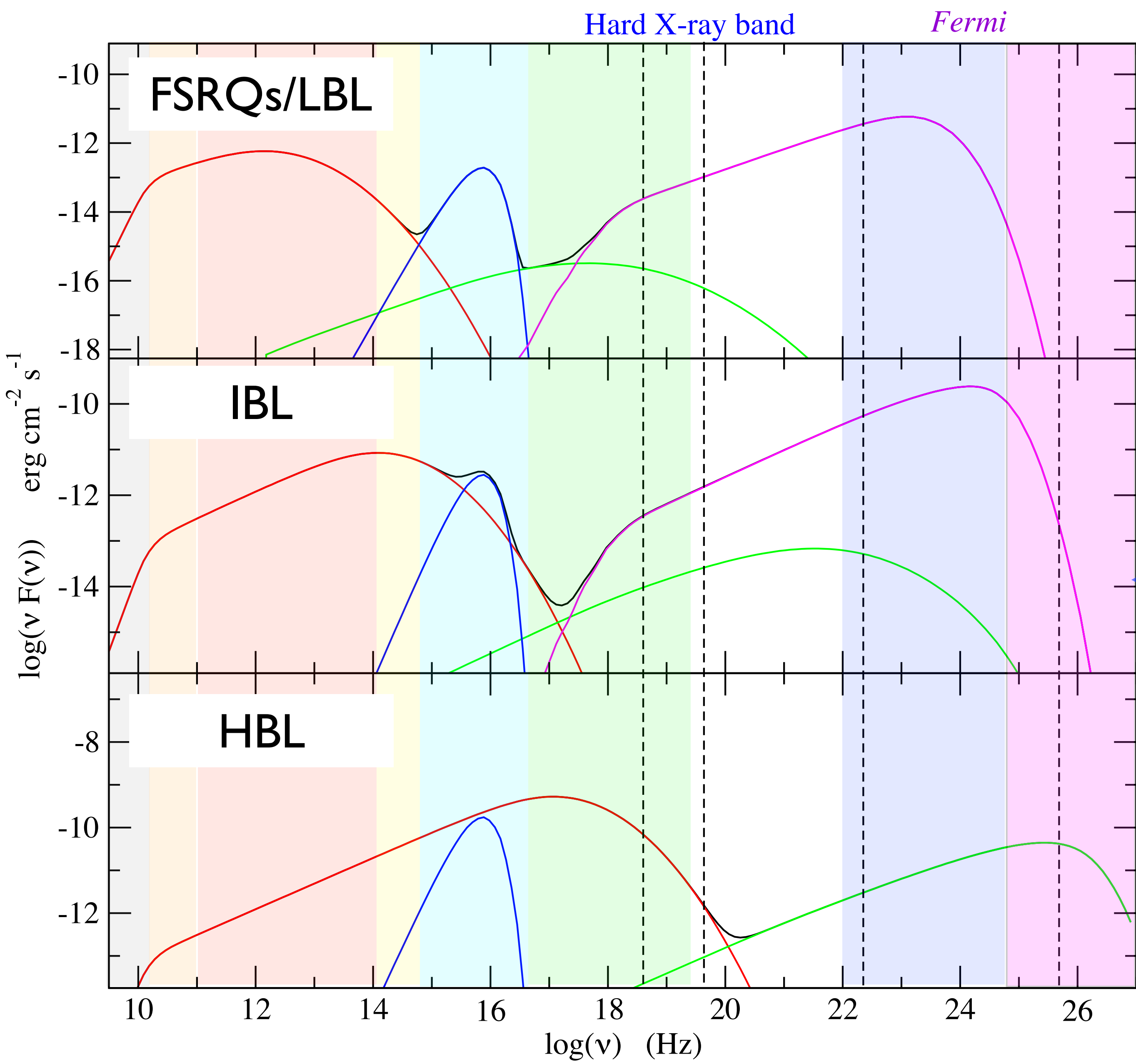


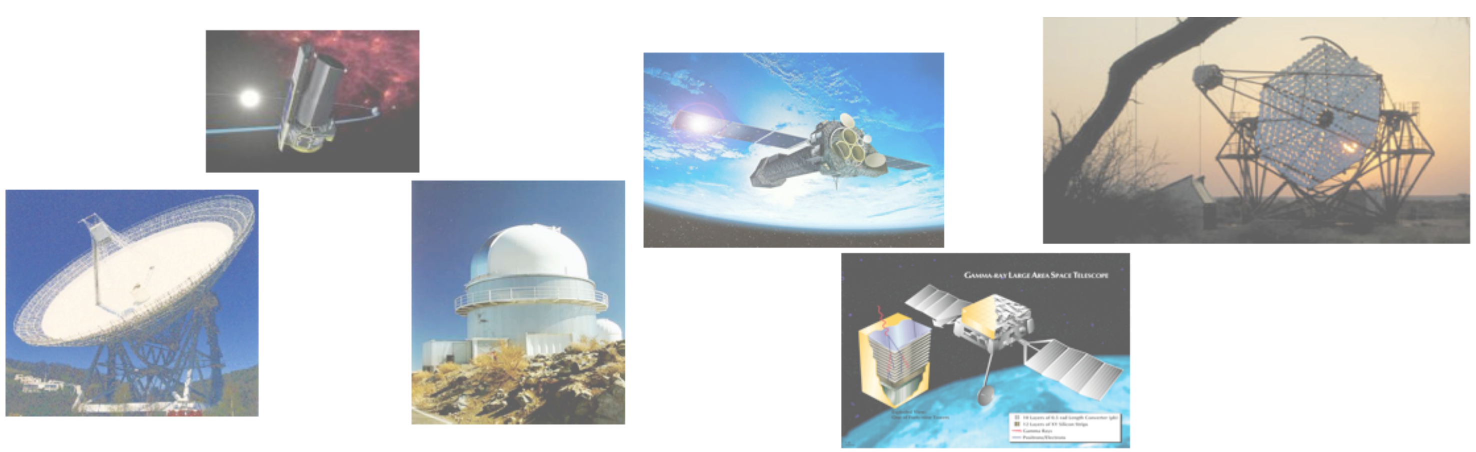
Tramacere +2011



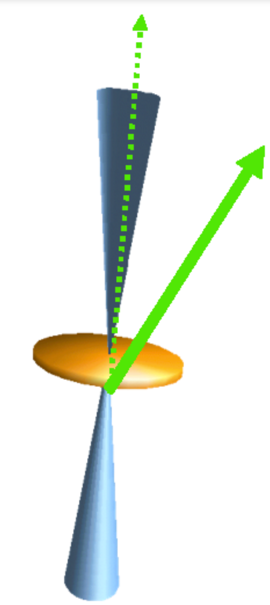


blazars in a nutshell



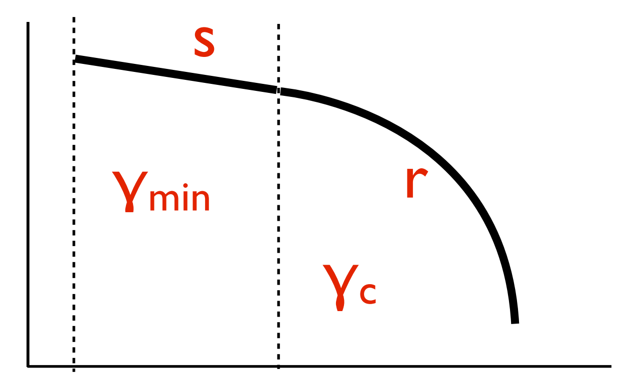


jet/disk



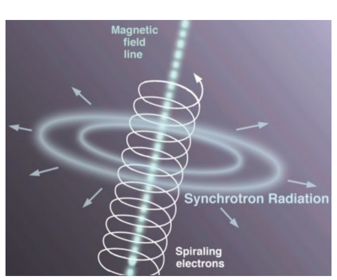
+

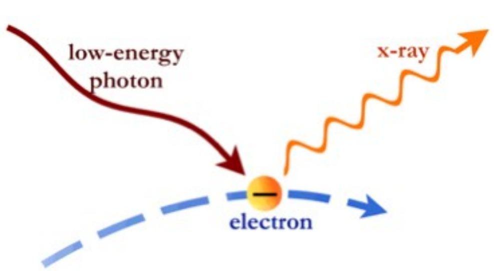
acc.



+

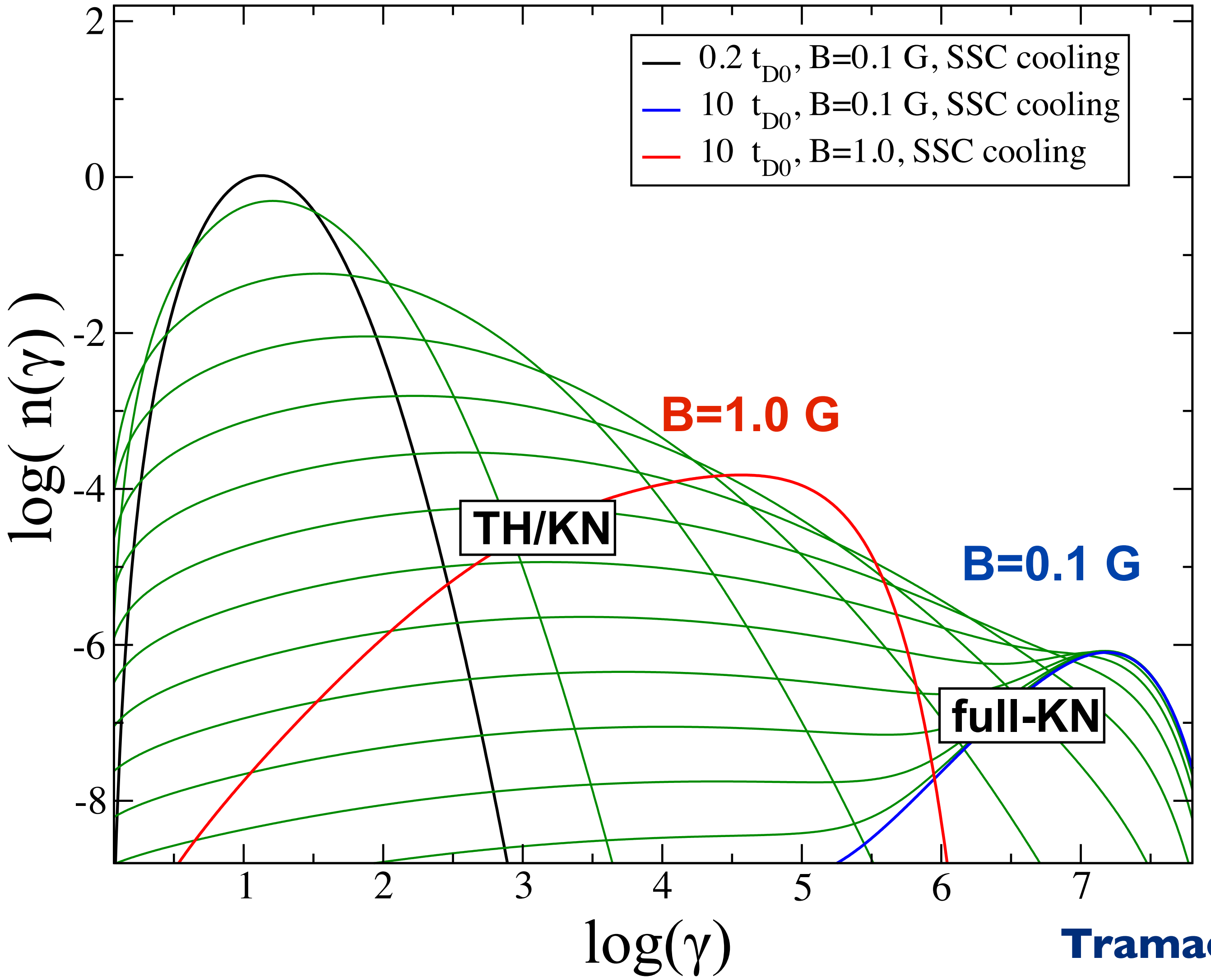
em.



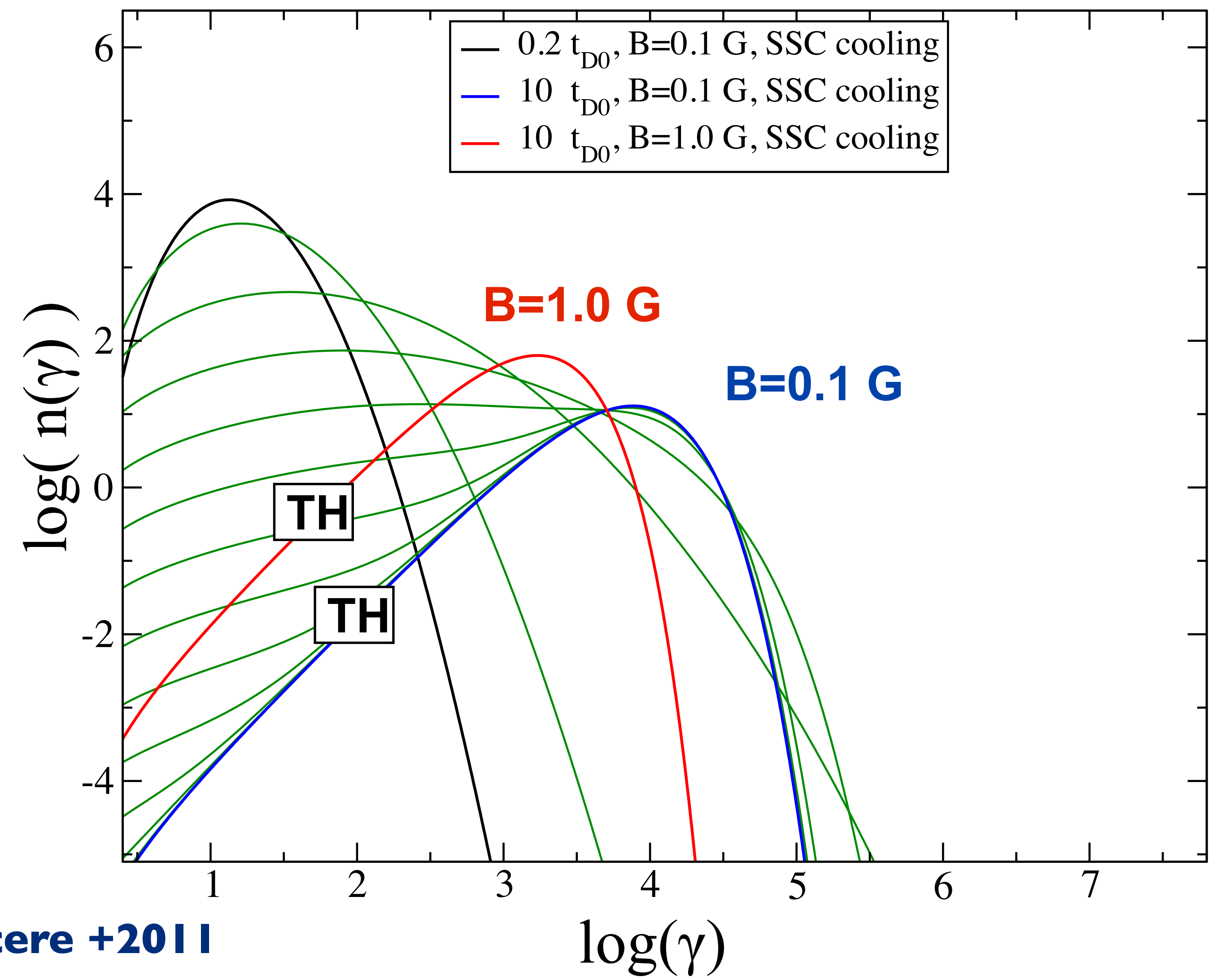


IC cooling and equilibrium

R= 1x10^{15} cm



R= 5x10^{13} cm



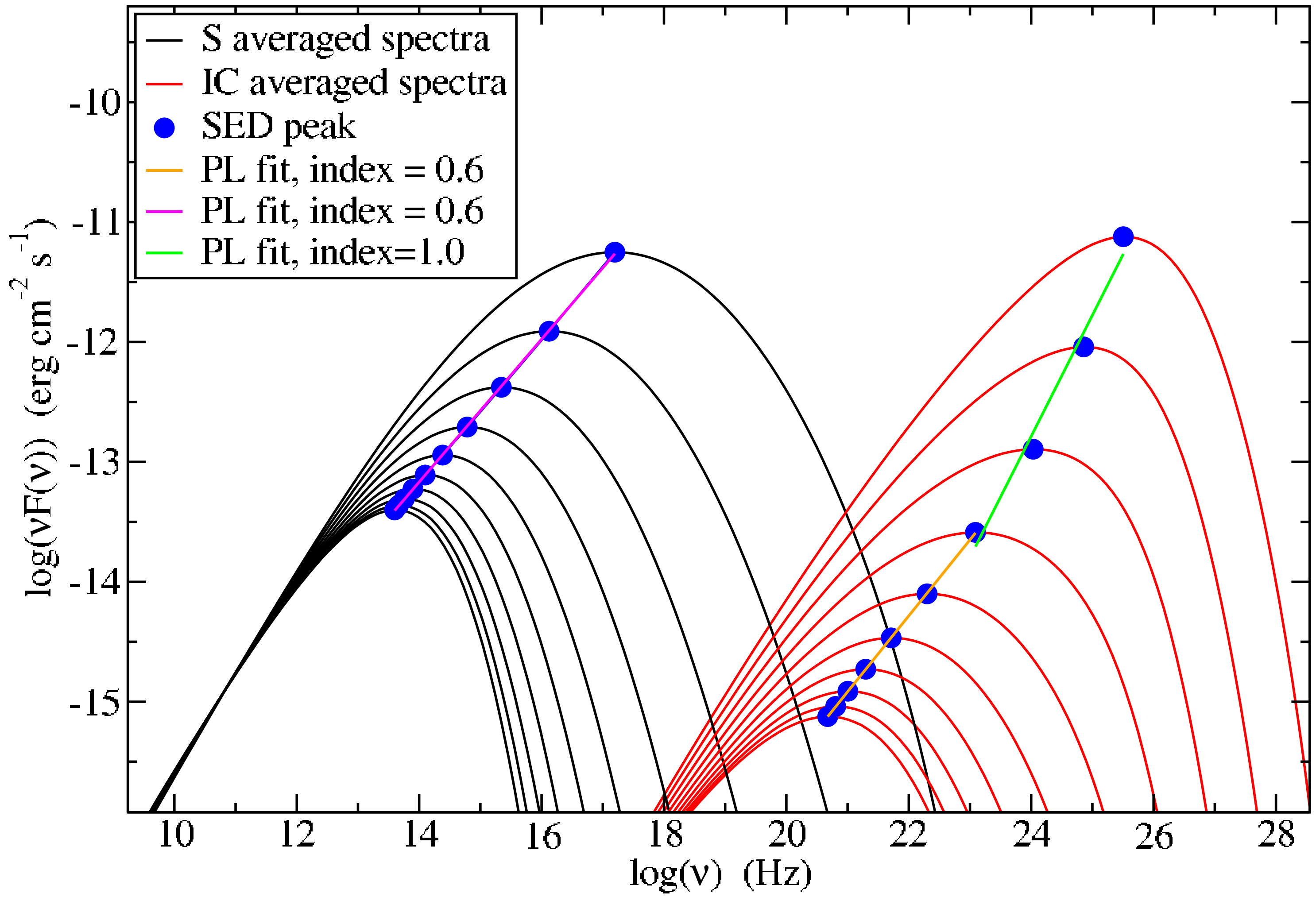
Tramacere +2011

U_{ph} (R= 1x10^{13} cm**) >> U_{ph} (**R= 1x10^{15} cm**)**

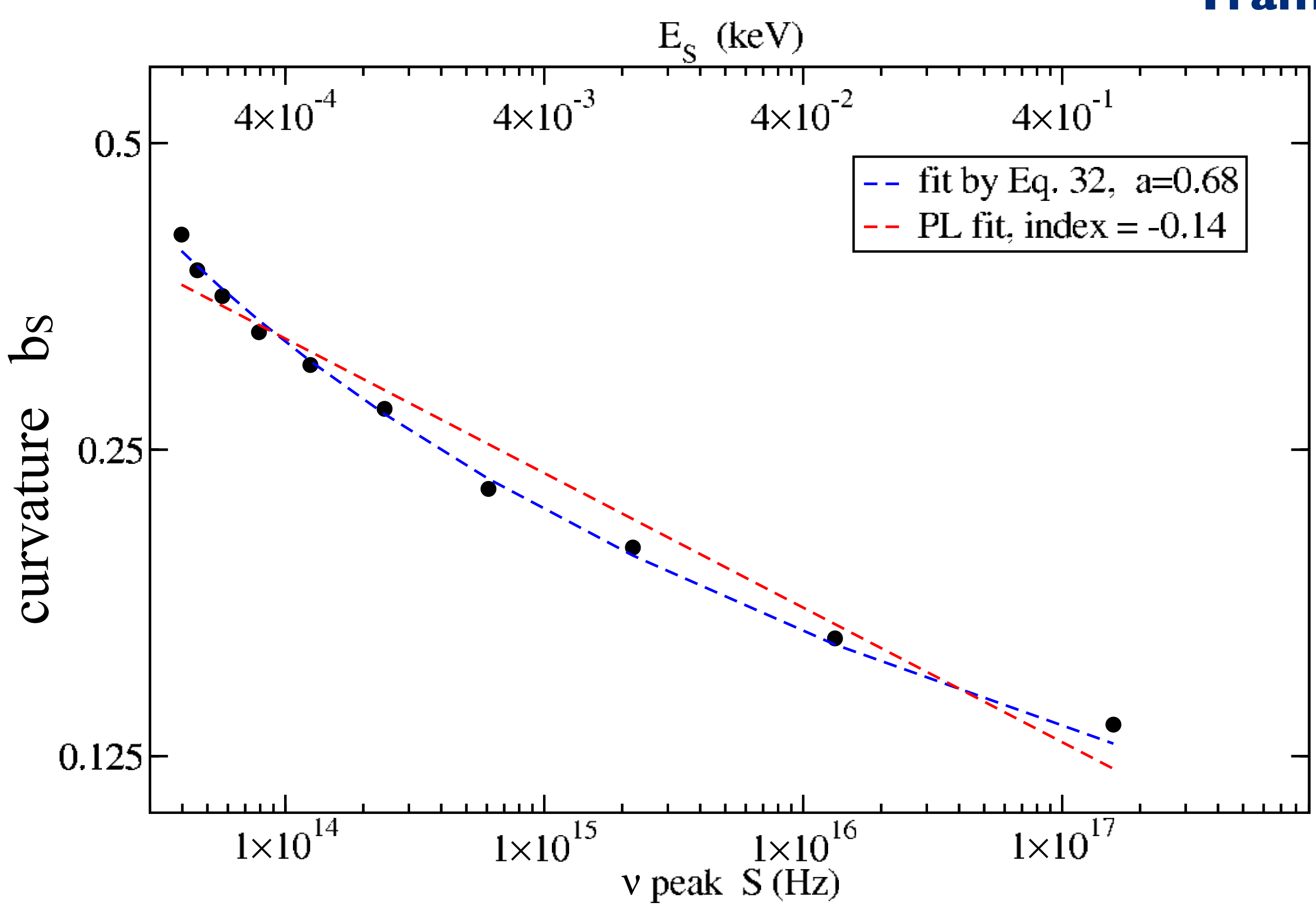
IC prevents higher energies in more compact accelerators (if all the parameters are the same) **Impact on rapid TeV variability!**

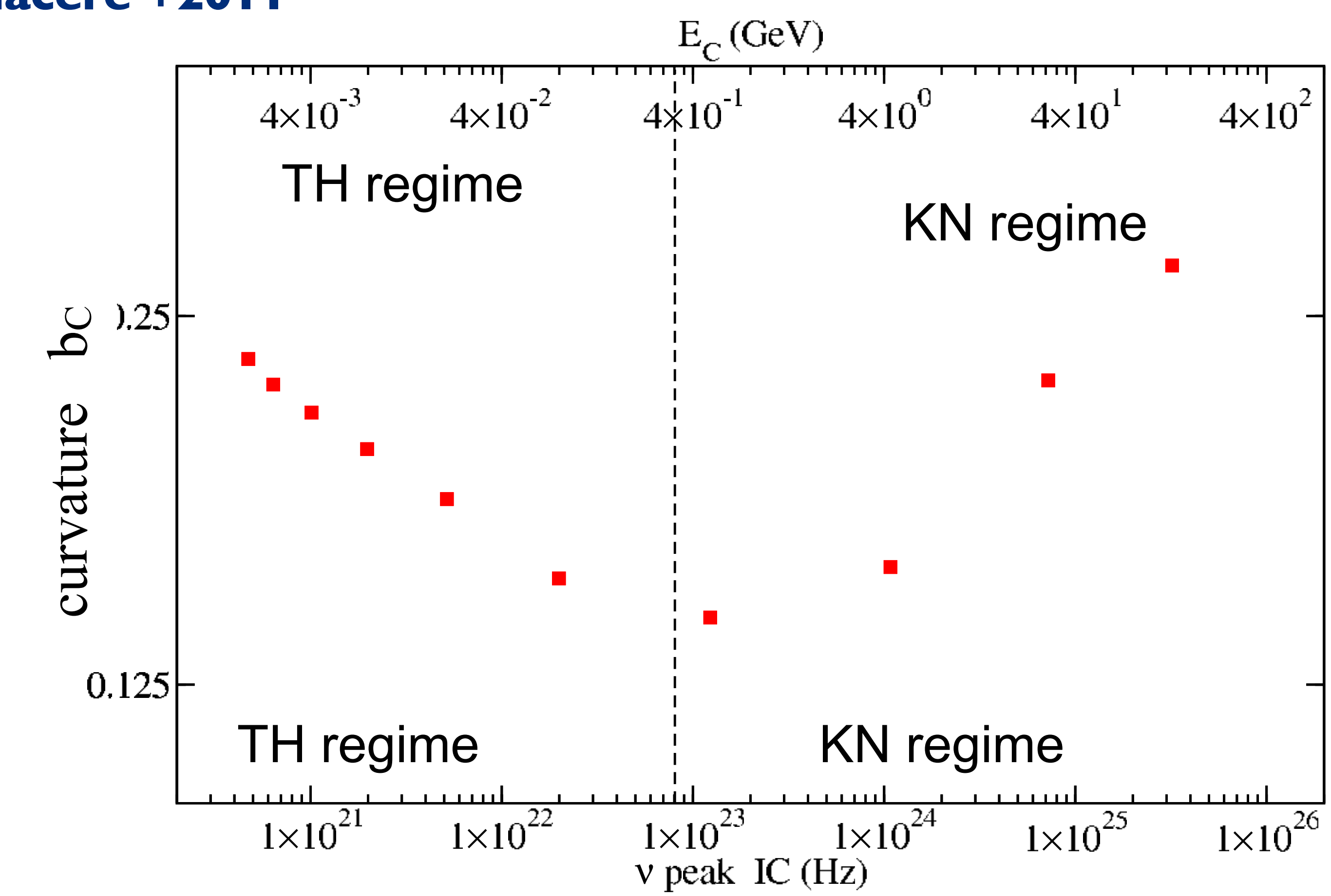
S vs IC

Tramacere+2011

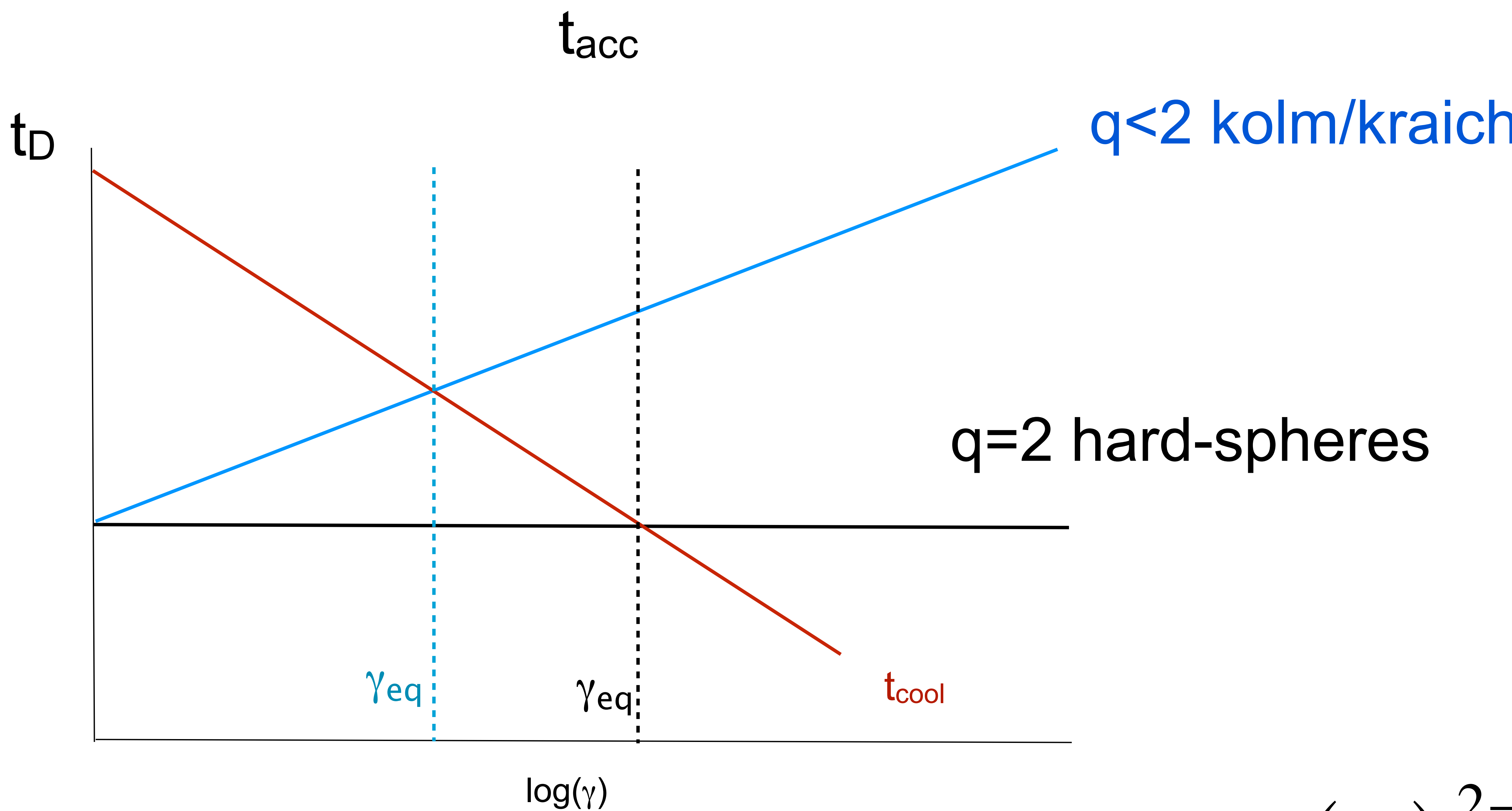


Tramacere +2011





effect of the turbulence index q

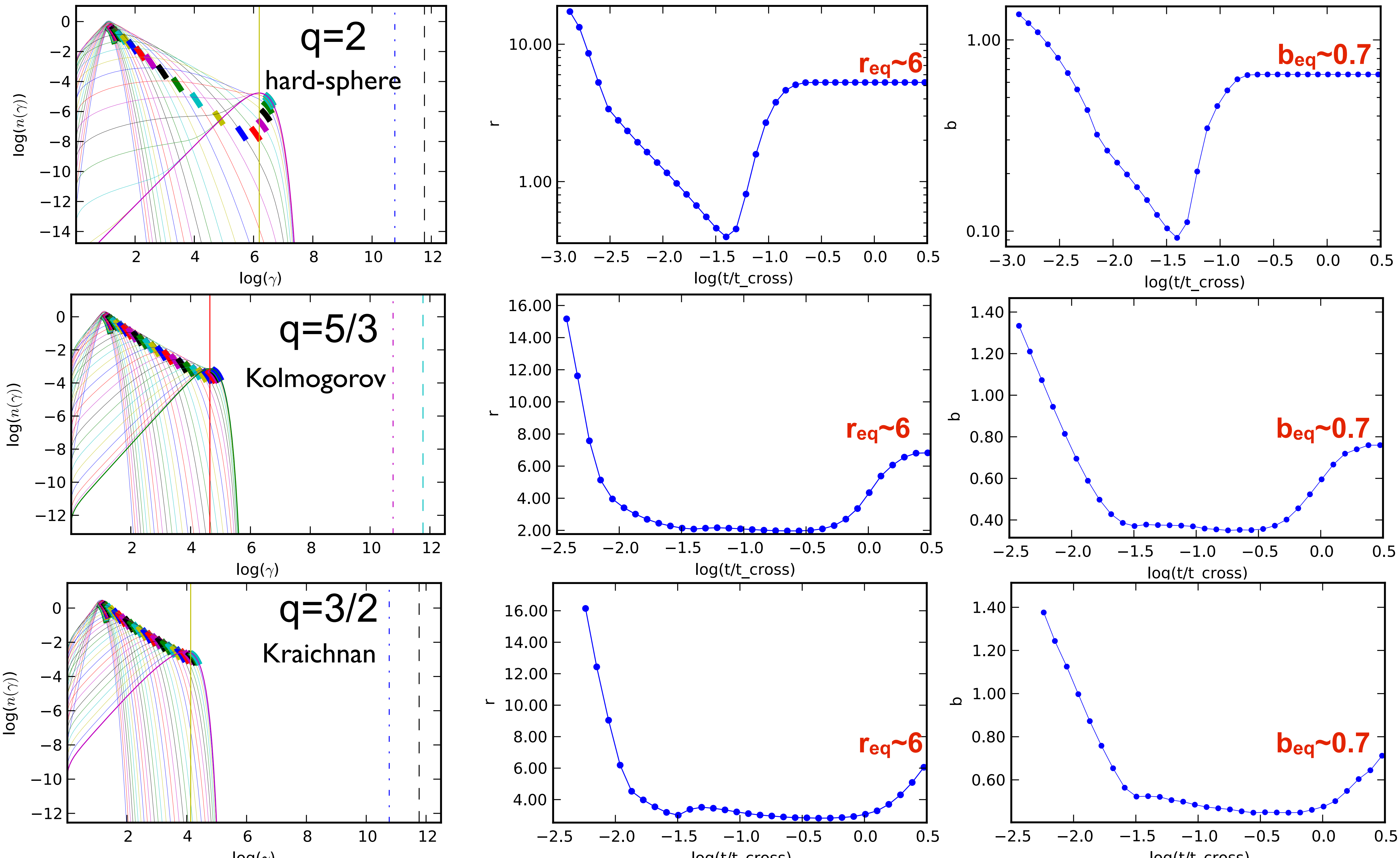


$$t_D = \frac{1}{D_{p0}} \left(\frac{\gamma}{\gamma_0} \right)^{2-q}$$

effect of the turbulence index q

B=1.0 G, $t_{D0}=10^3$, R=5×10^{15} cm

n(γ) curvature | **synch. peak curvature**



log-parabola is not a “new" model...

KARDASHEV 1962

At first, for simplicity, we consider the effect of each process viewed separately on the energy spectrum, and then the simultaneous effect of two or more processes.

Spectra of Isolated Processes

1. Random and Systematic Acceleration.

The kinetic equation is

$$\frac{\partial N}{\partial t} = \alpha_1(t)\frac{\partial}{\partial E}\left(E^2\frac{\partial N}{\partial E}\right) - \alpha_2(t)\frac{\partial}{\partial E}(EN).$$

Let the energy distribution be specified, at each instant of time t_0, by the δ-function in the neighborhood of energy E_0:

$$N(E, 0) = N_0\delta(E - E_0)$$

and

$$\int_0^\infty N(E, 0)\,dE = N_0.$$

Then, utilizing the techniques developed, e.g.,in [13], we may find that

$$N(E, t) = \frac{N_0}{\sqrt{\pi}E2\sqrt{a_1}}e^{-\left(\ln\frac{E_0}{E}+a_1+a_2\right)^2/4a_1}, \quad (1)$$

where

$$a_1 = \int_{t_0}^{t}\alpha_1(t)\,dt, \qquad a_2 = \int_{t_0}^{t}\alpha_2(t)\,dt.$$

increases c
The quanti
to expansi
the quanti
sistently p
creasing E
and conve
correspond
For th
$= KE_0^{-\gamma}$ is
$E_{min} \leq E_0$
initial con

$$\int_{E_{min}}^{E_{max}} K$$

$=$

where

x

At E_{max}

Tramacere+2011

statistical approach

$$n(\gamma) = \frac{N_0}{\gamma \sigma_\gamma \sqrt{(2\pi)}} \exp\left[\frac{-\left(\ln(\gamma/\gamma_0) - n_s\left[\ln \bar{\varepsilon} - \frac{1}{2}(\sigma_\varepsilon/\bar{\varepsilon})^2\right]\right)^2}{2 n_s (\sigma_\varepsilon/\bar{\varepsilon})^2} \right]$$

$$\boxed{\log(n(\gamma)) \propto \frac{(\log \gamma - \mu)^2}{2\sigma_\gamma^2} \propto r \ [\log(\gamma) - \mu]^2}$$

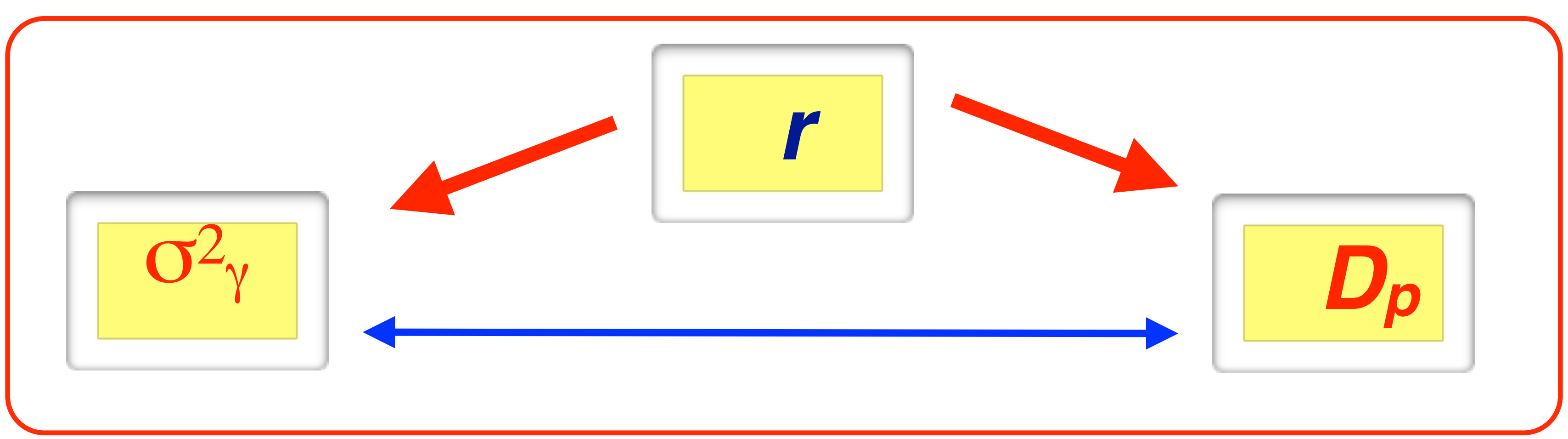
diffusion equation approach

$$n(\gamma, t) = \frac{N_0}{\gamma \sqrt{4\pi D_{p0} t}} \exp\left\{ - \frac{[\ln(\gamma/\gamma_0) - (A_{p0} - D_{p0})t]^2}{4 D_{p0} t} \right\}$$

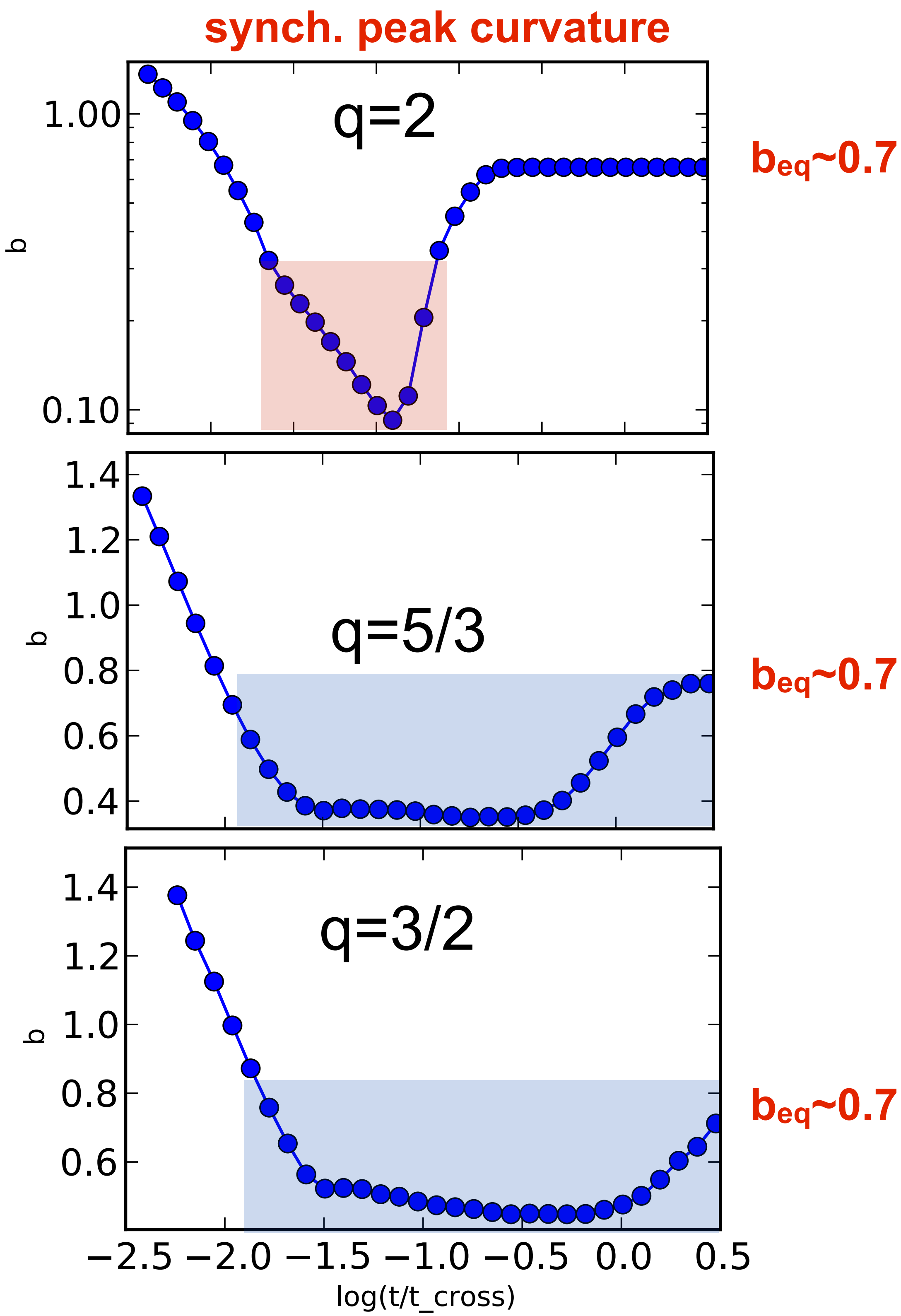
$$r \propto \frac{1}{D_{p0} t} \rightarrow D_{p0} \propto \left(\frac{\sigma_\varepsilon}{\bar{\varepsilon}}\right)^2$$

The curvature *r* is inversely proportional to ***t*=>*n*$_s$** and ***D*$_p$=>**σ_ε

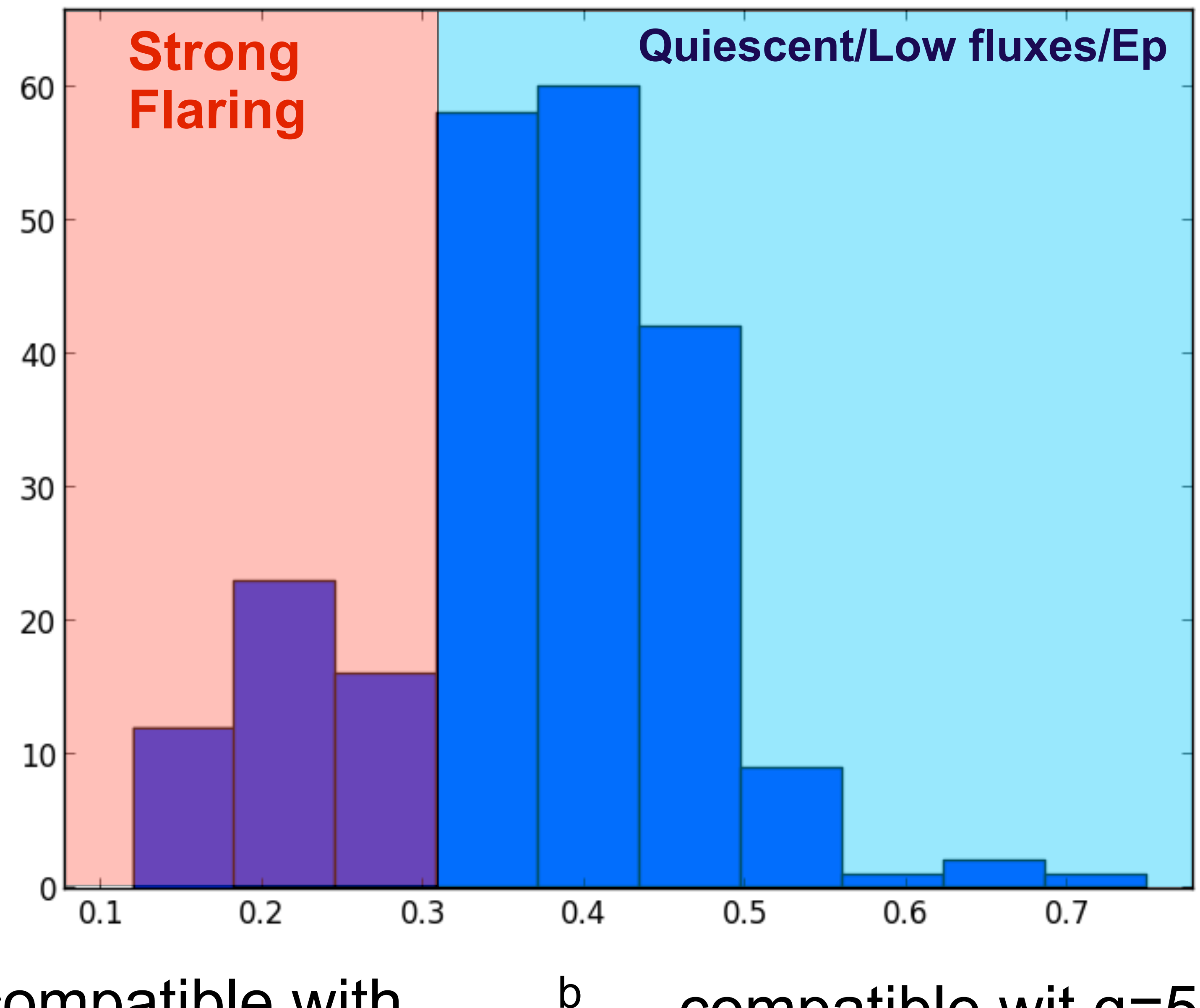
log-parabolic shape natural consequence of dispersion



b distributions and q



both flaring and quiescent seem to be far from equilibrium b eq.~[0.7-1.0] (if full KN or S)



compatible with q=2 far from equilibrium constraint on B

compatible wit q=5/3 constraint on B, and duration, or TH/KN

q=2 require more fine tuning,especially on duration

self-consistent approach: acc+cooling

$$t_D = \frac{1}{D_{p0}} \left(\frac{\gamma}{\gamma_0} \right)^{2-q}$$

$$t_{DA} = \frac{1}{2D_{p0}} \left(\frac{\gamma}{\gamma_0} \right)^{2-q}$$

observed values

$E_{p1}/E_{p2} \sim 5$

$\Delta t \sim$ few ks

→

values compatible with Tammi & Duffy 2009

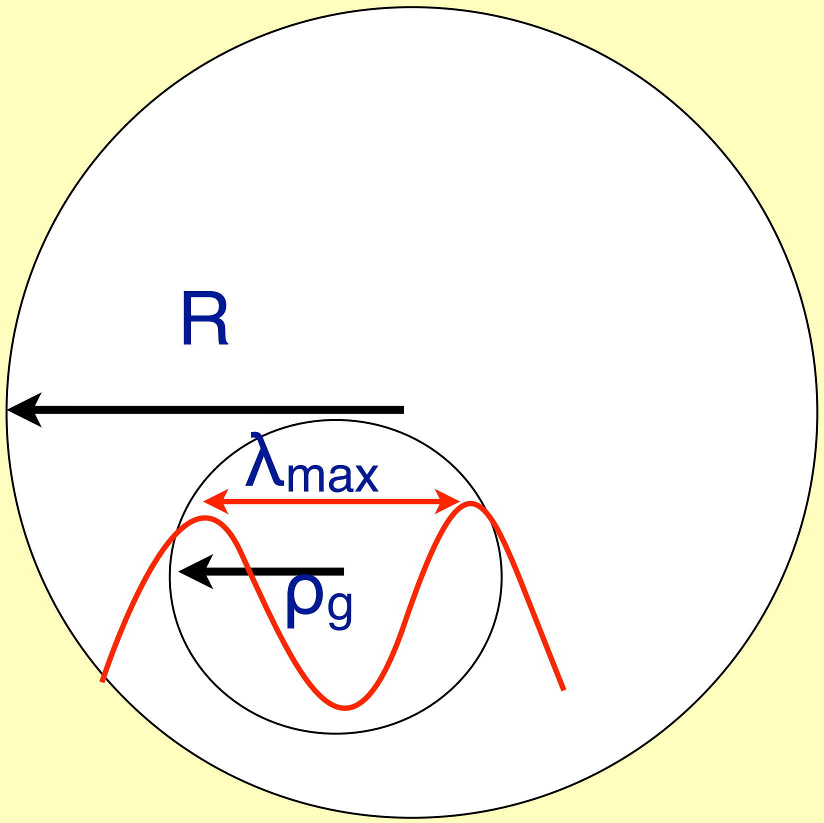
$t_{DA} \sim< 5$ ks

$t_D \sim< 10$ ks

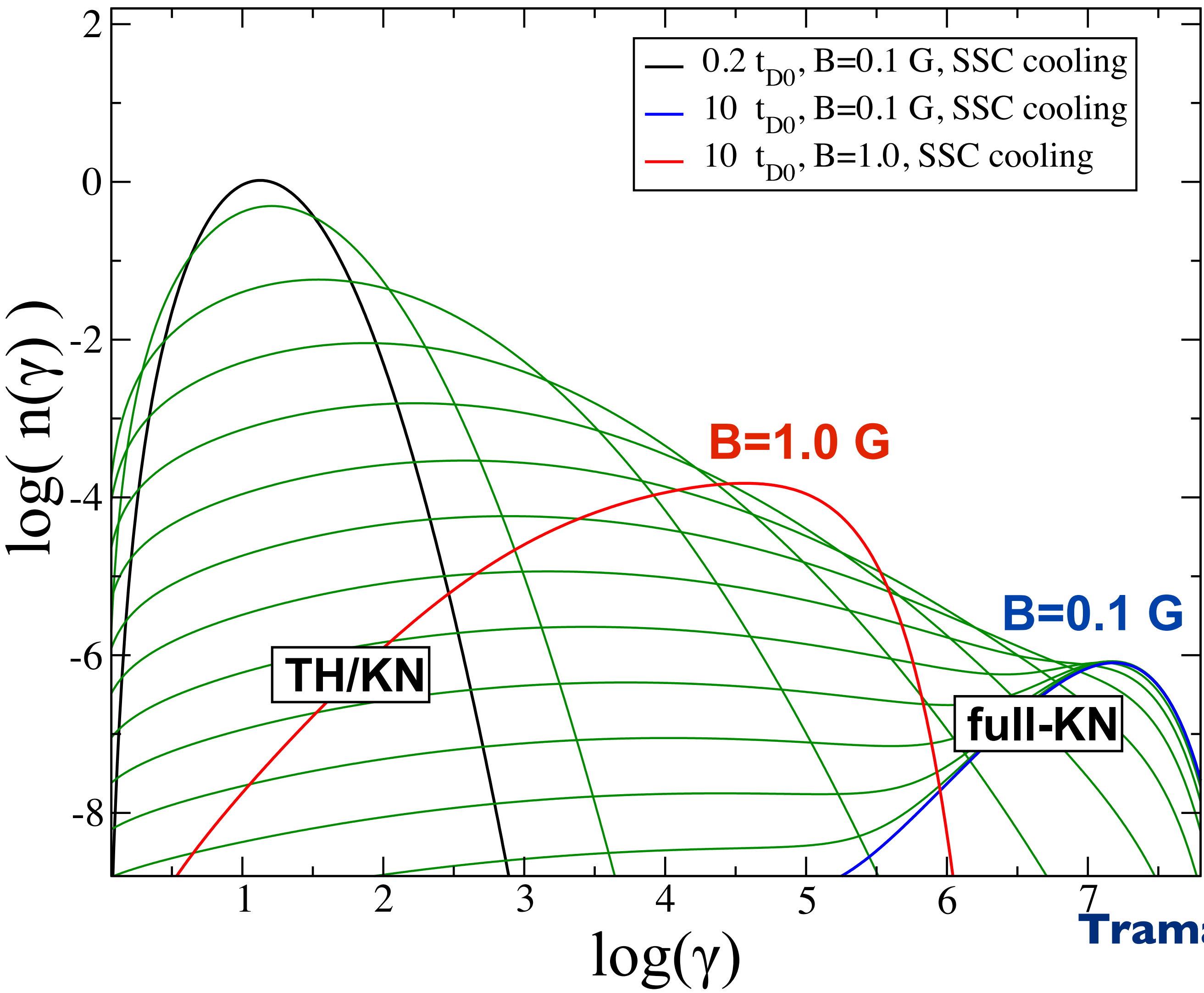
set-up of the accelerator

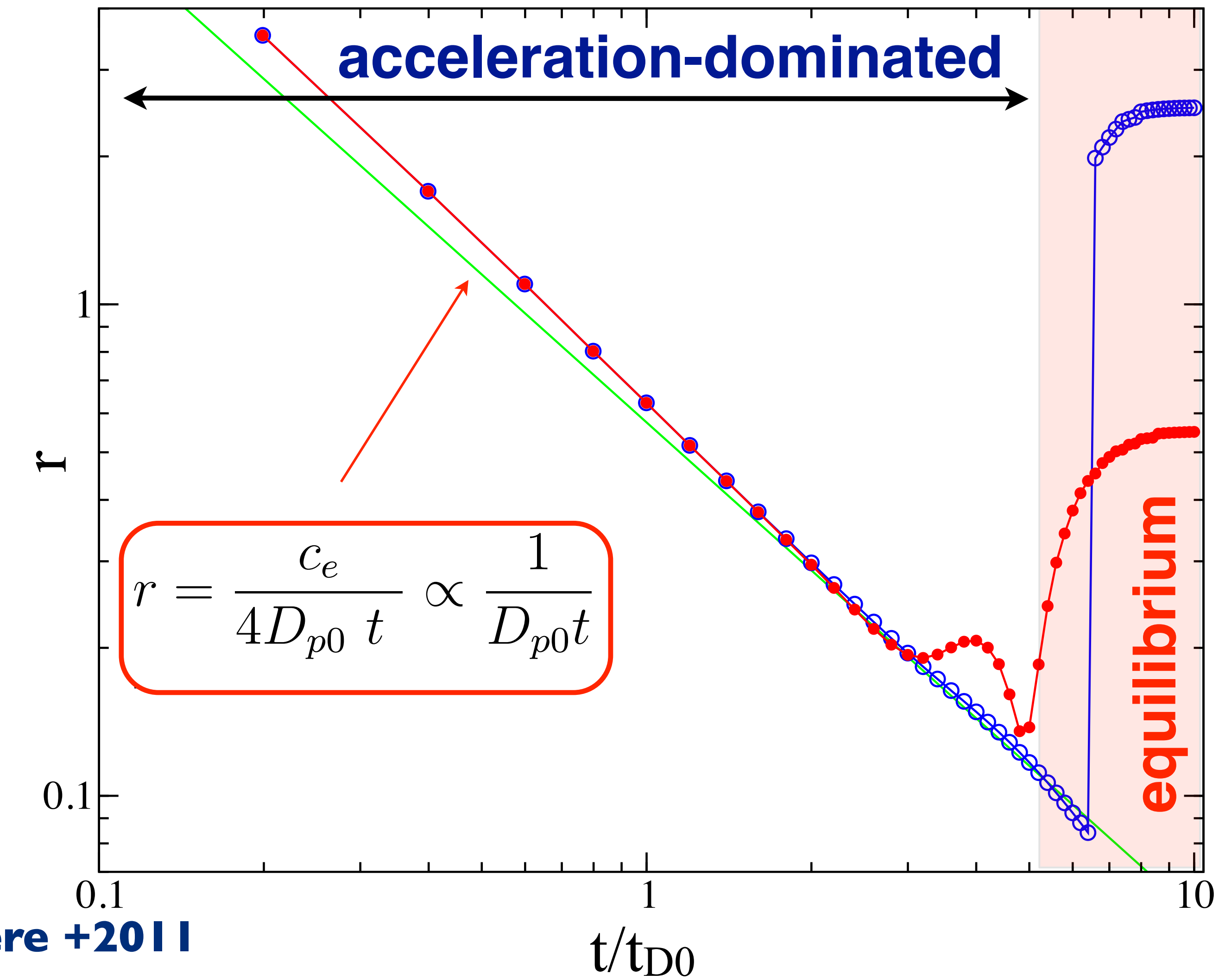
- $R \sim 10^{13}$-10^{15} cm
- $\delta B/B << 1$, $B \sim [0.01$-$1.0]$ G
- $\beta_A \sim 0.1$-0.5
- $\lambda_{max} < R$ => $\sim 10^{[9-15]}$ cm
- $\rho_g < \lambda_{max}$ => $\gamma_{max} \sim 10^{7.5}$

→ **$t_D \sim< 10^4$ ks**



Flare: acc.-dominated-vs-equil.,R= 10^{15} cm, q=2

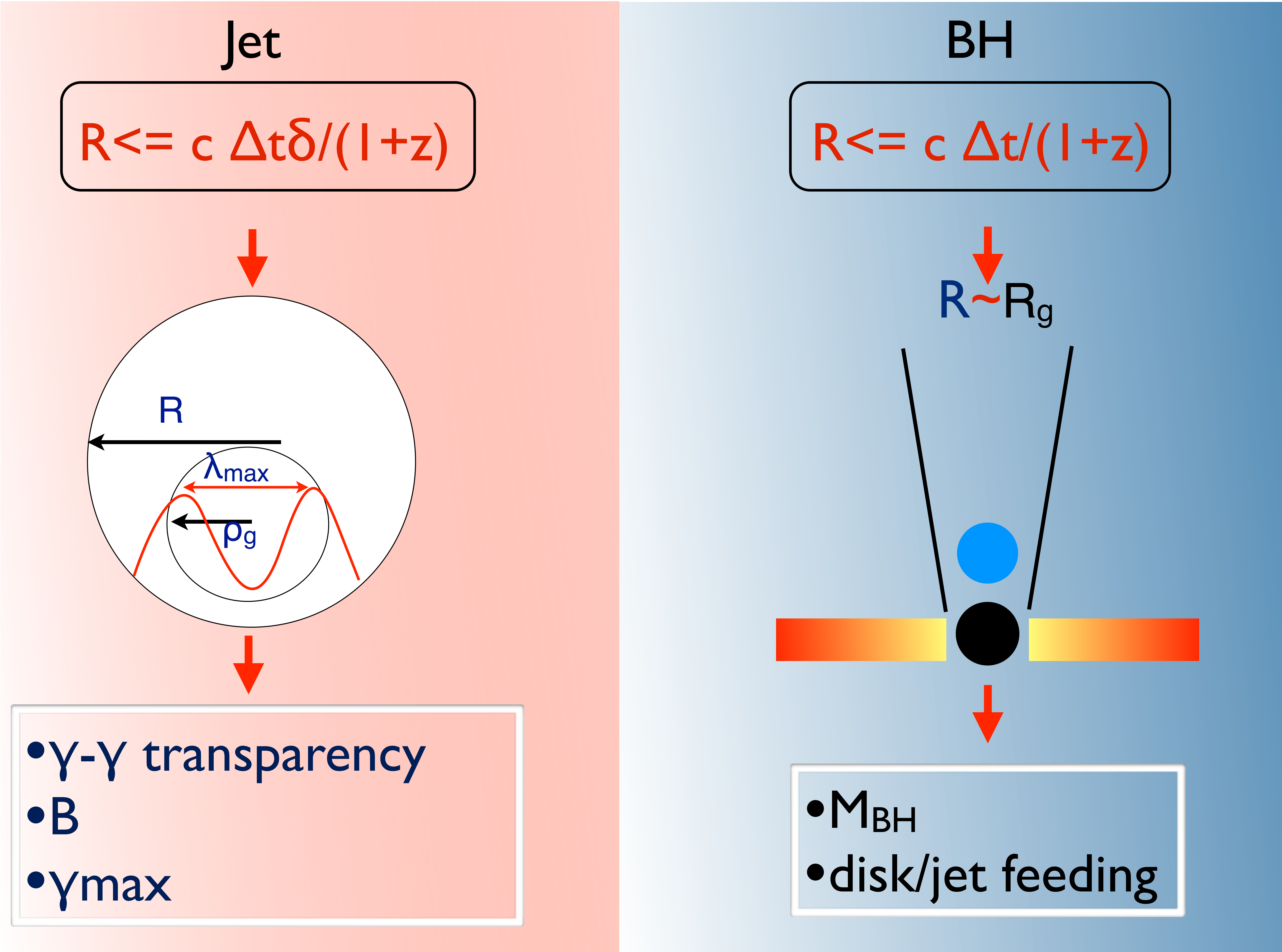


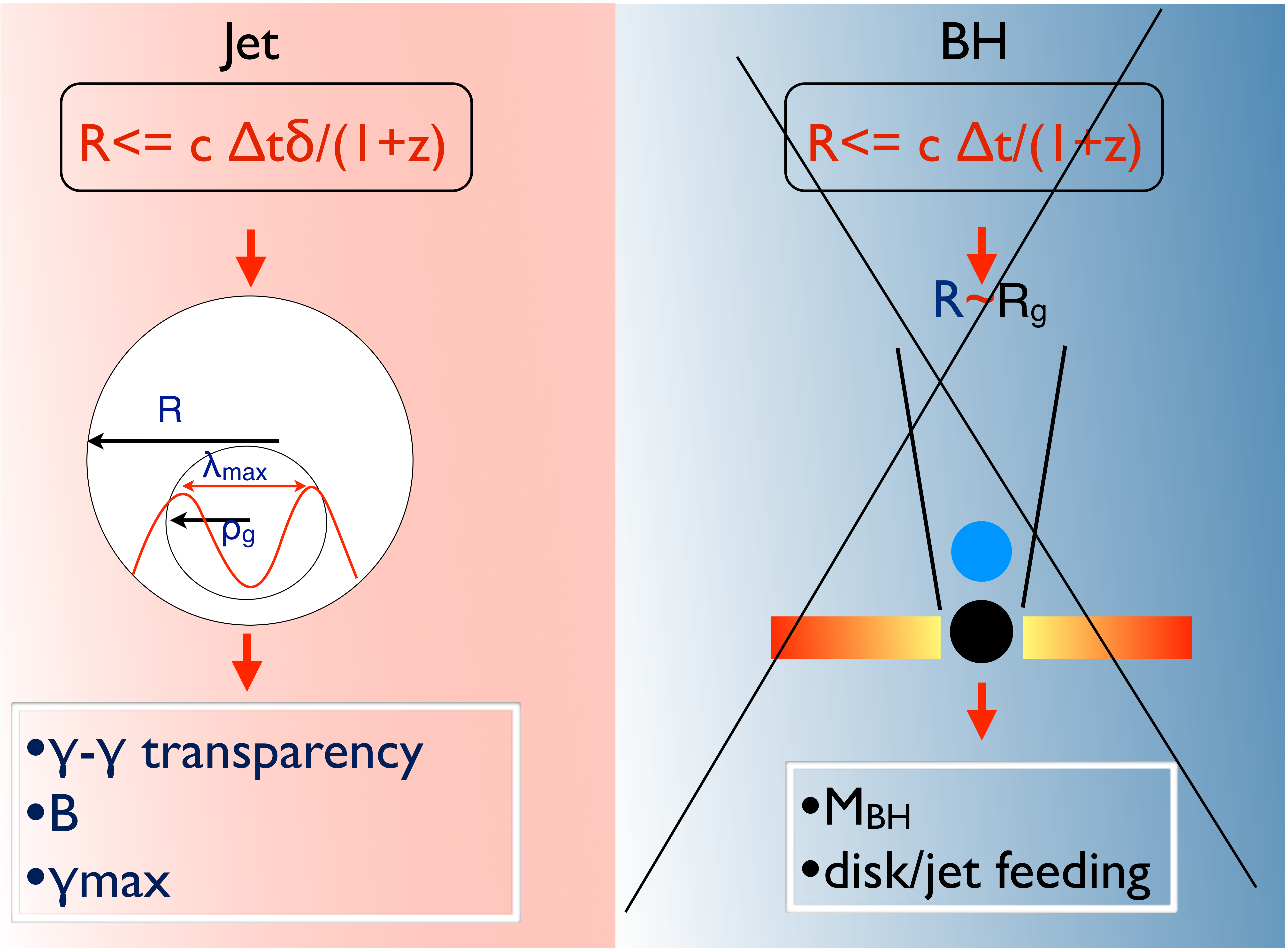


Tramacere +2011

- mono energetic inj., $t_{inj} << t_{acc}$, $t_{inj} << t_{sim}$
- we measure r@peak as a function of the time
- two phase: **acceleration-dominated**, **equilibrium**
- equil. distribution:
    - **f=1 for q=2 and S, full TH, or full KN**
- **equil. curv.: r~2.5, (r_{3p}~6.0) for TH or full KN**
- **equil. curv.: r~0.6, (r_{3p}~4.0) for TH-KN**

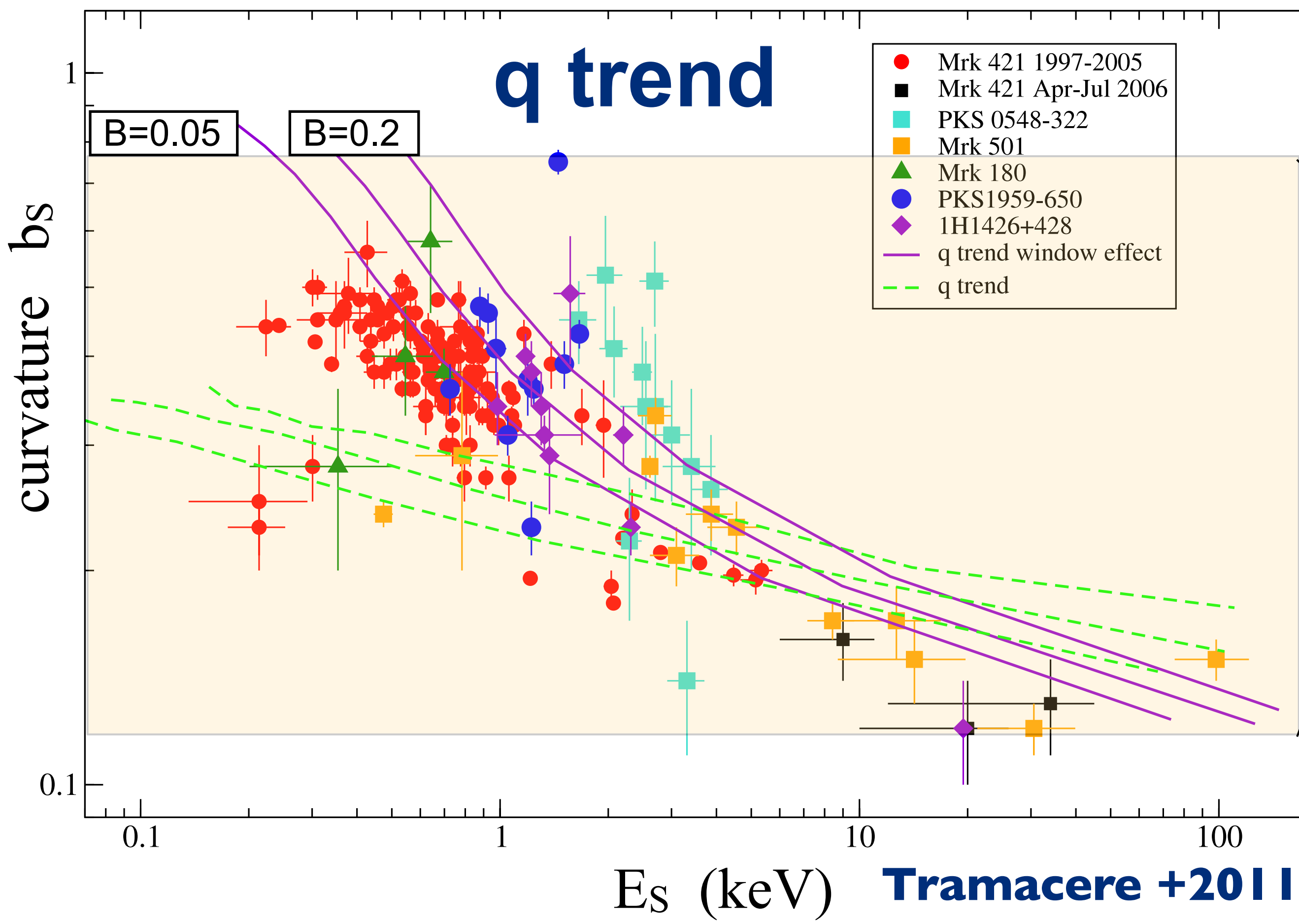
$$n(\gamma) \propto \gamma^2 \exp\left[\frac{-1}{f(q,\dot{\gamma})}\left(\frac{\gamma}{\gamma_{eq}}\right)^{f(q,\dot{\gamma})}\right]$$

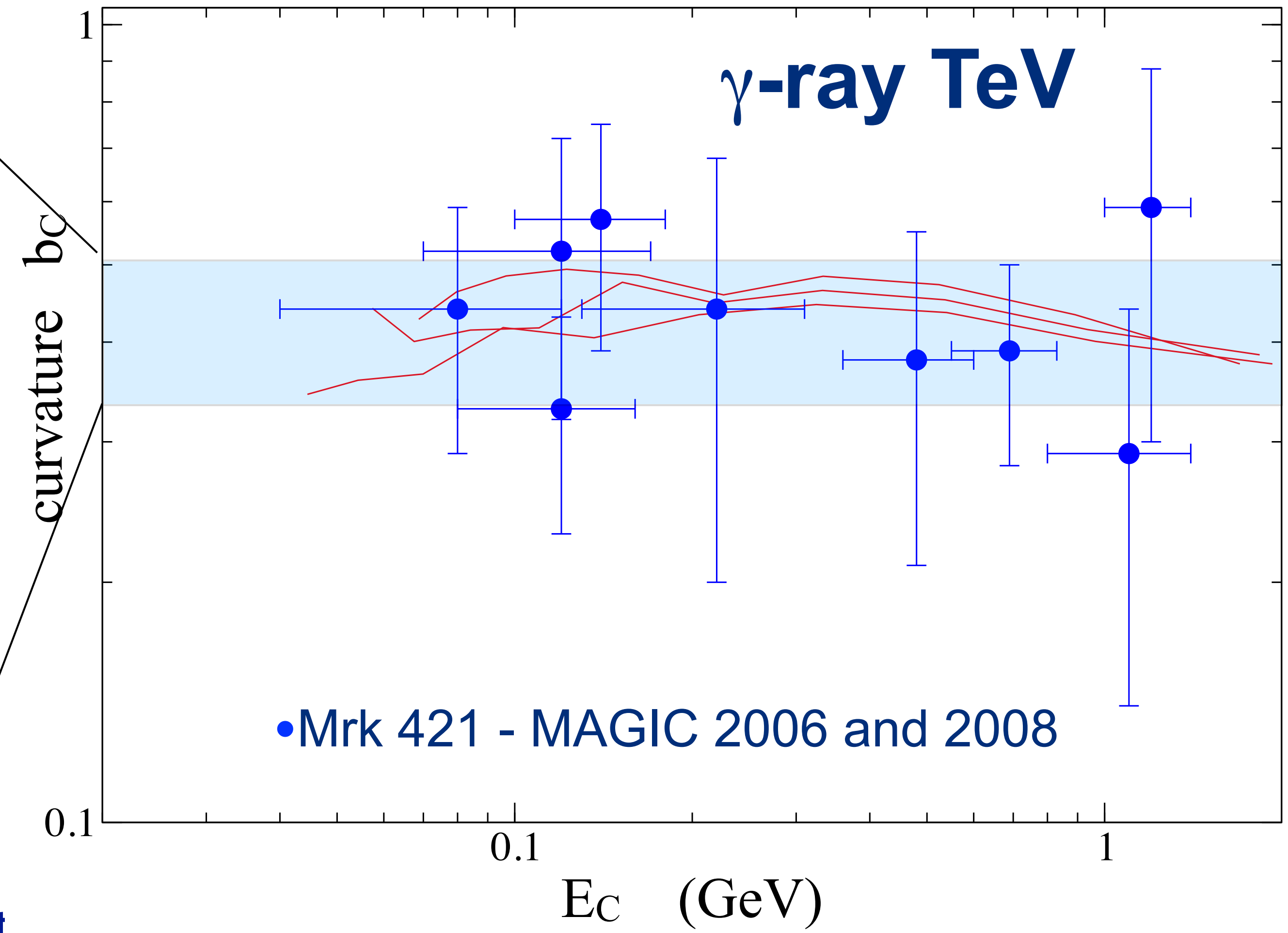

Jet
R<= c ΔtΔδ/(1+z)
R
λmax
ρg
•γ-γ transparency
•B
•γmax
BH
R<= c Δt/(1+z)
R~Rg
•MBH
•disk/jet feeding



Jet
R<= c ΔtΔδ/(1+z)
R
λmax
ρg
•γ-γ transparency
•B
•γmax
BH
R<= c Δt/(1+z)
R~Rg
•MBH
•disk/jet feeding

E_s-b_s X-ray trend and γ-ray predictions

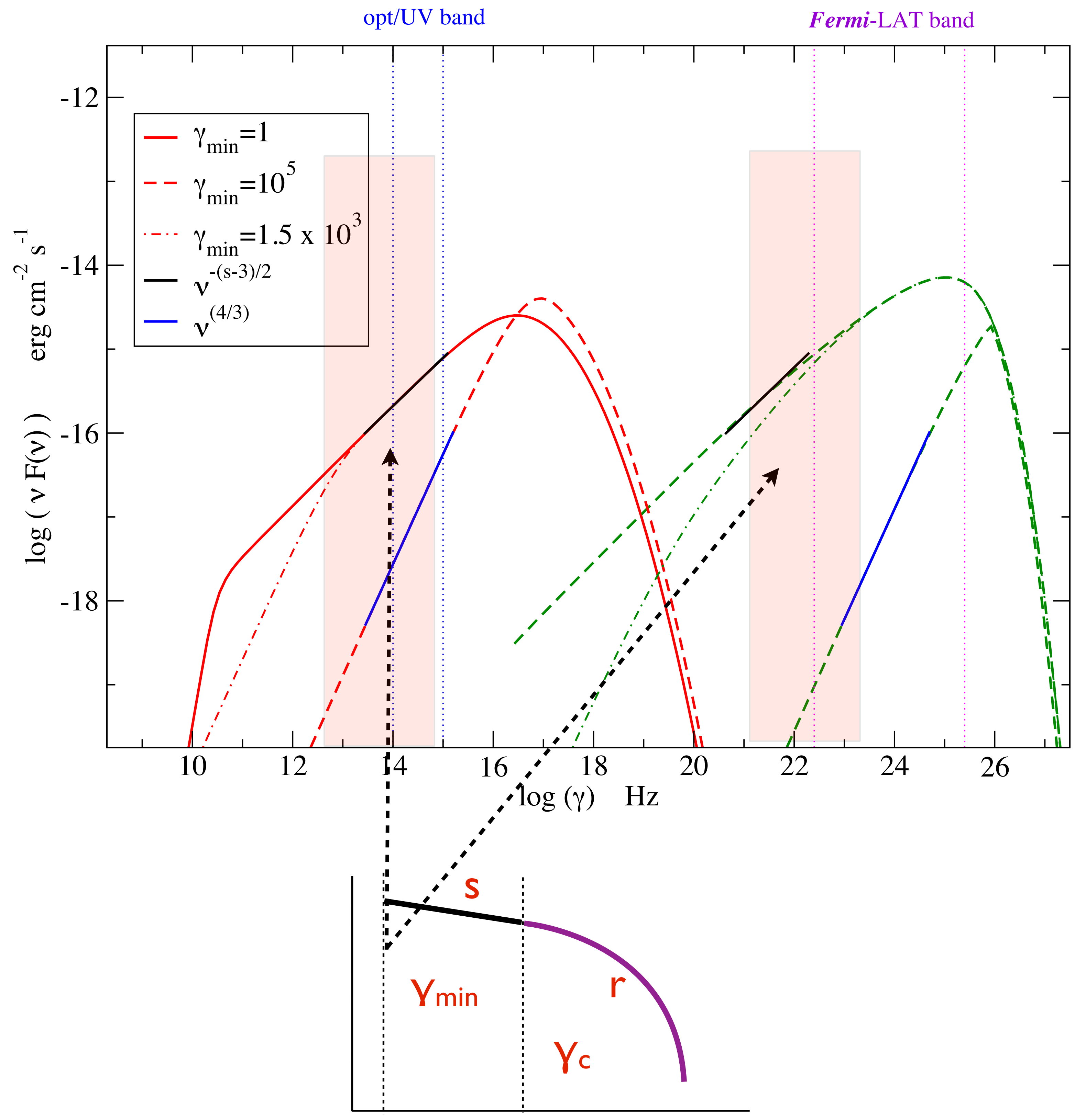




- data span **13 years**, both flaring and quiescent states
- We are able to reproduce these long-term behaviours, by changing the value of only one parameter (q)
- curvature values imply distribution far from the equilibrium (b~[0.7–1.0])
- More data needed at GeV/TeV, curvature seems to be cooling-dominated

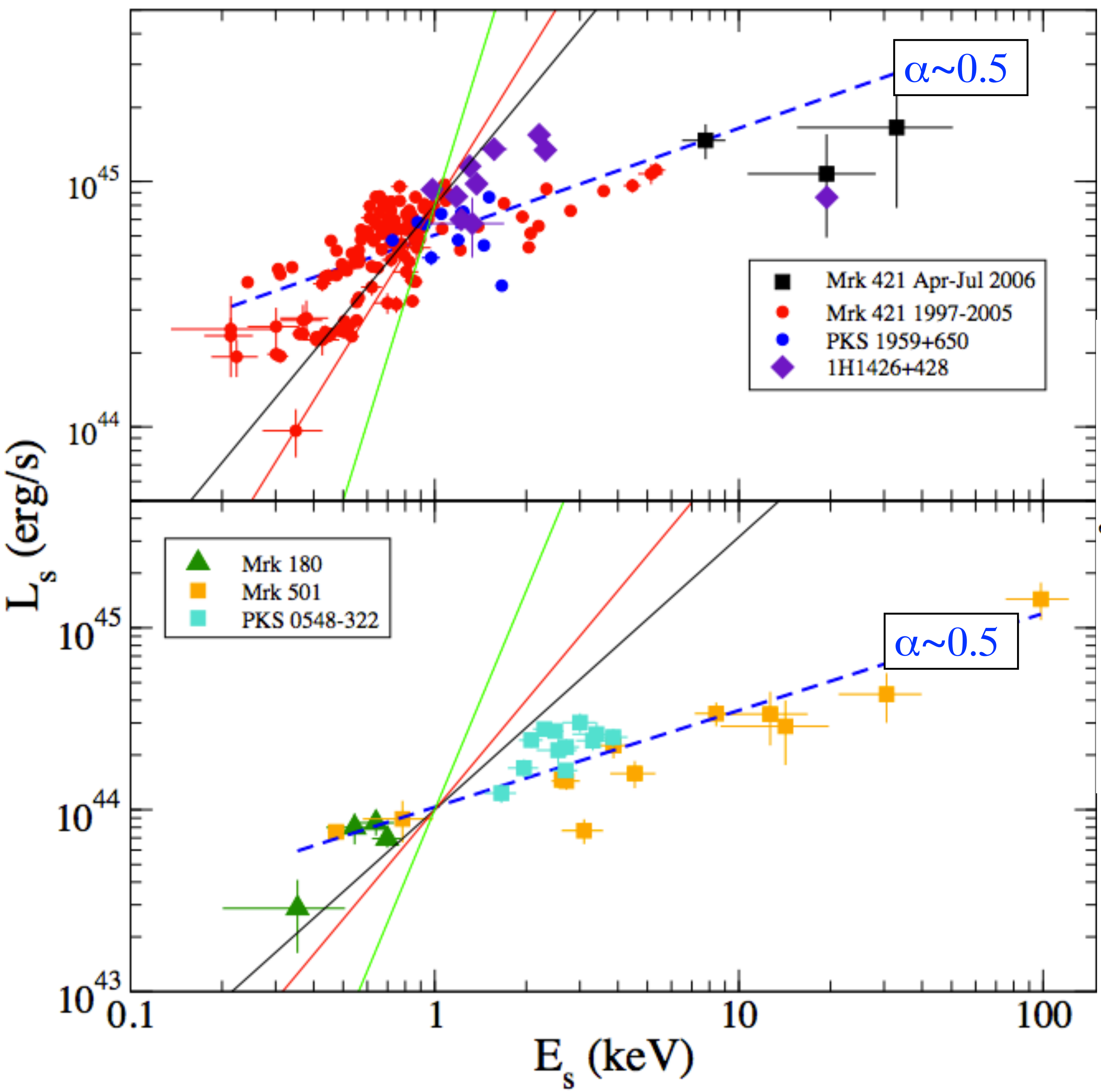
$L_{\rm inj}$ (E_s–b_s trend)	(erg s^{-1})	5×10^{39}
$L_{\rm inj}$ (E_s–L_s trend)	(erg s^{-1})	5×10^{38}, 5×10^{39}
q		[3/2, 2]
t_A	(s)	1.2×10^3
$t_{D_0} = 1/D_{P0}$	(s)	$[1.5 \times 10^4, 1.5 \times 10^5]$
$T_{\rm inj}$	(s)	10^4
$T_{\rm esc}$	(R/c)	2.0

HBLs case


opt/UV band
Fermi-LAT band
-12
-14
-16
-18
10
12
14
16
18
20
22
24
26
log (ν F(ν)) erg cm^{-2} s^{-1}
log (γ) Hz
γ_{min}=1
γ_{min}=10^5
γ_{min}=1.5 x 10^3
$\nu^{-(s-3)/2}$
$\nu^{(4/3)}$
s
r
γ_{min}
γ_c

acceleration signature in the E_S-vs-L_s trend

long-trend main drivers

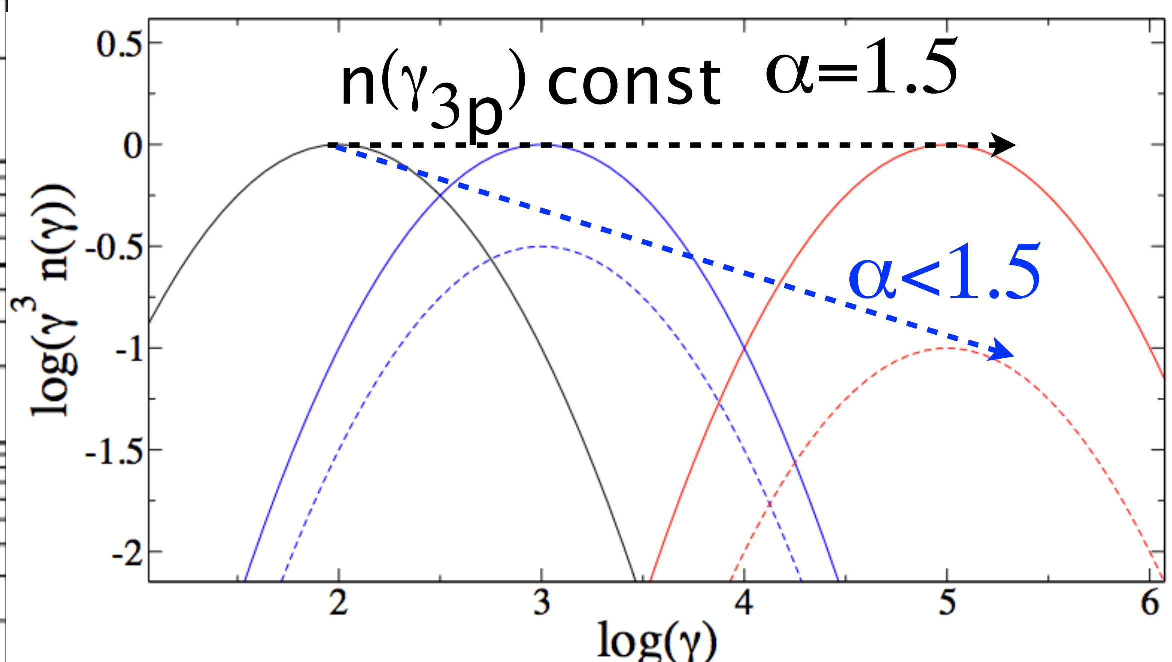


Tramacere+2009

$$S_s(E_s) \propto n(\gamma_{3p})\gamma_{3p}^3 B^2 \delta^4$$
$$E_s \propto \gamma_{3p}^2 B\delta.$$

↓

$$S_s \propto (E_s)^\alpha$$



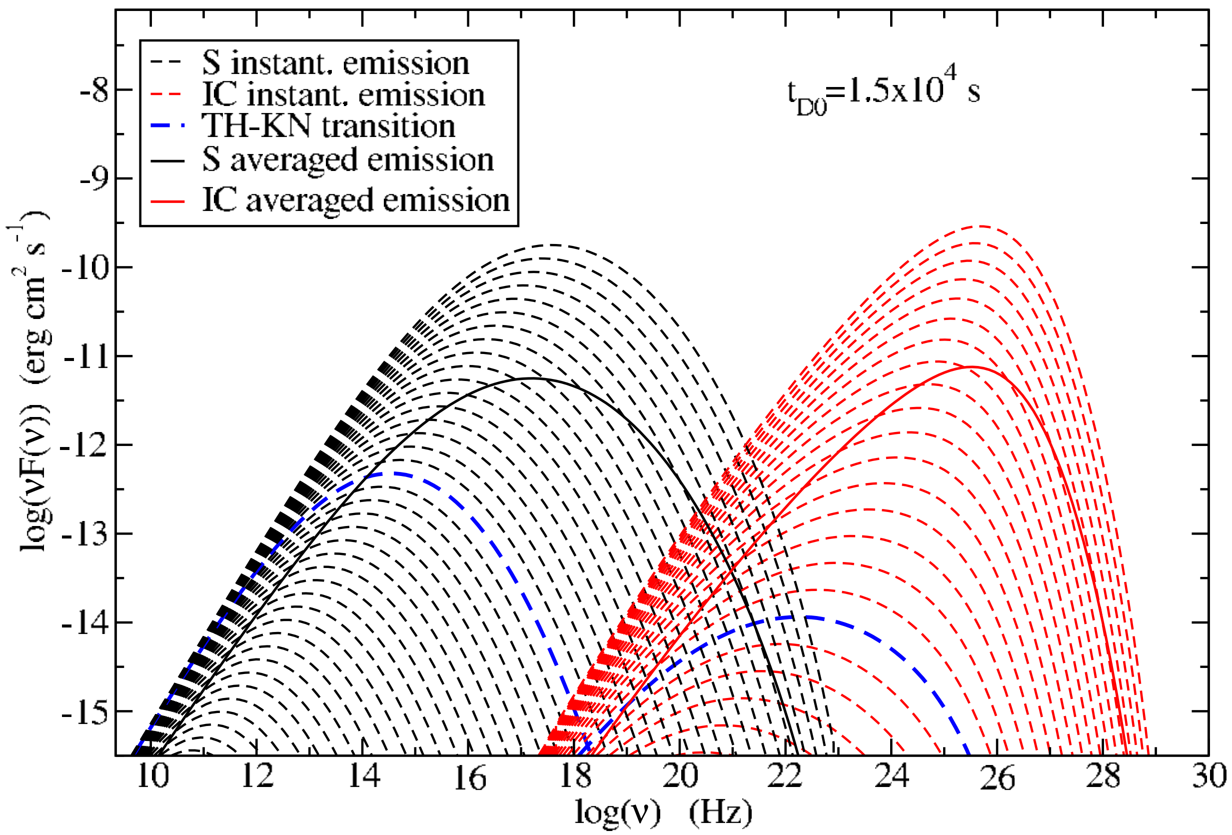
- γ_{3p}↑ **and $n(\gamma 3p)$↓ =>** α<1.5

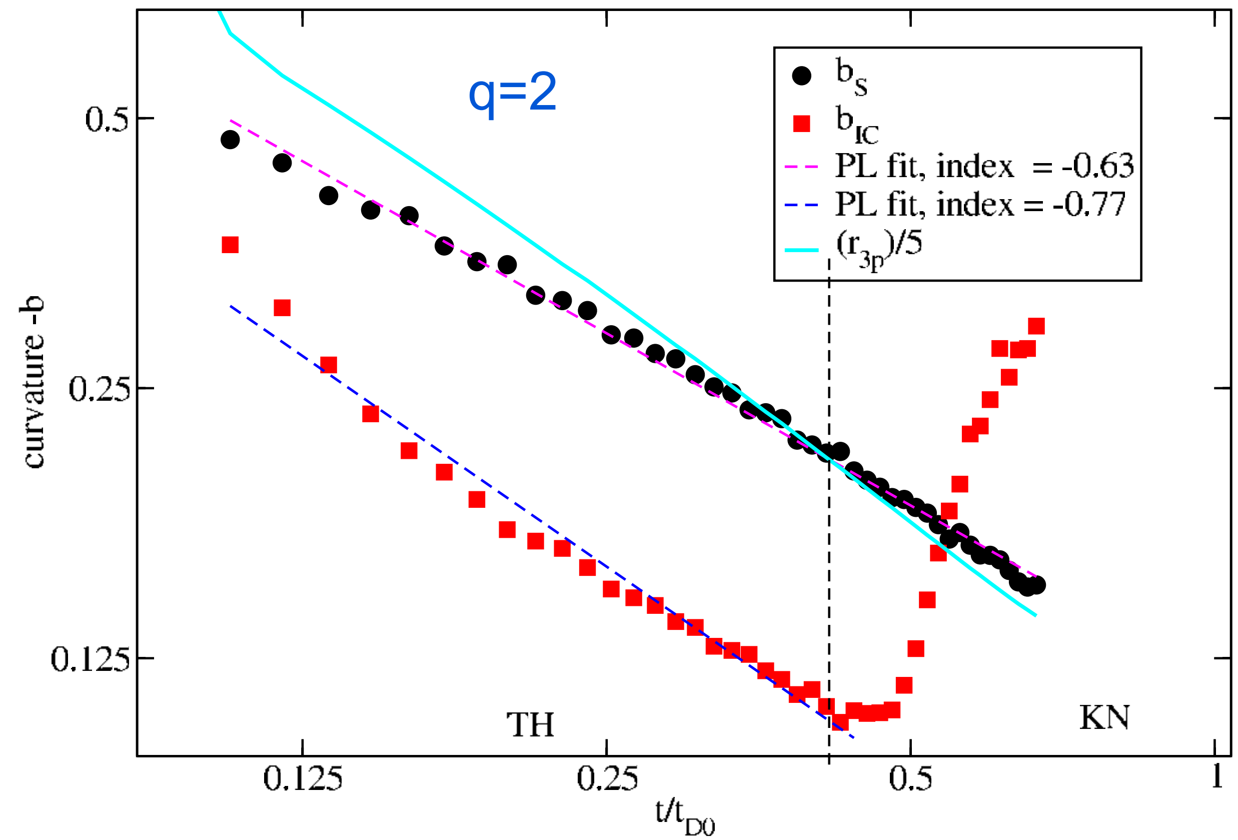
acceleration+energy conservation

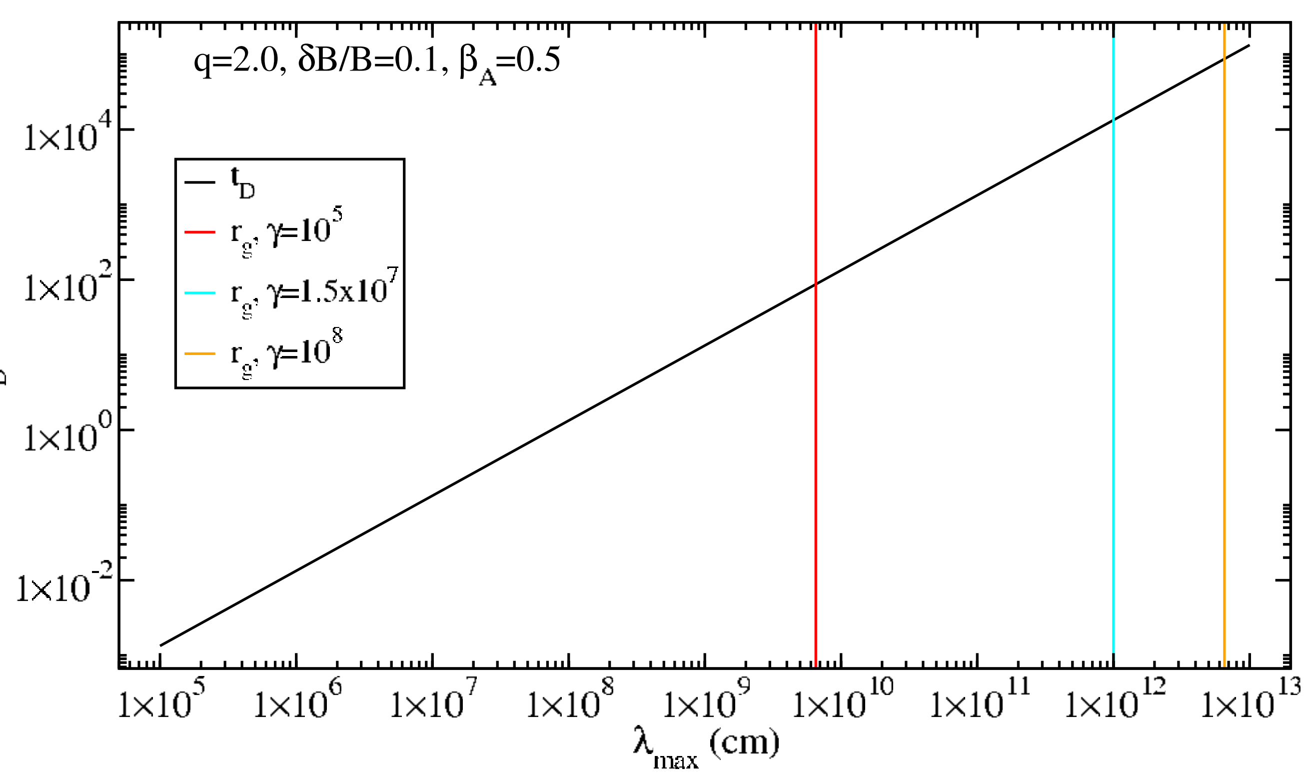
- B->α=2.0,
- δ->α=4

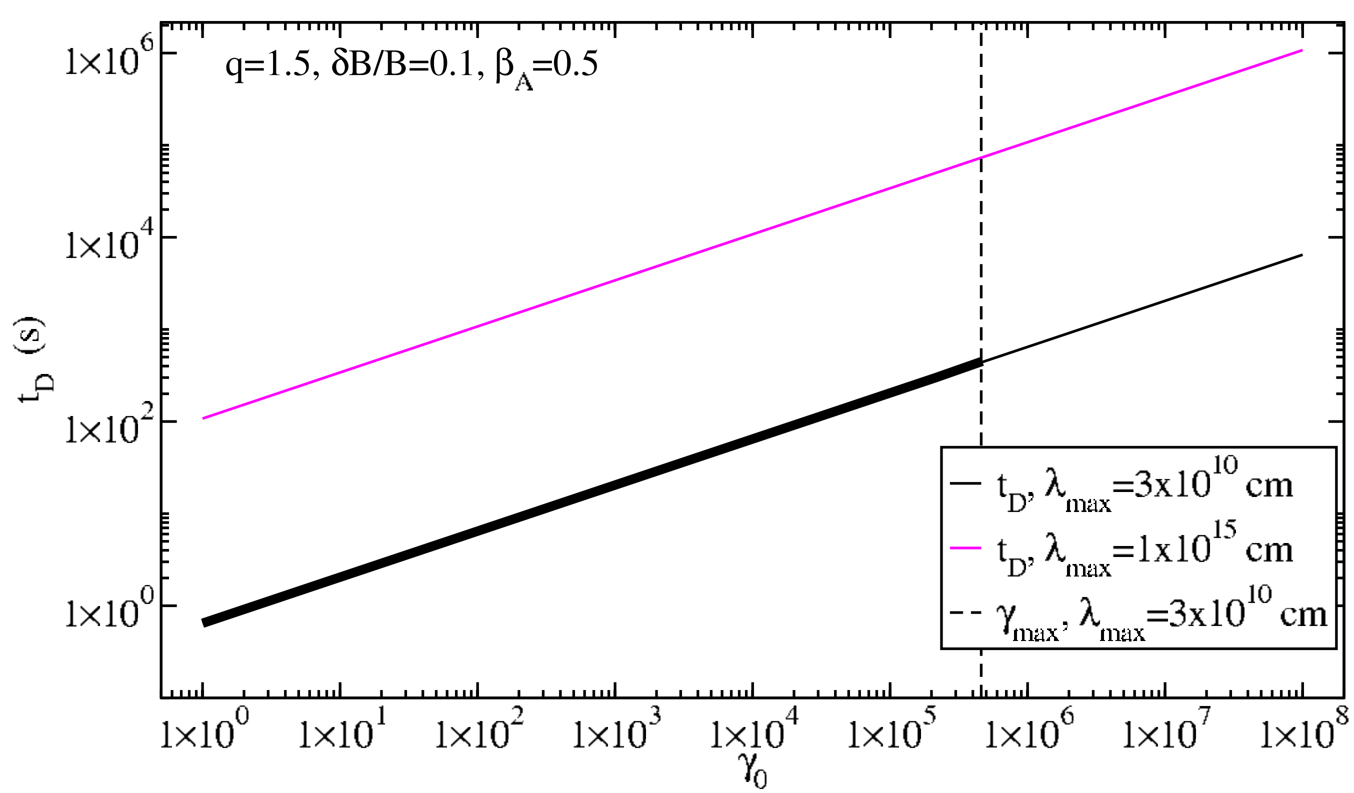
incompatible as long-trend main driver

SEDs evolution

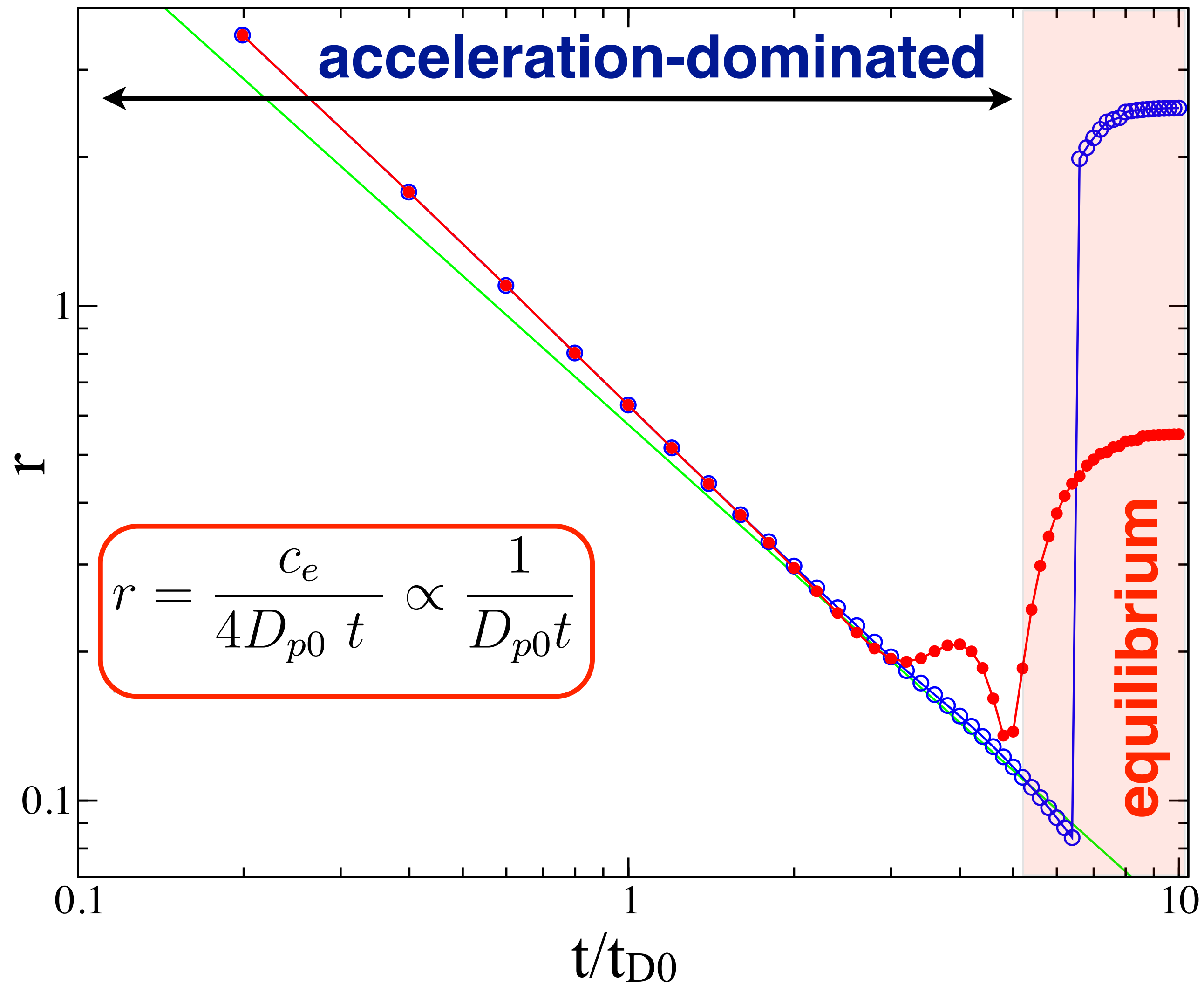


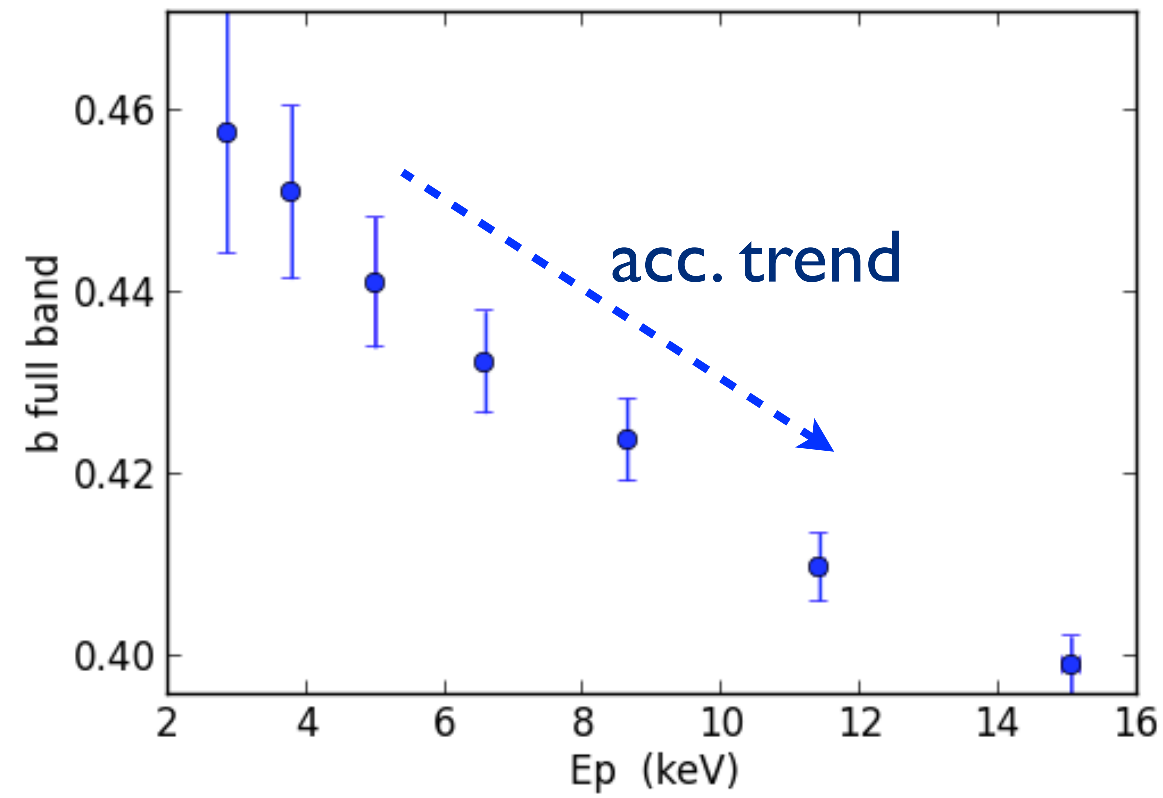


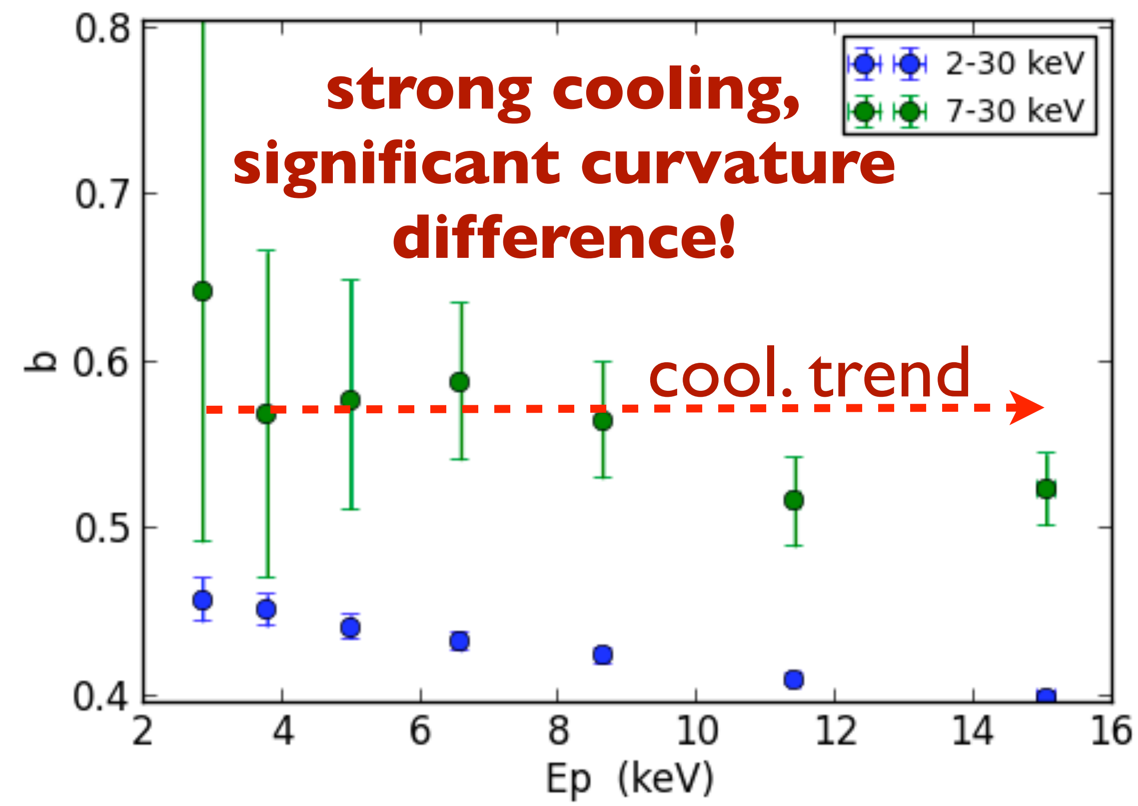




Strong cooling



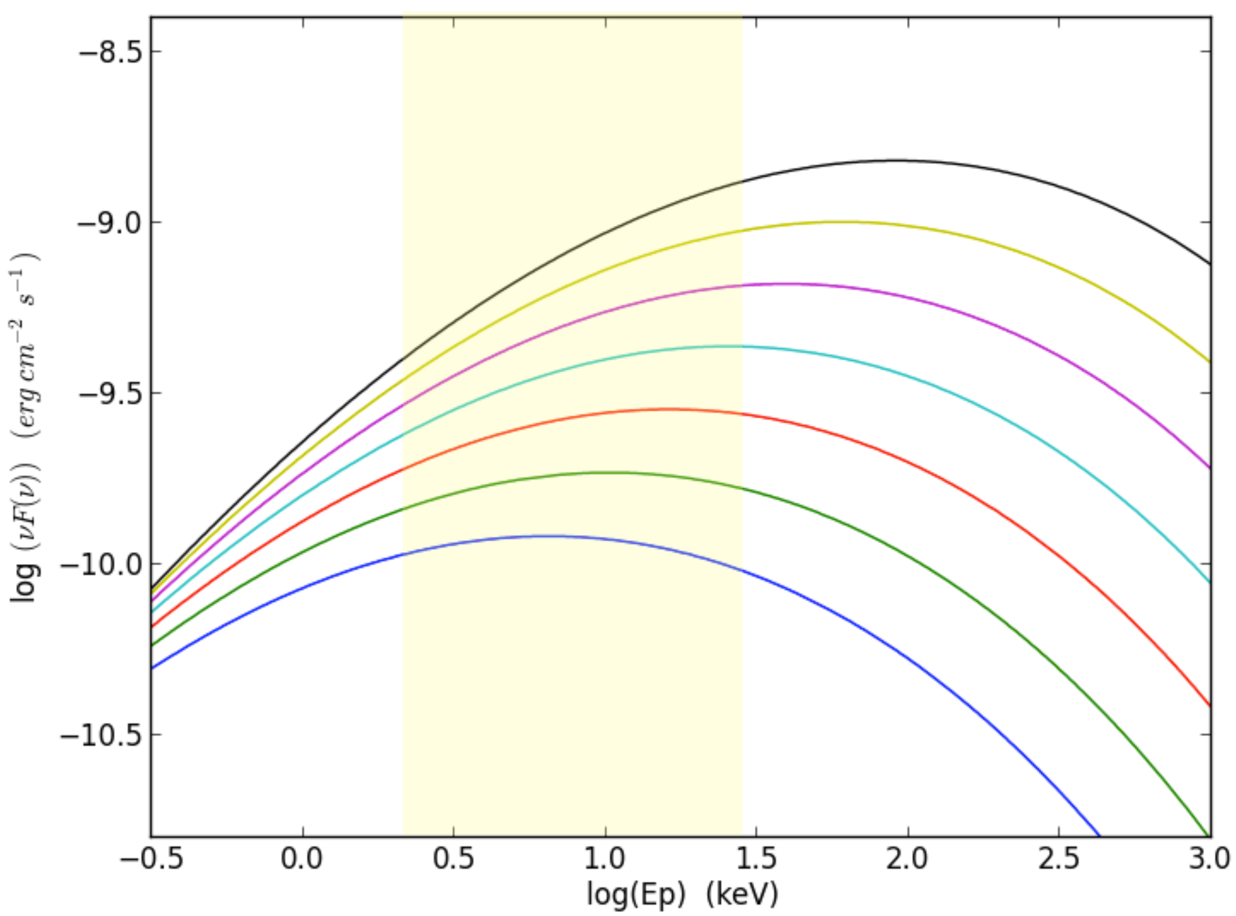




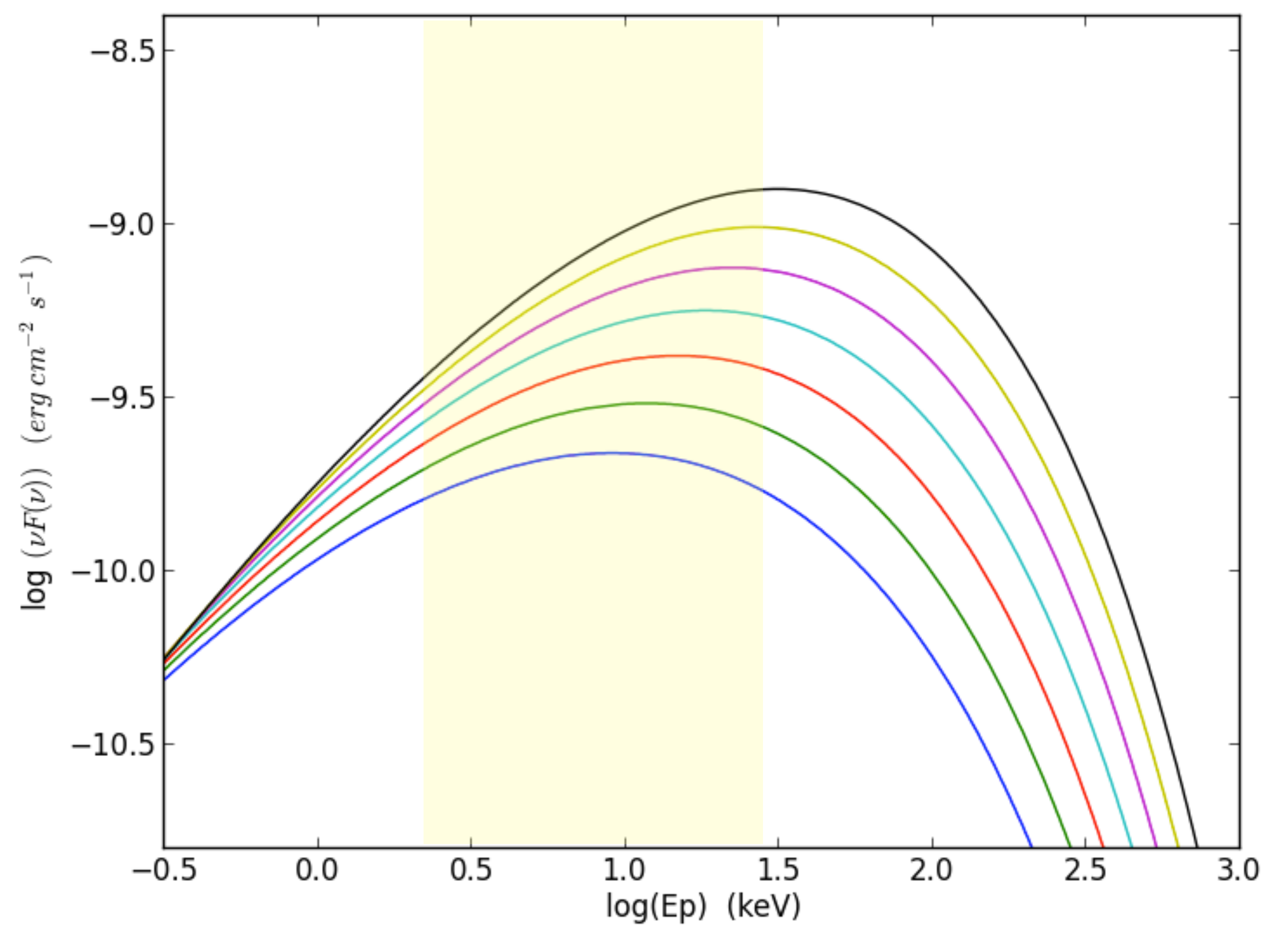
- Full bands curvature related to EED broadness, acceleration signature
- High energy band, dominated by cooling, moving towards the equilibrium

Moving Ep above 30 keV

Low cooling



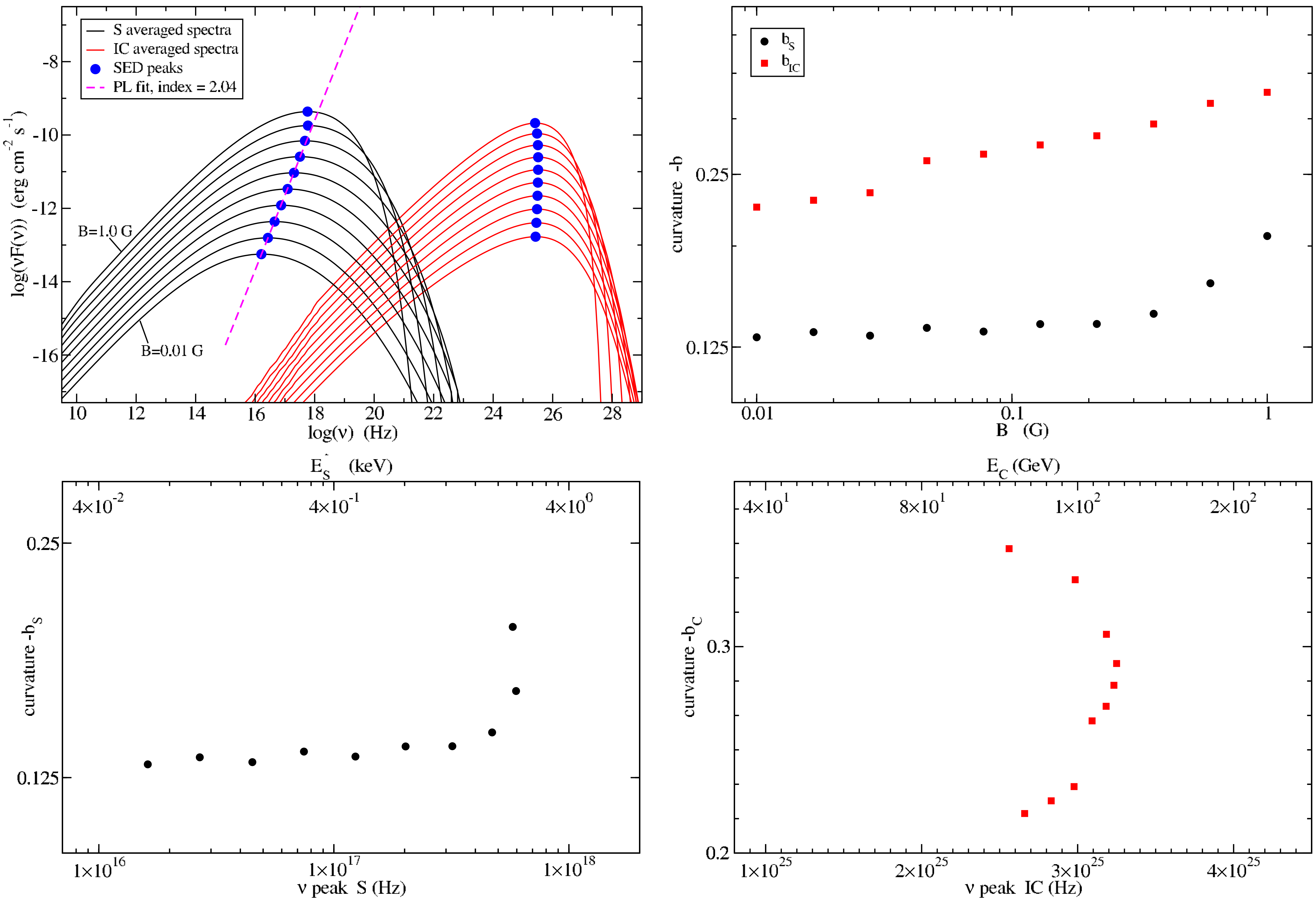
Strong cooling



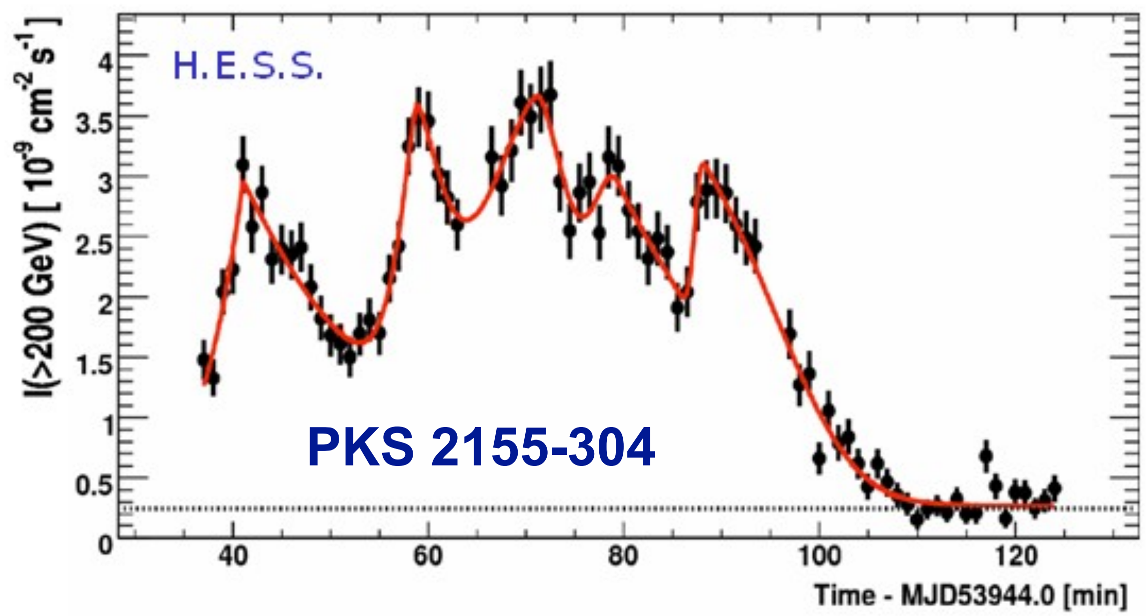
B	0.2/1.0	*G*
R	3×10^{15}	*cm*
$\mathbf{L_{inj}}$	5×10^{39}	*erg/s*
q	2	
$\mathbf{t_A}$	1.2×10^{3}	*s*
$\mathbf{t_D}$	$2.2\text{-}\times10^{4}$	*s*

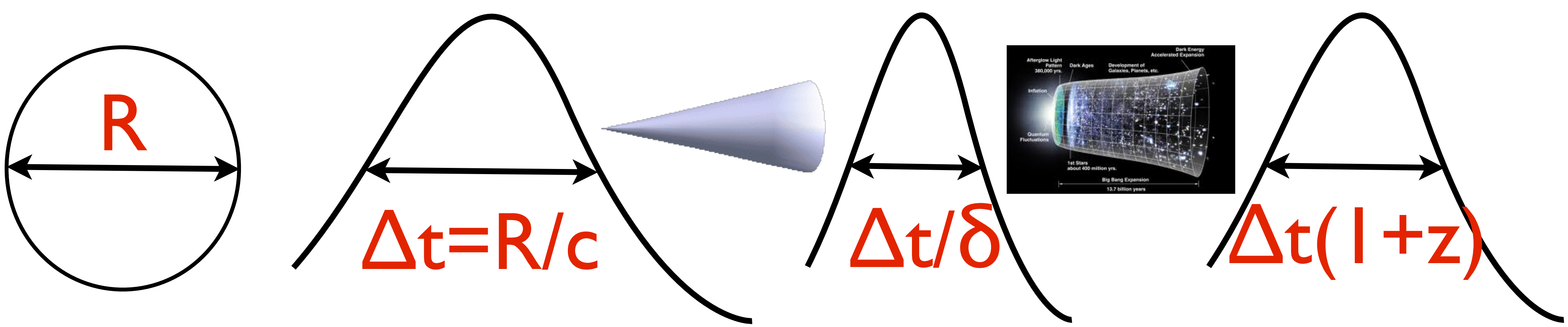
- SEDs are rescaled in order that the **brightest** state matches the flux of 10^{-09} erg cm^{-2} s^{-1} [2-10] keV
- during the flares, the fluxes range in ~ 1×10^{-10}-10^{-9} erg cm^{-2} s^{-1}
- **1 ks integration time**

Effect o B on SEDs



Rapid Variability

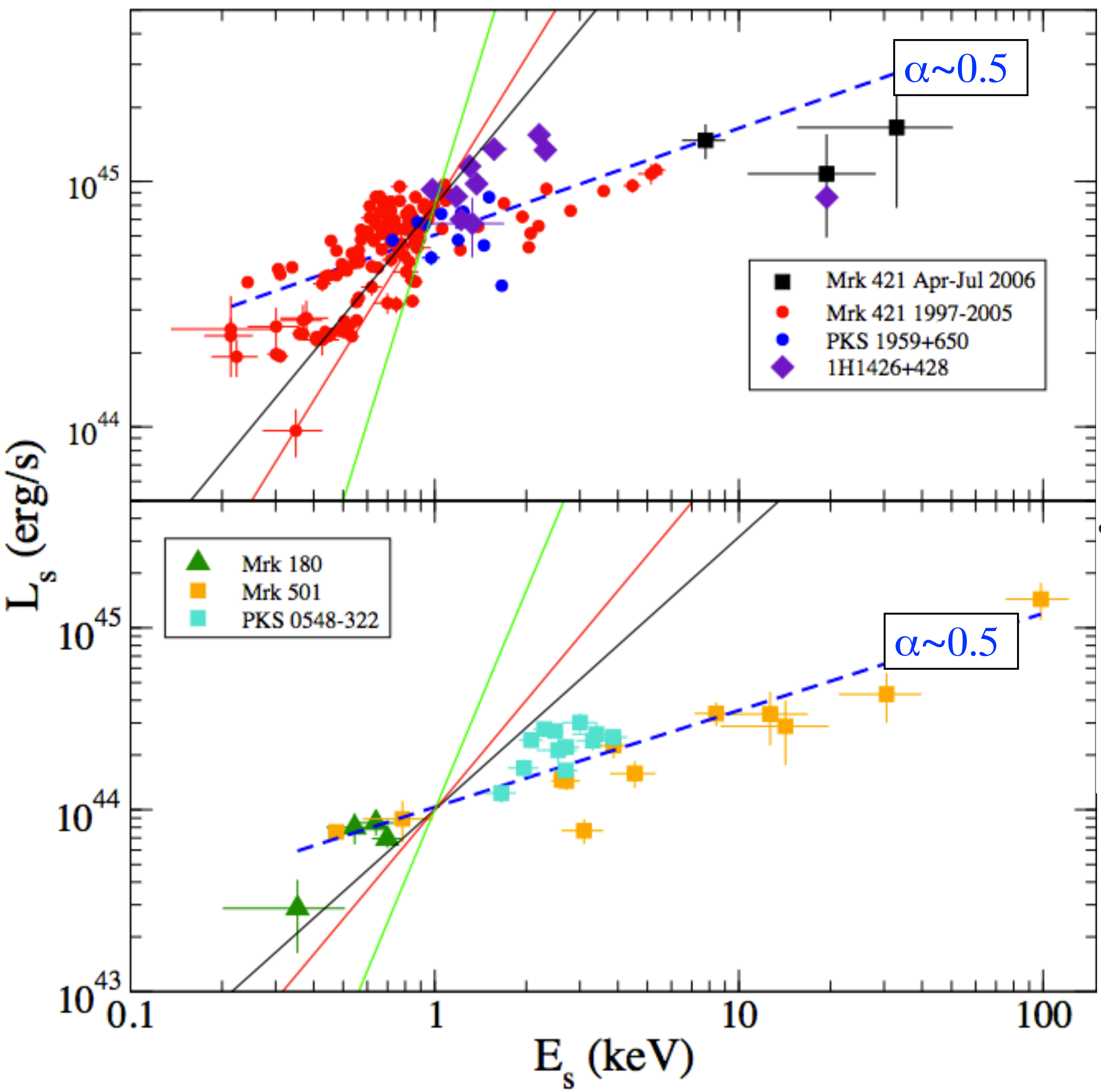




$$\boxed{R \leq c\,\Delta t\,\delta/(1+z)}$$

acceleration signature in the E_S-vs-L_s trend

long-trend main drivers

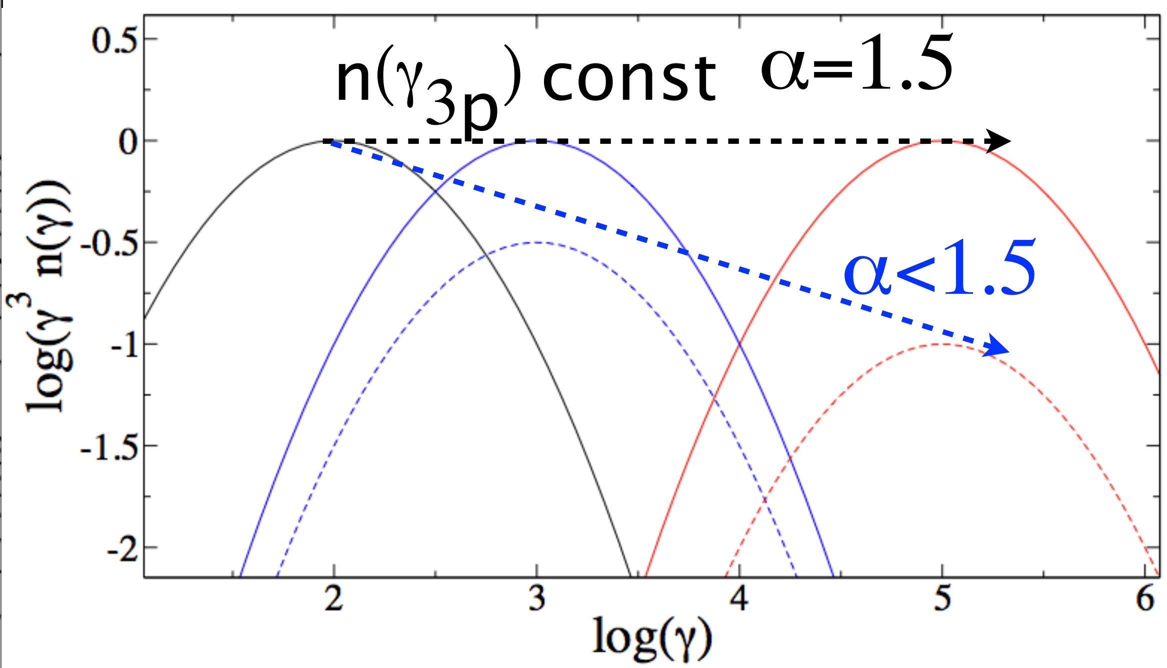


Tramacere A., et al.2009A&A...501

$$S_s(E_s) \propto n(\gamma_{3p})\gamma_{3p}^3 B^2\delta^4$$

$$E_s \propto \gamma_{3p}^2 B\delta.$$

$$S_s \propto (E_s)^\alpha$$

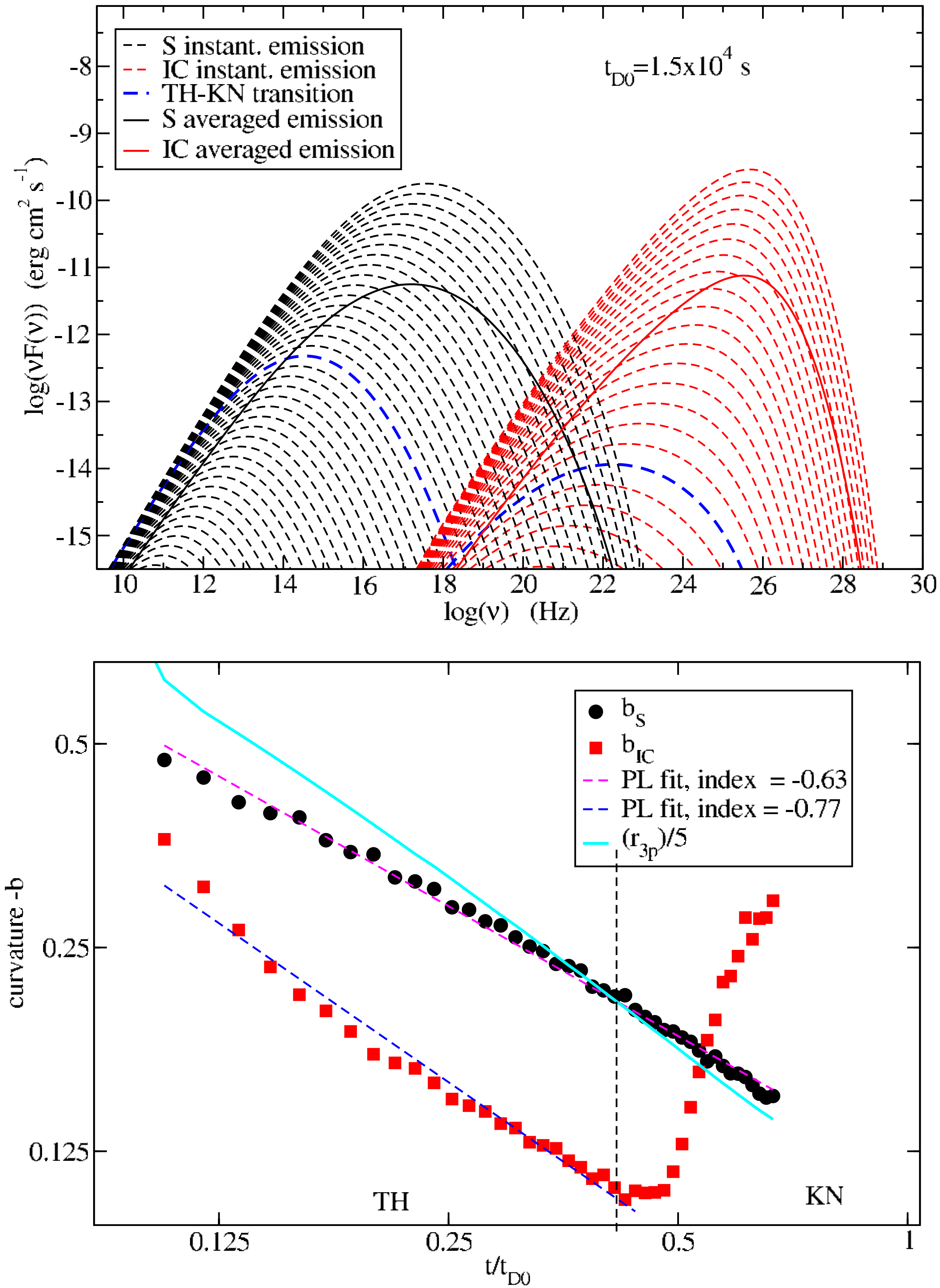


- γ_{3p}↑ and n(γ3p)↓ => α<1.5

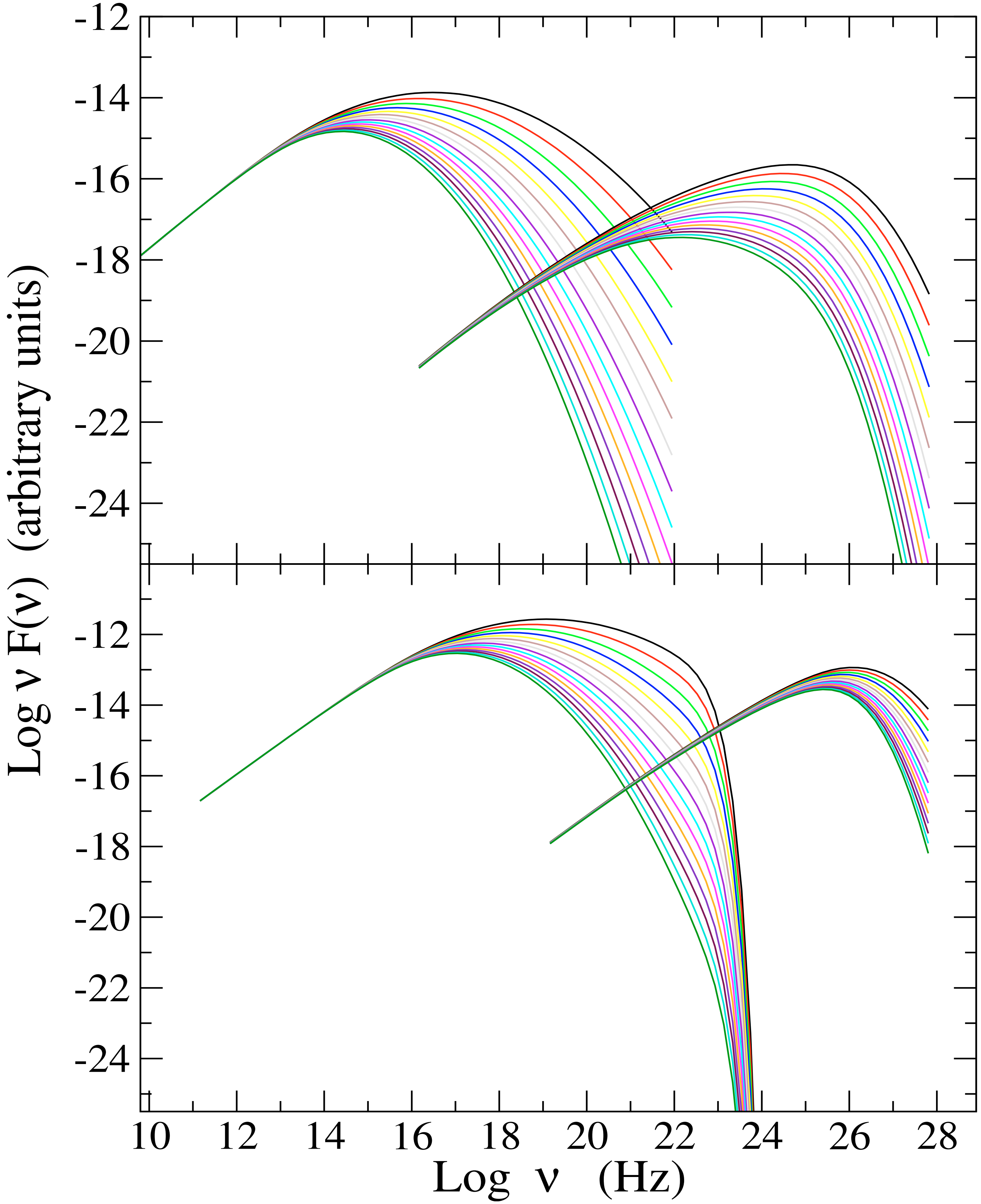
acceleration+energy conservation

- B−>α=2.0,
- δ−>α=4

incompatible as long-trend main driver

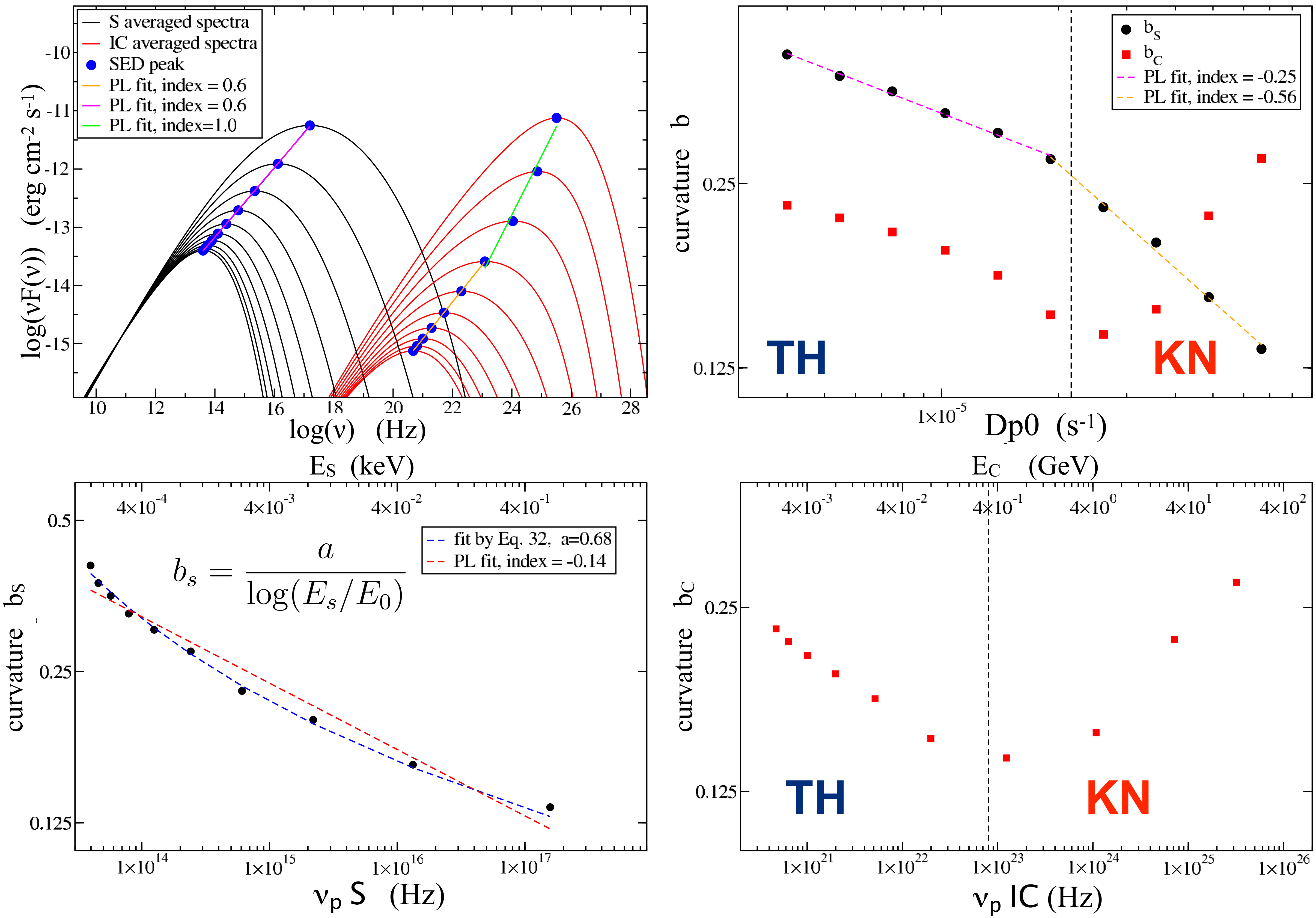


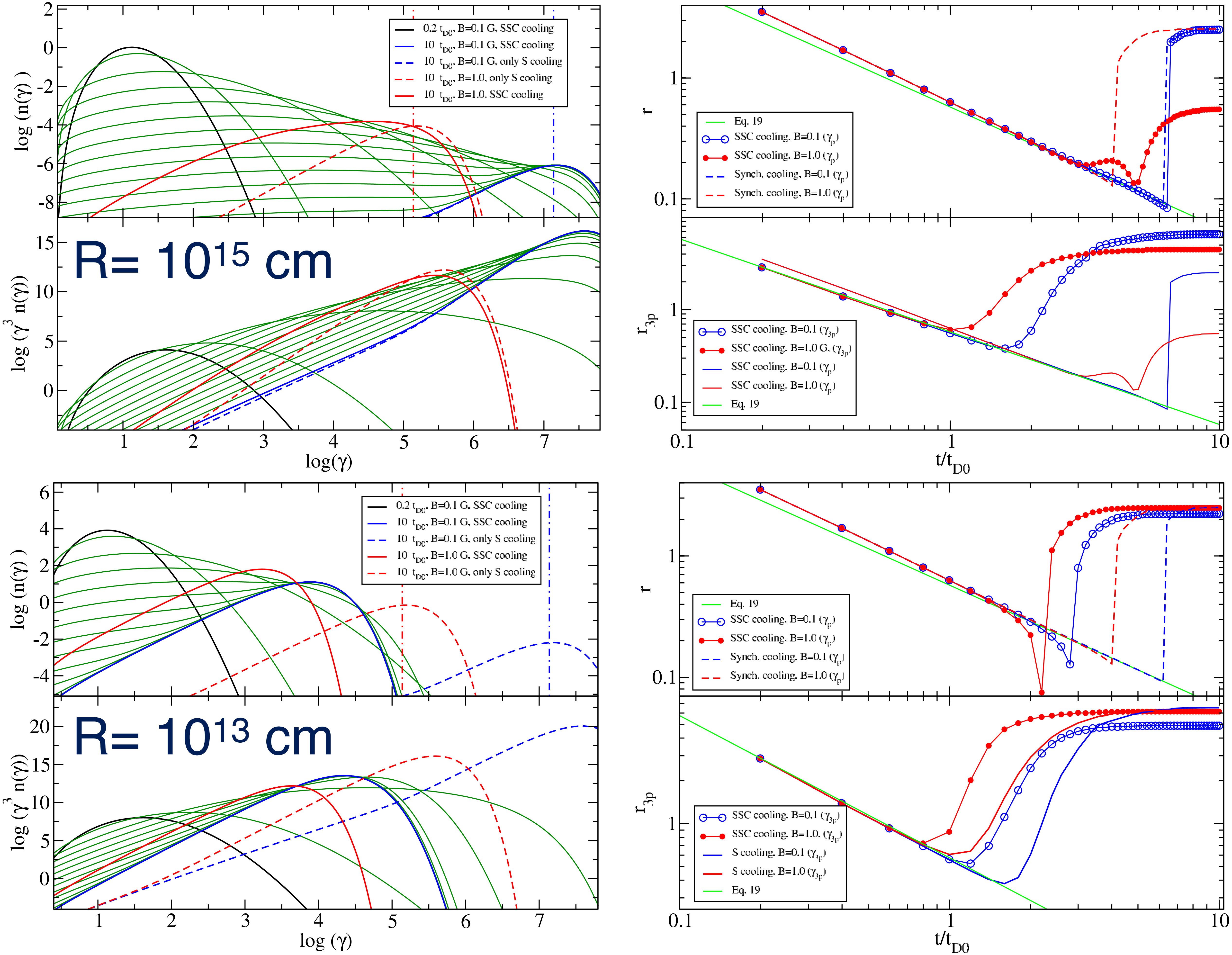
S instant. emission
IC instant. emission
TH-KN transition
S averaged emission
IC averaged emission
t_{D0}=1.5x10^4 s
log(νF(ν)) (erg cm^2 s^{-1})
log(ν) (Hz)
b_S
b_{IC}
PL fit, index = -0.63
PL fit, index = -0.77
$(r_{3p})/5$
curvature -b
TH
KN
t/t_{D0}



Log ν F(ν) (arbitrary units)
Log ν (Hz)
-12
-14
-16
-18
-20
-22
-24
10
12
14
16
18
20
22
24
26
28

D_p-driven trends $t_{D=}[1.5x10^4-1.5x10^5]$, L_{inj}=const.

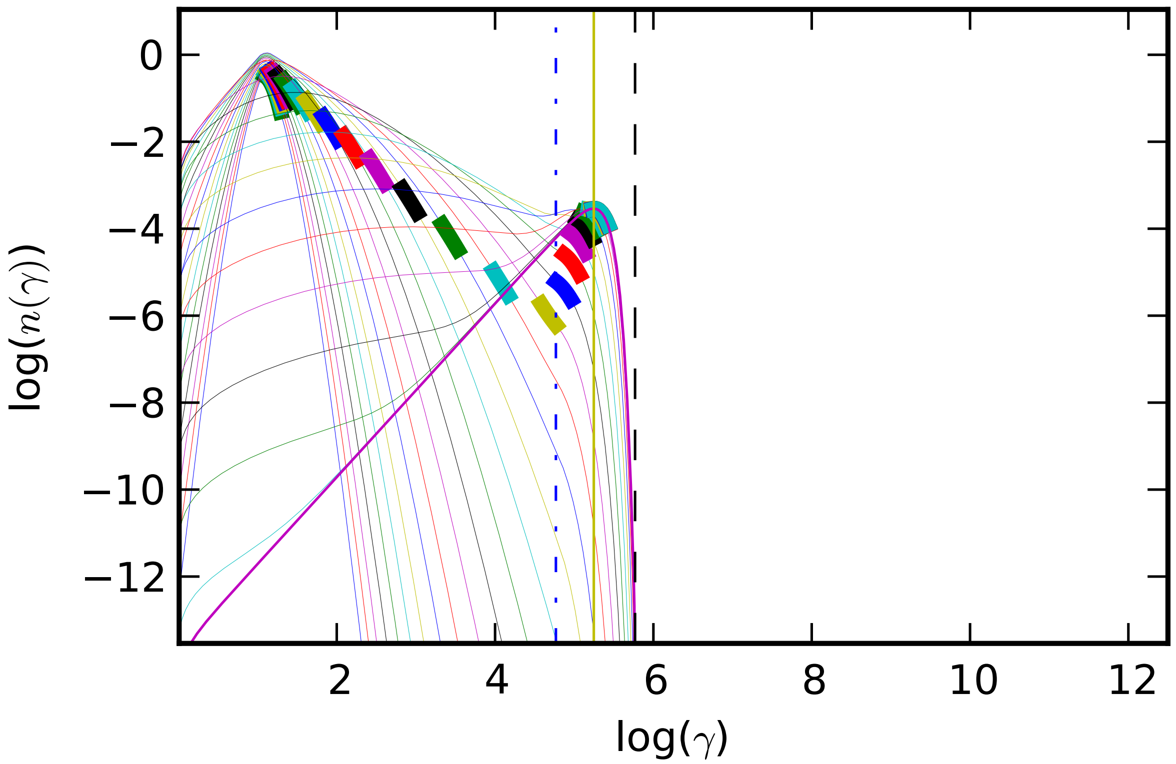




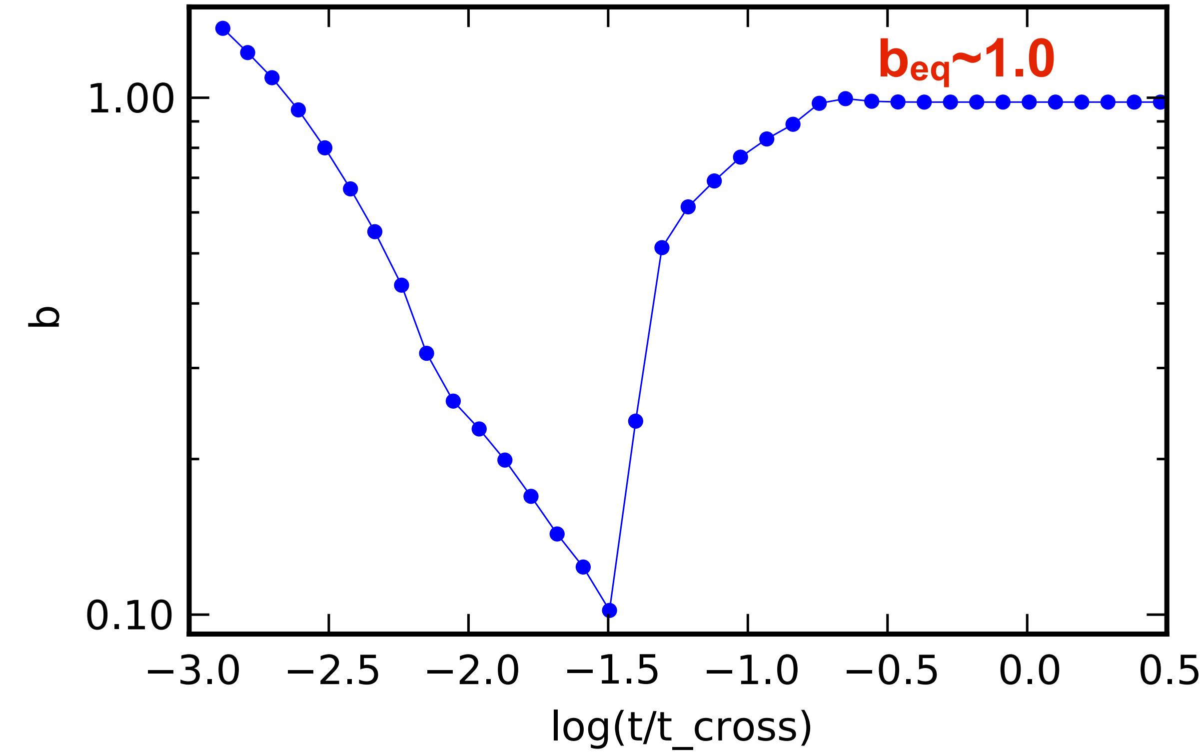
R= 10^15 cm
R= 10^13 cm
log (n(γ))
log (γ^3 n(γ))
log(γ)
log (γ)
r
r_3p
t/t_D0
0.2 t_D0, B=0.1 G, SSC cooling
10 t_D0, B=0.1 G, SSC cooling
10 t_D0, B=0.1 G, only S cooling
10 t_D0, B=1.0, only S cooling
10 t_D0, B=1.0, SSC cooling
10 t_D0, B=1.0 G, SSC cooling
10 t_D0, B=1.0 G, only S cooling
Eq. 19
SSC cooling, B=0.1 (γ_p)
SSC cooling, B=1.0 (γ_p)
Synch. cooling, B=0.1 (γ_p)
Synch. cooling, B=1.0 (γ_p)
SSC cooling, B=0.1 (γ_3p)
SSC cooling, B=1.0 G, (γ_3p)
SSC cooling, B=1.0, (γ_3p)
S cooling, B=0.1 (γ_3p)
S cooling, B=1.0 (γ_3p)

effect of λ_{max}, λ_{coher}

B=1.0 G, t_{D0}=1E3 s,q=2.0, **$\lambda_{max}=10^9$ cm**



synch. peak curvature



B=1.0 G, t_{D0}=1E3 s,q=2.0, **$\lambda_{max}=10^{15}$ cm**